Structure and Fragmentation of a high line-mass filament: Nessie

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ABSTRACT

Context. An increasing number of hundred-parsec scale, high line-mass filaments have been detected in the Galaxy. Their evolutionary path, including fragmentation towards star formation, is virtually unknown.

Aims. We characterize the fragmentation within the hundred-parsec-scale, high line-mass Nessie filament, covering size-scales between $\sim 0.1 - 100$ pc. We also connect the small-scale fragments to the star-forming potential of the cloud. Methods. We combine near-infrared data from the VISTA Variables in the Via Lactea (VVV) survey with mid-infrared Spitzer/GLIMPSE data to derive a high-resolution dust extinction map for Nessie . We then apply a wavelet decomposition technique on the map to analyze the fragmentation characteristics of the cloud. The characteristics are then compared with predictions from gravitational fragmentation models. We compare the detected objects to those identified in ~ 10 times coarser resolution from ATLASGAL 870 μ m dust emission data.

Results. We present a high-resolution extinction map of Nessie (2" full-width-half-max, FWHM, corresponding to 0.03 pc). We estimate the mean line mass of Nessie to be $\sim 627 M_{\odot} \text{ pc}^{-1}$ and the distance to be $\sim 3.5 \text{ kpc}$. We find that Nessie shows fragmentation at multiple size scales. The median nearest-neighbour separations of the fragments at all scales are within a factor of two of the Jeans' length at that scale. However, the relationship between the mean densities of the fragments and their separations is significantly shallower than expected for Jeans' fragmentation. The relationship is similar to the one predicted for a filament that exhibits a Larson-like scaling between size-scale and velocity dispersion; such a scaling may result from turbulent support. Based on the number of young stellar objects (YSOs) in the cloud, we estimate that the star formation rate of Nessie is $\sim 371 \text{ M}_{\odot} \text{ Myr}^{-1}$; similar values result if using the number of dense cores, or the amount of dense gas, as the proxy of star formation. The star formation efficiency is 0.017. These numbers indicate that by its star-forming content, Nessie is comparable to the Solar neighborhood giant molecular clouds like Orion A.

Key words. Stars: formation - Infrared: ISM - ISM: clouds - dust, extinction

Star formation is an important process in the evolution of galaxies and the Universe. It plays a crucial role in gas-tostars conversion through parameters such as star-forming rate and -efficiency, and the initial mass function (e.g., Mc-Kee & Ostriker 2007; Hennebelle & Falgarone 2012; Padoan et al. 2014). Star formation takes place in dense regions of molecular clouds, which appear to be commonly composed of filamentary structures (Schneider & Elmegreen 1979; Arzoumanian et al. 2011; Hacar et al. 2013; Schisano et al. 2014: Li et al. 2016: Kainulainen et al. 2017: Stutz & Gould 2016, see André et al. 2014 for a review). Filaments are observationally defined as any elongated structures with an aspect ratio larger than ~ 5 and a clearly higher density than their surroundings (Myers 2009). Given the link between filamentary structures and star formation, the processes driving the formation and evolution of filaments are linked with star formation rate and -efficiency. However, these processes are still not well understood.

Especially, the physics of filament fragmentation are not well known. This is mostly because determining the basic

characteristics of filaments is observationally challenging, as the cold molecular hydrogen is invisible to observations. Therefore, different tracers and techniques are needed to determine its distribution and properties (e.g., Lombardi & Alves 2001; Goldsmith et al. 2008; Goodman et al. 2009; André et al. 2014). Each of the techniques is sensitive to different density regimes and has different spatial resolution. For studies of the structures related to star formation, the resolution should clearly resolve the Jeans' length. This is about 0.1 pc for typical conditions of a molecular cloud (gas temperature T = 15 K, average density $\overline{n}(H) = 10^5$ cm⁻³). This currently limits the observations mostly on nearby (< 500 pc) clouds. Interferometric observations can increase this resolution farther, but they have their own caveats (e.g. spatial filtering, slow mapping speed).

However, the nearby clouds that can be systematically mapped in high-enough resolution are mainly low-mass clouds, containing mostly low line-mass filaments (mass per unit length of $(M/l) \lesssim$ a few ×10 M_{\odot}) forming almost exclusively low-mass stars. An exception to this is the integral shaped filament in the Orion A cloud (at distance 414 pc, Menten et al. 2007) whose fragmentation have been

analyzed in high-resolution using interferometric data (e.g., Takahashi et al. 2013; Teixeira et al. 2016; Kainulainen et al. 2017). But in general, our current observational picture of filaments is mostly built by data on low-mass clouds. Filaments that have much higher line-masses $((M/l) \gg 100 M_{\odot})$, which may also be able to form high-mass stars, have been identified in numbers, but they are typically located at further distances (e.g., Jackson et al. 2010; Hernandez et al. 2012; Busquet et al. 2013; Kainulainen et al. 2013; Ragan et al. 2014; Wang et al. 2014; Beuther et al. 2015; Abreu-Vicente et al. 2016; Henshaw et al. 2016; Li et al. 2016; Wang et al. 2016). Modern facilities are only approaching the ability to study them systematically in resolution that resolves the Jeans' scale.

Recently, Kainulainen & Tan (2013) developed a dust extinction based method that allows studying infrared dark molecular clouds at a resolution of $\sim 2''$ over a wide dynamic range of column densities, using a combination of near- and mid-infrared observations (see also Lombardi & Alves 2001; Kainulainen et al. 2011; Butler & Tan 2012). This method allows us to analyze the internal structure of clouds up to several kpc distance at ~ 0.1 pc resolution, enabling fragmentation studies of high line-mass filaments.

