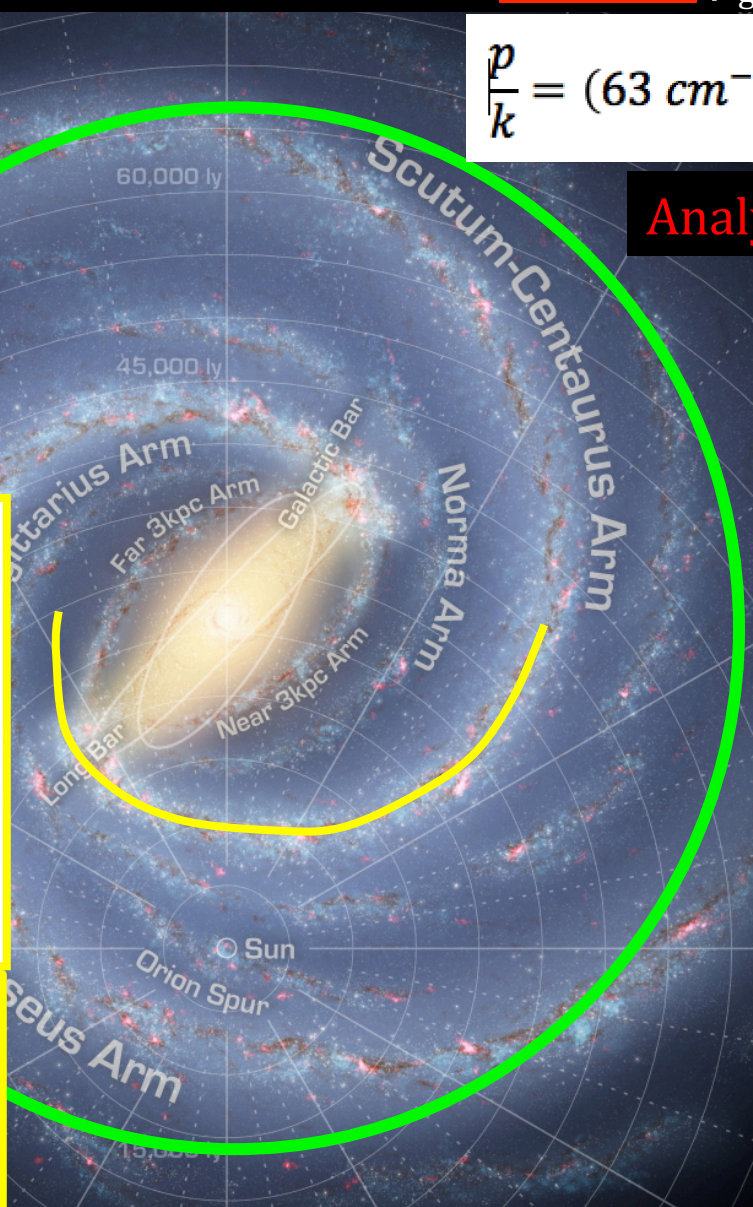
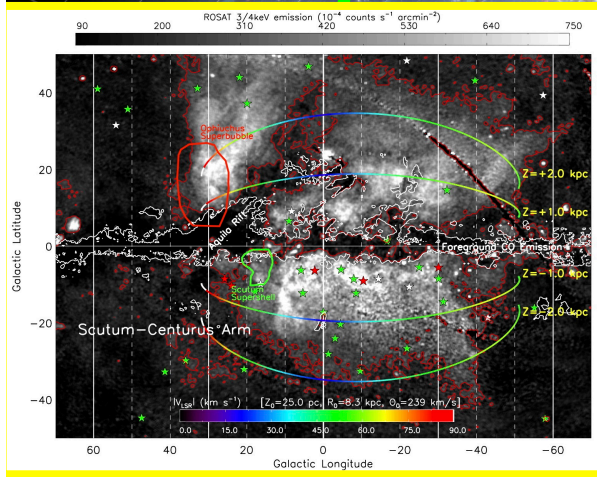
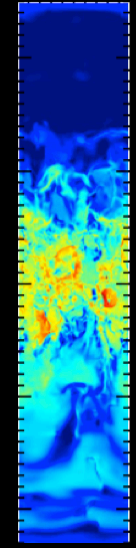


Hydrostatic pressure/SF in the MW: We measure $\rho_{\text{gas}}(R,\theta,z)$ and $\rho_*(R,\theta,z)$!

$$\frac{p}{k} = (63 \text{ cm}^{-3} \text{ K}) \Sigma_{\text{gas}} \Sigma_* \left(\frac{c}{1+c} \right) \quad c = \frac{Zg}{Z_*}$$

Analytical

3D AMR
MHD
models
with SNR
driven
turbulence



1. The Scutum-Centaurus arm, a.k.a. Molecular Ring, is under higher pressure than you would think because of the weight of the “invisible” ionized gas and possible outflow/wind.

2. The density of the stellar disk “truncates” at $R=13.5\pm 1$ kpc, exactly where the molecular gas dies away. Did you know this?

