

**discussion**



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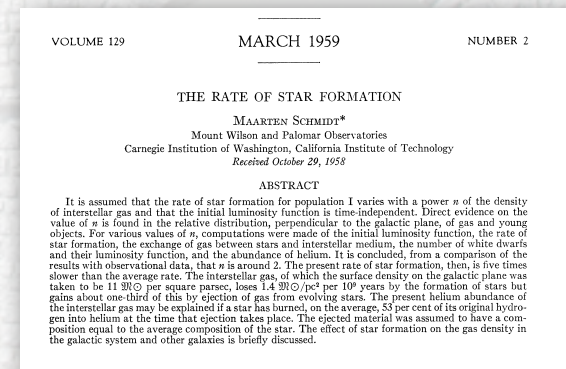


# Star formation relations

- Is there a one-to-one relation between gas density and star formation?
- If (really) yes, what does it mean? What does it tell us about the physical processes that govern stellar birth? What does a slope of 2, 1.5, 1.0, or 0.8 tell us?
- What may be the hidden parameters that the typical KS plot does not reveal?

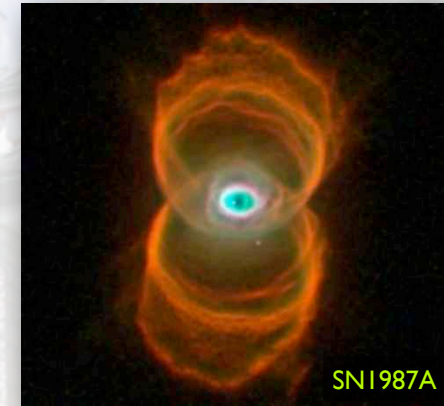
## Origin of turbulence

- What drives the supersonic ISM turbulence?
- Can we measure that? How can we distinguish observationally between different mechanisms?
- What are good statistical measures to characterize turbulent flows in the ISM?



# Importance of feedback

- Processes
  - protostellar outflows
  - H<sub>2</sub> regions
  - radiation pressure
  - supernovae
- How important is feedback for cloud evolution and dynamics?
  - On what scales / evolutionary phases are they important?
  - Why do SF clouds look identical to non-SF clouds?
- Is SF really self-regulated?  
(accelerated SF vs. slow equilibrium evolution)





# Importance of cooling and heating

- Do we appreciate the importance of the thermodynamics in the ISM sufficiently?
- Dust? Do we understand how dust works? Dust acts as thermostat for high-density gas.
- What are good measures for the age / evolutionary state of clouds?

## Energy balance in the ISM

- Competition between many different processes
  - turbulent kinetic energy density
  - magnetic energy density
  - thermal energy density
  - cosmic rays! (largely neglected!)

$$\frac{1}{2} \frac{d^2 I_E}{dt^2} = 2(E_{\text{kin}} + E_{\text{int}}) - 2(\tau_{\text{kin}} + \tau_{\text{int}}) + M + \tau_M - W - \frac{1}{2} \frac{d\Phi}{dt}$$

$$\frac{1}{2} \frac{d^2 I_L}{dt^2} = 2(E_{\text{kin}} + E_{\text{int}}) - 2\tau_{\text{int}} + M + \tau_M - W,$$

Ballesteros-Paredes (2006)

# Simulations and Observations

- What can we learn from numerical simulations?
- How can we compare models with observational data?
- Do we need every increasing resolution or do we need more physics?
- How can small-scale simulations best provide subgrid-scale models for galactic simulations?