With the high-resolution mapping technique in hand, we can address a basic question related to filament fragmentation: What are the fragmentation characteristics of massive filaments and are they in agreement with gravitational fragmentation models?

In this paper, we take the advantage of the highresolution provided by the Kainulainen & Tan (2013) extinction mapping technique and analyze the fragmentation characteristics of a ~ 100 pc long, high line-mass filamentary cloud known as "Nessie" (Jackson et al. 2010). It is supposedly located within the Scutum-Centaurus Arm of the Milky Way (Goodman et al. 2014; Ragan et al. 2014; Zucker et al. 2015; Abreu-Vicente et al. 2016). The high resolution allows us to characterize the cloud structure and to gauge the fragmentation processes over a wide range of scales ($\sim 0.1 \text{ pc} - 100 \text{ pc}$). We will use the dust extinction mapping technique in conjunction with the nearinfrared (NIR) data from the ESO/VISTA telescope and mid-infrared (MIR) data from the *Spitzer* satellite. We then analyze the derived column density map with a hierarchical structure-identification technique and examine the fragmentation of the cloud over multiple size-scales. The results will then be compared with theoretical models and other clouds in literature. Finally, we compare our identified small scale structures to clumps identified in low resolution (~ 20") dust emission maps by Csengeri et al. (2014). This demonstrates how structures identified from data with ten times coarser resolution fragment when viewed in finer detail.

2. Data

2.1. Infrared data and data reduction

We employ NIR imaging data from the VVV (VISTA Variables in the Via Lactea) survey (Saito et al. 2012) at the 4.1 m VISTA telescope of the Paranal Observatory. The calibrated and reduced data are publicly available in the ESO archive. Specifically, we used the J, H, K_S spectral bands of the tiles d069 and d068. For each filter band there are two $t_{\rm exp} = 80 \ s$ exposures and additionally there are

8 and 12 $t_{exp} = 16 \ s$ exposures of tiles d069 and d068 in the K_S band, respectively. The pixel size of the images is $0.34" \times 0.34"$. Detailed information about the observations can be found in Table A.1 in the Appendix. We stacked the observations and performed PSF photometry with the daophot package (Stetson 1987) using the IRAF software. The PSF model was created from bright isolated stars with the model radius of $r_{\text{PSF}} = 1.5$ ". The different spatial resolutions of the single observation epochs has no significant effect on the photometry as we show in the Appendix B. The daophot algorithm identifies and extracts extended sources and cosmic rays, and we expect only a very low contamination of the data by galaxies, because we are looking through the galactic mid-plane. The zero-point magnitudes were defined by comparing the resulting magnitudes of the stars with the corresponding stars of 2MASS, that are flagged as good photometric quality (Skrutskie et al. 2006; Cutri et al. 2003). This resulted in zero-points $J_{zpt} = 21, 21 \text{ mag}$, $H_{zpt} = 21,22 \text{ mag}, K_{S,zpt} = 20,88 \text{ mag}.$ The resulting data shows the expected shape in the near-infrared colorcolor scatter plot (Fig. 1), with a bump for the main sequence stars and an elongated distribution for stars with varying reddening. We also tested the photometry measurements for completeness by adding artificial stars. We could identify all artificial stars up to a magnitude of about $J_{com} = 16.5 \text{ mag}, H_{com} = 15.5 \text{ mag}, \text{nd } K_{S,com} = 15.0 \text{ mag}.$



Fig. 1. Near-infrared color-color diagram of all sources in the mapped area extracted from the VVV survey with the photometric errors lower than 0.02 mag. The blue crosses indicate non-redded intrinsic colors of stars (Bessell & Brett 1988). The arrow shows the reddening for an extinction of $A_V = 10$ mag

We also employ MIR 8 μ m imaging data from the *Spitzer*/GLIMPSE survey, data release 5 (Benjamin et al. 2003; Churchwell et al. 2009). The pipeline-reduced (S13.2.0 1v04) images were retrieved from the IRSA¹ database and used as such. The 8 μ m image has a spatial resolution of 2.4" and a pixel size of 1.2" times 1.2". The used tile is centered around RA = 16: 43: 14.08, DEC = -16: 00: 15.92. The effective integration time of the tile is 1.2 s.

¹ http://irsa.ipac.caltech.edu/data/SPITZER/GLIMPSE/

2.2. ATLASGAL data

We also use data from the APEX telescope large area survey of the galaxy (ATLASGAL, Schuller et al. 2009) for a comparison with our extinction data. The survey was obtained by the Millimeter and Submillimeter Group of the Max-Planck-Institut für Radioastronomie from 2007 to 2010 at the Atacama Pathfinder Experiment (APEX) located on Chajnantor in Chile. The survey instrument was the Large APEX Bolometer Camera (LABOCA) observing at 870 μm , which traces the thermal dust emission. The resolution of the survey is $\Omega = 19.2$ " with a sensitivity in the range of 40 - 70 mJy/beam. The maps covering the Nessie filament are centered at $l = -22.5^{\circ}$, $b = 0.0^{\circ}$ and $l = -19.5^{\circ}$, $b = 0.0^{\circ}$ and were observed on August 18th and 21st of 2007. The flux per beam, F_{ν} , of the ATLASGAL map can be used to estimate the hydrogen column density $N(H_2)$ under the assumptions of a constant gas-to-dust ratio of R = 100 and a dust opacity of $\kappa_{345 \text{ GHz}} = 1.85 \text{ cm}^2 \text{g}^{-1}$, which was extrapolated by Schuller et al. (2009) based on the work of Ossenkopf & Henning (1994),

$$N(H_2) = \frac{F_{\nu}R}{B_{\nu}(T_d)\Omega\kappa_{\nu}\mu_{H_2}m_H} .$$
⁽¹⁾

 $B_{\nu}(T_d)$ is the Planck function at the dust temperature T_d , m_H is the mass of a hydrogen atom, and μ_{H_2} the mean molecular weight of the interstellar medium with respect to hydrogen molecules, which is 2.8 (Kauffmann et al. 2008).

Csengeri et al. (2014) have identified clump-like structures from the ATLASGAL data using 2D Gaussian fitting (Gauss Clump Source Catalog, GCSC). It provides the position, peak flux F'_{ν} and integrated flux S_{ν} , the half maximum major and minor axes and the position angle of the clumps. We then calculated the masses of the clumps from (Schuller et al. 2009):

$$M = \frac{S_{\nu}R\,d^2}{B_{\nu}(T_d)\kappa_{\nu}} \,, \tag{2}$$

where R is the gas-to-dust ratio and d the distance towards the clump.

3. Extinction mapping technique

We employ the technique from Kainulainen & Tan (2013), which is based on combining extinction maps made at two wavelength regimes: in near-infrared using NICER (Near-Infrared Color Excess Revisited, Lombardi & Alves 2001) and in mid-infrared using the absorption against the Galactic background (e.g., Peretto & Fuller 2009; Butler & Tan 2012). Below, the implementation of the two techniques is explained in detail.

3.1. NICER-Method

We use the NICER method in conjunction with JHK_S photometric data of the VVV survey. The method is based on near-infrared color measurements of stars shining through the molecular cloud and comparison of those with stars of a reference field that is (optimally) free from extinction. The observed reddening towards the cloud region is used to estimate the extinction by adopting a wavelength dependent reddening law. The extinction values towards each star are then used to derive a spatially smoothed dust extinction map.

This method is straightforward to apply for nearby clouds (d < 500 pc, e.g., Lombardi et al. 2006; Froebrich et al. 2007; Juvela et al. 2008; Goodman et al. 2009; Kainulainen et al. 2009), where the contamination due to stars between the cloud and the observer is small. The extinction towards more distant clouds might be underestimated because of these (mostly unreddend) foreground stars, especially in high extinction regions where the fraction of foreground sources is high (Lombardi 2005). The foreground stars do not trace the dust reddening caused by the cloud, but only the reddening along the line of sight until the cloud. Therefore, foreground sources should be removed as accurately as possible, which is challenging in practice, because of the degeneracy between the intrinsic colors of stars and reddening caused by extinction.

The subtraction of the foreground is also necessary for the reference field (see, Kainulainen et al. 2011). Due to diffuse dust in the Galactic plane stars in the reference field, located at the same distance as stars behind the cloud, are redder than the ones at closer distance. Therefore, foreground stars shift the mean color of the reference field towards blue, which leads to an overestimation of the extinction. For the implementation of the NICER method we have to find a reliable way to remove the effect of the foreground stars. This is described in the following.

First, we derive a "dirty" extinction map using arbitrary reference colors and use this map to identify low- and high-extinction regions. The low-extinction region (Fig. 2; $338.39^{\circ} < l < 338.58^{\circ}; -0.36^{\circ} < b < -0.21^{\circ})$ is then used as control field to estimate the reference colors, indicating the average star colors without dust reddening by the cloud. In the regions of high extinction, identifying foreground stars is simple: they appear as a distinct feature in the frequency distribution of individual extinction measurements (cf., Kainulainen et al. 2011). For regions of lower extinction the feature is less distinct, but under the assumption of uniformly distributed foreground stars the position and width of the frequency distribution remains the same; this fact can be used to statistically subtract the contribution of foreground stars to the reference field colors. To do this, we fit a Gaussian function G_{fg} , to the peak of the foreground stars in the extinction histogram $H(A_V)$ (Fig.3) and subtract these stars in a statistical sense from the distribution. To achieve this, we add a weighting term $(W_{\rm fg}(A_V^{(n)}))$, see Fig. 3) into the original NICER method. This weighting term suppresses the contribution of stars that might be foreground stars, and it is calculated in the following way

$$W_{\rm fg}(\widehat{A}_V^{(n)}) = \frac{H(A_V) - G_{\rm fg}}{H(A_V)} .$$
 (3)

The weighting term is introduced into equation (15) of Lombardi & Alves (2001) as shown here:

$$W^{(n)} = \frac{W(\theta - \theta^{(n)}) \cdot W_{\text{fg}}(\widehat{A}_V^{(n)})}{\operatorname{Var}(\widehat{A}_V^{(n)})} , \qquad (4)$$

where $W^{(n)}$ is the weighting of the *n*th star, $W(\theta - \theta^{(n)})$ is the weight for the distance between the actual location θ and the location of the *n*th star $\theta^{(n)}$, $W_{\text{fg}}(\widehat{A}_V^{(n)})$ is the foreground weight based on the estimated extinction of the *n*th



Fig. 2. Extinction maps of Nessie derived using the NIR data of the VVV survey (top), mid-infrared data of the *Spitzer* Space Telescope (center) and their combination (bottom). The black areas indicate regions of bright mid-infrared emission that hampers extinction mapping. The red rectangle marks the area used for estimating the reference colors for the NICER method. The white circle marks the high extinction region used to estimate the mid-infrared foreground emission.

star, and $\operatorname{Var}(\widehat{A}_V^{(n)})$ is variance of the estimated extinction of the *n*th star.

With this method the contribution of foreground stars was subtracted statistically from the mean color of the reference field to calculate an estimate of the mean color of the stars in the background of the cloud. The statistical subtraction is done in the JHK-color-color space, where the density of foreground stars was subtracted from the density of the reference field stars in each color-color bin. Then the foreground-corrected number of stars per bin was calculated from the resulting density in the reference field. The foreground-corrected mean color was calculated from this sample of stars, which is also the estimate of the background color. The JHK-color-color histograms of the reference field before and after correction are shown in Appendix C.

With the foreground-corrected reference color and the method for extracting foreground sources the "true" near-infrared extinction map was calculated. The spatial reso-



Fig. 3. Left: The black line shows histogram of the calculated extinction from a high extinction region. The red line marks the Gaussian fitted to the peak of foreground stars. **Right:** The black line shows the empirical weighting function, which is derived like shown in Eq. 3. The red line shows the fitted function, which is then introduced into the weighting function of the NICER method (Eq. 4).

lution of the map is given by the width of the Gaussian smoothing function that is used to smooth the pencil-beam measurements towards the stars onto the map grid. The pixel size is chosen following the surface number density of background sources so that even in high extinction regions, where the density is lower, each pixel covers at least two stars. For the VVV data we concluded that a pixel size 24" is sufficient, which leads to a beam width of 48".

3.2. Mid-infrared Extinction Measurement

We use the MIR imaging data from the GLIMPSE survey to estimate extinction through the cloud at 8 μ m. Generally, the technique is based on the extinction of the diffuse MIR emission from the Galactic plane by the dust of the cloud (see, e.g., Johnstone et al. 2003; Peretto & Fuller 2009; Butler & Tan 2012). Consider a simplistic geometry in which the intensity of radiation behind the cloud is I_0 . Then, the intensity right in front of the cloud is $I_1 = I_0 e^{-\tau_8}$, in which τ_8 refers to the optical depth at the Spitzer 8 μ m band. An observer detects the intensity $I_{\rm obs,1}$, which in addition to I_1 contains the intensity $I_{\rm fg}$ that is emitted from between the cloud and the observer, i.e., $I_{\rm obs,1} = I_1 + I_{\rm fg}$. A line-of-sight off the cloud does not exhibit extinction and the observed intensity is $I_{\rm obs,0} = I_0 + I_{\rm fg}$. Combining these relations, one can solve the optical depth

$$\tau_8 = \ln \frac{I_{\rm obs,0} - I_{\rm fg}}{I_{\rm obs,1} - I_{\rm fg}}.$$
(5)

Thus, the optical depth along the line of sight can be estimated through measurements of the off-cloud and foreground intensities.

Various approaches have been used in the past to estimate the off-cloud and foreground intensities (see, e.g., Johnstone et al. 2003; Peretto & Fuller 2009; Ragan et al. 2009; Butler & Tan 2012). We follow an approach similar to Butler & Tan (2012) to which we refer to for a thorough description and discussion; we describe here only the implementation of the technique in our case. The off-cloud intensity is estimated using a median-filtered 8 μ m map. Prior to the filtering, the most prominent dark features are masked from the map by using a threshold intensity of 46 MJy sr⁻¹. The filter size defines the upper limit of the structures the map is sensitive to. However in our case, we will later combine the MIR-derived map with the NIR-derived map that probes spatial scales larger than 24". Therefore, the filter function width is not a crucial choice for us, as long as there is some overlap of scales probed by the MIR and NIR maps. Following the discussion in (Ragan et al. 2009), we chose the filter width of 3'.

The foreground intensity is estimated with the help of the pixels with lowest intensities (i.e., highest extinctions) in the 8 μ m data. If several independent high-extinction regions show similar intensities, one can assume that such locations are opaque and the intensity towards them is a reasonable estimate of the foreground intensity. The smallest intensities detected in the cloud area are $I_{obs,1} = 24.6$ $MJy sr^{-1}$. There are three independent locations in the cloud where the intensity is within $2\sigma_{\rm rms}$ of this value (the rms noise, $\sigma_{\rm rms}$, of the GLIMPSE data is ~0.6 MJy sr⁻¹. Reach et al. 2005). One of them $(l, b = 337.895^{\circ}, -0.563^{\circ})$ is extended, containing tens of pixels, which indicates that the region indeed is saturated. The number of the saturated regions is relatively low given the large extent of the cloud on the sky; it would be preferable to have numerous saturated regions along the cloud. Regardless, we adopt the value of 24 MJy sr^{-1} for the foreground intensity. We note that the resulting fraction of foreground emission, i.e., $I_{\rm fg}/I_{\rm obs,0} \approx 45\%$, well in the range of the foreground intensities typically determined for IRDCs (e.g., Butler & Tan 2012).

Following the estimation of the off-cloud and foreground intensities, Eq. 5 is used to compute an optical depth map for Nessie. Finally, the map is converted into units of visual extinction by adopting the ratio between 8 μ m and V band optical depths (based on Cardelli et al. 1989; Ossenkopf & Henning 1994, see Kainulainen & Tan 2013)

$$A_{\rm V} = 33.6\tau_8.$$
 (6)

The resulting extinction map is shown in Fig. 2.

3.3. Combined Near- and Mid-infrared Extinction Measurement

We have now derived the near-infrared and mid-infrared extinction maps; both show some advantages and disadvantages. The near-infrared data is sensitive to low column densities, but have a low resolution. The mid-infrared data have a good resolution, but are much less sensitive. Therefore, we now want to combine them and use the near-infrared data to recalibrate the mid-infrared data, thus prevailing high spatial resolution of the mid-infrared data while imposing the good calibration of the near-infrared data on them. The combination of near- and mid-infrared extinction maps follows the scheme described in Kainulainen & Tan (2013). The combined maps deliver a higher dynamic range of extinction compared to maps computed from nearor mid-infrared data alone (Fig. 2). The correlation between the two maps is shown in the Appendix D.

The combined map is then converted to molecular hydrogen column density by applying the conversion of Savage et al. (1977); Bohlin et al. (1978); Rachford et al. (2002):

$$N(\mathrm{H}_2) = A_V \cdot 0.94 \cdot 10^{21} \mathrm{cm}^{-2} \mathrm{mag}^{-1} , \qquad (7)$$

using a typical reddening constant of $R_V = 3.1$ (Schultz & Wiemer 1975) and assuming all hydrogen atoms are in molecular form.

4. Results

4.1. Distance determination

The foreground star density measurements (see Section 3.1) allow us to estimate the distance of Nessie independently of previous, kinematic distance estimates. We can compare the measured surface density of foreground stars with a distance-dependent stellar surface density model of the Galaxy. We used the Besancon Galactic stellar distribution model (Robin et al. 2003) to estimate the distance, see Fig. 4. For a more detailed description of the method see Kainulainen et al. (2011); Ioannidis & Froebrich (2012). The most important input parameter of the stellar distribution model is the extinction caused by diffuse interstellar dust. We used the measurements by Marshall et al. (2006) to estimate the mean extinction along the line of sight towards Nessie. For an estimate of the uncertainty we also estimated the minimum and maximum extinction, which indicate the upper and lower limit of the surface density (Fig. 4). We neglected other, potentially significant uncertainties in our distance calculations such as the uncertainty of the measured number surface density of the foreground stars or of the stellar distribution model. Therefore, the uncertainty of the distance is underestimated and it is more likely in the order of 15 % corresponding to $\Delta d \approx 0.5$ kpc (Kainulainen et al. 2011).

The result of our distance estimate is $d_{\rm extinction} = 3.5 \pm 0.5$ kpc, which is in agreement with the kinematic distance estimations of Jackson et al. (2010), $d_{\rm HCN} = 3.1$ kpc. We find also dynamical distance measurements from Wienen et al. (2015) for 14 ATLASGAL sources likely embedded in the Nessie cloud. Their distances range between 3.0 kpc and 3.5 kpc, which is also in agreement with our estimate. The distance of ~ 3.5 kpc suggests that Nessie is associated with the Scutum-Centaurus spiral-arm of the Milky Way as suggested by Goodman et al. (2014) and Ragan et al. (2014).



Fig. 4. Predicted stellar surface density based on the Besançon stellar distribution model (Robin et al. 2003). The blue area indicates the uncertainty arising from the scatter in the diffuse extinction measurements. The horizontal line represents the measured foreground star surface density and the vertical lines the resulting estimates of the distance and its uncertainty.

4.2. The large-scale structure

The combined near- and mid-infrared extinction map of the Nessie cloud is shown in Fig. 5 and zoom-ins in Figs. 6, 7, and 8. For comparison, Fig. 2 shows the near-infrared based map, mid-infrared based map, and their combination.

The filament has a length of ~ 1.1° following the central, dense main axis (neglecting inclination) and perpendicular a width of ~ 0.05°. This corresponds to a physical size of 67 pc × 3 pc at the distance of d = 3.5 kpc. The width of the extinction structures, defined at the column density contours of about $A_V = 3$ mag, varies along the filament. This can be seen in the zoomed in map of Nessie (Fig. 6). In the region between $338.57^{\circ} < l < 338.95^{\circ}$ the low column density material is located only towards the south of dense main axis, between $338.23^{\circ} < l < 338.30^{\circ}$ towards north and south and the rest of the filament shows almost no surrounding low column density material. These two low column density regions show also some less dense structures, which are mainly orientated almost perpendicular to the main filament.

We need to identify which structures that we see in the map are actually part of Nessie. This is difficult because we miss information about the line-of-sight velocities of the structures. However, the Nessie filament was confirmed as a velocity coherent structure by Jackson et al. (2010). Additionally, some areas lack the mid-infrared extinction data and cannot be used in the further analysis, such as the HIIbulb at $(l; b) = (337.95^\circ; -0.46^\circ)$ (Fig. 5), which is part of Nessie in Jackson et al. (2010). Therefore, the map needs to be cropped to the Nessie filament. To do this, we introduce a polygon around the cloud (see Fig. 5). The area selection is mainly based on by eye inspection of the derived column density map with orientation on the $A_V = 3$ mag contour and the observations published by Jackson et al. (2010).

We derive an estimate of the total cloud mass from the column density map, given by:

$$M_{\text{Nessie}} = \sum_{i,j} (N(\text{H}_2)_{i,j}) \cdot p^2 \cdot m_{\text{H}} \cdot \mu_{\text{H}_2} , \qquad (8)$$

where $N(\text{H}_2)_{i,j}$ is the column density of the (i, j) pixel of the map, $p = \tan(1.2^{\circ}) \cdot d_{\text{Nessie}}$ is the physical size of a pixel, m_{H} is the mass of the hydrogen atom, and $\mu_{\text{H}_2} = 2.8$ the mean molecular weight of the interstellar medium (Kauffmann et al. 2008). The total mass of the Nessie cloud within the polygon (Fig. 5) is $M_{\text{Nessie}} = 4.2 \cdot 10^4 \text{ M}_{\odot}$.

From the length and mass we calculate the mean line-mass of the filament (mass per unit length along the main axis of the filament). The mean line-mass of Nessie is $(M/l) = 627 \ {\rm M}_{\odot} \ {\rm pc}^{-1}$. As we neglected an inclination of the filament, which would increase its length, the derived line-mass is an upper limit. We note that there are variations in the line mass along the filament, both at large scales due to the varying amount of diffuse extinction and at small scales due to the substructure of the cloud.

4.3. Fragmentation Analysis

We analyzed fragmentation of Nessie simultaneously over a wide range of spatial scales using an algorithm explained in Kainulainen et al. (2014), which employs wavelet filtering to identify structures at various spatial scales. In short, the algorithm uses a spatial filtering algorithm based on the à Trous wavelet transform (Starck & Murtagh 2002) to decompose the column density map into scale-maps that describe structure at different scales. The different scales are defined as 2^i pixels, with $2 \leq i \leq 8$, where the limits are given by the pixel size for small scales and



Fig. 5. Column density map of the Nessie filament. The white polygon marks the area chosen for the mass estimate of the cloud. The green rectangles show the positions of the zoom-ins shown in Figs. 6, 7, and 8.



Fig. 6. Zoom-in number one of the column density map (Fig. 5). The black contours indicate the levels of 5, 10, 20, 30, 40, 50, 60 \cdot 10²¹ cm⁻². The white contour indicates the smoothed $A_V = 3$ mag level. Additionally, the Class1 ('x') and Class2 ('+') YSOs are marked in white. Article number, page 8 of 30



Fig. 7. Zoom-in number two of the column density map (Fig. 5). The black contours indicate the levels of 5, 10, 20, 30, 40, 50, $60 \cdot 10^{21} \text{ cm}^{-2}$. The white contour indicates the smoothed $A_V = 3$ mag level. Additionally, the Class1 ('x') and Class2 ('+') YSOs are marked in white. Article number, page 9 of 30



Fig. 8. Zoom-in number three of the column density map (Fig. 5). The black contours indicate the levels of 5, 10, 20, 30, 40, 50, $60 \cdot 10^{21} \text{ cm}^{-2}$. The white contour indicates the smoothed $A_V = 3$ mag level. Additionally, the Class1 ('x') and Class2 ('+') YSOs are marked in white. Article number, page 10 of 30

the cloud size for large scales. Individual structures are then identified from each scale-map using the clumpfind-2D algorithm (Williams et al. 1994). This provides the position, the size in x and y direction, and the total amount of column density of the structures $N(\rm H)_{tot}$.

For reliable detection of structures, it is necessary to estimate the noise level of each scale-map. The noise level is estimated as the standard deviation σ of an (almost) extinction free area. The size of the area corresponds to the size-scale of the largest scale map. To test the robustness of the structure identification, we tested the clumpfind-2D algorithm for contour level separations of 1.5σ , 3σ , 4σ and 5σ with the lowest level at 3σ . The results do not show a significant difference and we chose the level separation of 3σ .

The numbers of structures identified at each scale using the chosen technique are listed in Table 1. The number of structures increases towards smaller scales, but drops significantly for the smallest scales (i = 2, see Table 1). This behavior was seen for all tested algorithm parameters and therefore, it is not likely to be an artifact. In the data these smallest structures trace only the densest clumps, which are predominantly located along the dense spine of the filament, but not in the surrounding low column density gas. This suggests, that only in the densest parts the filament is able to fragment into the smallest scales.

Table 1 shows the properties of structures at each scale *i*: the total number of identified structures $N_{\rm strc}$, the total mass of these structures $\sum (M_{\rm strc})$, the median hydrogen number density $\tilde{n}({\rm H})$, and the median separation \tilde{s} . The sum of the masses over all scales, including scale i > 8, results in a total cloud mass of about $M_{\rm Nessie}^{\rm scales} = 4.9 \cdot 10^4 {\rm M}_{\odot}$. This is slightly higher than the mass derived from the combined column density map (see Section 4.2). The difference is a consequence of the used spatial filtering algorithm, which may not reproduce the true shapes of the structures accurately.

We include in the fragmentation analysis all structures identified at scales i = 2 - 8. We only include structures within the Nessie filament area (see the polygon in Fig. 5). We computed the projected nearest neighbor distances of the structures. The separation distributions of the scales i = 2, 3 are shown in Fig. 9. They are non-Gaussian in shape and we adopt the median separation as a diagnostic of the separations (given in Table 1).

For the fragmentation analysis an estimate of the structure density is interesting; we estimate this from the outputs of the clumpfind-2D algorithm. The size of a structure was given by clumpfind-2D as the number of pixels, $N_{\rm pix}$, in the FWHM area. For the calculation of the structure volume we assume the shape of a prolate spheroid, that has been found to be among the shapes that best quantify the structures at the scales we are looking at (e.g., Kainulainen et al. 2014). The depth of the prolate spheroids is estimated as the shorter of the projected x and y dimensions. Therefore, the volume of a fragment is

$$V = 4/3 \pi \cdot x \cdot y \cdot \min\{x, y\} . \tag{9}$$

The average column density, $\overline{N}(H)$, is given by: $\overline{N}(H) = N(H)_{tot}/N_{pix}$, and therefore, the hydrogen number density of one structure is: $n(H) = \overline{N}(H) \cdot \pi \cdot x \cdot y/V$. The median number density and the 95% interval for structures at each scale are shown as a function of their median separation in Fig. 10.

Additionally, we estimated the median separation and density from the HNC molecular line observations of Jackson et al. (2010). We used the shown positions to estimate their separation at the distance of d = 3.5 kpc. The density was calculated assuming a spherical geometry with a radius of $r = \sqrt{\Omega/\pi}$, using the angular size Ω of the identified clumps, and their mass M. The hydrogen number density is given by:

$$n(H) = \frac{M}{\mu_{\rm H} m_{\rm H} (4/3\,\pi r^3)} , \qquad (10)$$



Fig. 9. Distributions of the separations (left) and densities (right) of the structures identified from the scale maps i = 2 (top) and i = 3 (bottom). The dashed line indicates the median and the dotted lines the 95% quantiles of each distribution.

where $\mu_{\rm H} = 1.4$ is the mean molecular weight of the interstellar medium with respect to atomic hydrogen and $m_{\rm H}$ is the mass of a hydrogen atom.

We estimated the uncertainty of the median separations and median mean densities using bootstrapping, because their probability distributions are not Gaussian (see Fig. 9). For the separation and mean density on every scale, we drew a new sample of values from among the observed values of separations and mean densities. This new sample had the same amount of data points as originally detected at that scale. We then calculated the median of these new, simulated samples. The resulting distribution of the median values then estimates the sampling function of the observed median and was used to estimate the uncertainty using the standard deviation. The uncertainties vary between 1% and 14% for the separation and between 1% and 25% for the density on scales of i = 3 and i = 8. The uncertainty values of all scales are given in Table 1.

The scatter shown in the separation density plot represent the 95% quantiles of the measured parameters. Large uncertainties, which are neglected here, are the opacity at different wavelength (J, H, K, 8 μ m) and their ratios contributing in the extinction measurement and the conversion factor from extinction to column density. For measuring masses also the uncertainty of the distance, as discussed before, introduces an significant contribution. For more detail see Kainulainen et al. (2011); Kainulainen & Tan (2013).

The density-separation relation (Fig. 10) shows a clear decrease of the mean densities for larger separations. We perform a linear least-square-fit in the log-log space to the data, which represents a power law of the form $\tilde{n}(\mathrm{H}) = A \cdot \tilde{s}^{\,p}$ as $log(\tilde{n}(\mathrm{H})) = p \cdot log(\tilde{s}) + log(A)$. The resulting parameters are $p = -0.96 \pm 0.05$ and $log(A) = 3.22 \pm 0.02$, which is $A = 1669^{+91}_{-86} \mathrm{cm}^{-3}$. The fitted model is shown as black line in figure 10.

A commonly used fragmentation model is the spherical Jeans' instability model (Jeans 1902), where the separation is linked to the mean density $\overline{\rho}$ via the Jeans' length

$$l_J = c_s (\pi/(G\overline{\rho}))^{1/2} , \qquad (11)$$

where c_s is the sound speed within the medium, and G the gravitational constant. We compute the prediction from this assuming a gas temperature of T = 15 K. At all scales, the observed mean separations are in agreement with the Jeans' scale within a factor of ~ 3 . However, for the smallest scales i = 2 - 4 the

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Fig. 10. Median number density of structures at different spatial scales as a function of their median separation. Measurements of this study are marked with crosses. The square marks the data point derived from HNC observations of Jackson et al. (2010). The error bars show the 95% quantiles of both measurements. The blue lines indicate the scale dependency of a infinite long cylinder in the non-turbulent case (solid), and non-thermal case (dash-dotted), and the dashed, red line indicates the scale dependency of Jeans' fragmentation. The black line shows a power law fit to the data.

measurements are systematically below the predicted relationship and for the largest scales systematically above (see the discussion about the slope of the relationship later in this section).

A more shallow slope of the Jeans' fragmentation can be achieved by assuming a non-isothermal medium (e.g. Takahashi et al. 2013). The innermost dense ($\sim 10^4 \text{ cm}^{-3}$) regions of the cloud are shielded from the interstellar radiation field and therefore, can reach temperatures down to 10 K. As the surrounding low density gas ($\sim 10^2 \text{ cm}^{-3}$) is exposed to the radiation, we assume a higher temperature of 20 K. This leads to a slope of about -1.7, which still does not solve the systematic deviations from the observation.

Another commonly used model describes the fragmentation an infinitely long, self-gravitating cylinder (Chandrasekhar & Fermi 1953; Inutsuka & Miyama 1992). This model predicts the separation, λ , depending on the scale-height $H = \sigma_{\rm v} (4\pi G \rho_c)^{-1/2}$, where ρ_c is the central density of a filament in virial equilibrium, $\sigma_{\rm v}$ the velocity dispersion of the medium, and G the gravitational constant. In the case of a non-turbulent medium, the velocity dispersion $\sigma_{\rm v}$ is given by the sound speed c_s within the medium (we assume T = 15 K to calculate the sound speed). In the regime of the filament radius $R \gg H$ the separation is given by $\lambda = 22 H$. If we assume a central density at the largest scale of $n_c(H) \approx 10^3$ cm⁻³, then we derive a scalehight of $H \approx 0.15$ pc. This is smaller than the typical radius of Nessie, $R \approx 1.5$ pc (see Section 4.2). Therefore, the separation is predicted to be

$$\lambda = 22 \cdot c_s (4\pi G\rho_c)^{-1/2} , \qquad (12)$$

which is shown in Fig. 10 and it is in agreement with the measurements within a factor of ~ 3 for scales larger than i = 5, but systematically above the measured densities. However, the model predicts central densities while we derived mean densities, and therefore, the model predicts an upper limit of the mean densities.

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The above models describe fragmentation in non-turbulent medium. However, observations show that high line mass filaments have a non-thermal linewidth (Jackson et al. 2010; Kainulainen et al. 2013), which is higher than the sound speed c_s in the non-turbulent case. Larson (1981) found an relation between the size of a molecular cloud and its observed linewidth. Such a linewidth-size relation might also apply to the here observed structures and therefore, we adopted a typical relation of $\sigma_{\rm v} = 0.72 \ {\rm km \ s^{-1}} \cdot (\lambda/1 \ {\rm pc})^{0.5}$ (Solomon et al. 1987; Heyer & Brunt 2004; Pillai et al. 2006; Shetty et al. 2012; Colombo et al. 2015), where the linewidth $\sigma_{\rm v}$ depends on the observed size scale λ . The non-thermal linewidth exceeds non-turbulent motion, given by the sound speed c_s , at large scales. But the linewidth-size relation can also be partially explained by the non-isothermal behaviour of the gas.

$$\lambda^{0.5} = 22 \cdot 0.72 \,\,\mathrm{km \, s}^{-1} (4\pi G \rho_c)^{-1/2} \,\,, \tag{13}$$

where ρ_c is the central density of a filament in virial equilibrium, and G the gravitational constant (Fig. 10).

Therefore, the relation between the central density and the separation is $\rho_c \propto \lambda^{-1}$, which is in agreement with the observed slope of $p = -0.96 \pm 0.05$. However, again we have to mention that the model predicts central densities while we derived mean densities. Additionally, without informations about the kinematics of the cloud, we cannot constrain the scaling velocity of the linewidth-size relation.

4.4. Comparison with ATLASGAL

We describe shortly how the parsec-scale structures identified in Nessie from ATLASGAL data (resolution of 18", Schuller et al. 2009) break down into substructures when viewed at about 10 times finer resolution of the extinction data. For this, we considered the 16 sources from the ATLASGAL GCSC catalog (Csengeri et al. 2014) that are likely embedded in the cloud. We calculated the number of structures within the FWHM ellipse of the ATLASGAL sources at the two smallest scales (i = 2, 3) of the extinction map (see Fig. 11). We also estimated the mass of the ATLASGAL clumps by adopting Equation (2) and assuming a dust temperature of $T_d \approx 15$ K. These masses are then compared to the total mass of the small scale structures. The resulting ratios are shown in Tab. 2.

In particular, we found that, on average, the number of small scale structures within the half power ellipse of the clump is $\overline{N}_{\text{strc},2} = 2.9$ and $\overline{N}_{\text{strc},3} = 2.8$. These contain 2% and 6% of the mass of the ATLASGAL clump. The half power ellipses of the clumps and the i = 2 structures identified within the clumps are shown in Fig. E.1 overlaid on the extinction map. While half of the ATLASGAL clumps are clearly associated with high extinction peaks, especially the four most massive ones (> 500 M_{\odot}) contain no or only low extinction peaks. This is dominantly because of the caveats of the extinction mapping technique. The massive clumps commonly exhibit MIR emission of Polycyclic Aromatic Hydrocarbons (PAHs) in the 8 μ m band (Benjamin et al. 2003); this interferes with the extinction mapping procedure. Also bright foreground stars cause a lack of mid-infrared extinction and influence our results. In total this likely leads to an underestimated number of substructures per clump and to underestimating some of their masses. It also shows that our method is excellent for identifying the youngest and densest regions, but it starts to fail as soon as star formation progresses and the regions show strong MIR emission.

Table 1. Results of the fragmentation analysis

Scale i	Scale [pc]	$N_{\rm strc}$	$\frac{\sum(M_{\rm strc})}{[10^3 {\rm M}_{\odot}]}$	$\overline{M}_{\rm strc}$ [M _☉]	$\widetilde{n}({\rm H})^{a}$ [10 ³ cm ⁻³]	$\frac{\sigma(\widetilde{n}(\mathrm{H}))}{[10^3 \mathrm{~cm^{-3}}]}$	\widetilde{s}^b [pc]	$\sigma(\widetilde{s})$ [pc]	$\lambda_{ m J}$ [pc]
>8	>5.2	1	34.5						
8	5.2	11	4.0	373	0.13	0.04	5.0	0.8	2.7
7	2.6	31	3.2	108	0.61	0.05	2.5	0.2	1.2
6	1.3	72	2.4	33.2	1.4	0.12	1.4	0.1	0.81
5	0.65	242	2.0	8.28	3.1	0.12	0.73	0.02	0.55
4	0.33	903	1.9	2.06	4.1	0.07	0.38	0.004	0.48
3	0.16	1751	1.2	0.66	6.3	0.08	0.23	0.002	0.38
2	0.08	523	0.20	0.40	14.2	0.47	0.17	0.004	0.26

Notes. ^(a) median of the mean density of the identified structures ^(b) median of the separation between identified structures

Table 2. ATLASGAL GCSC clumps (Csengeri et al. 2014) likely embedded in the Nessie cloud

Name	Size	PA	$M_{\rm clump}$	$N_{\rm strc, 2}$	$M_{\rm strc, 2}$	$\frac{M_{\text{strc, 2}}}{M_{\text{slump}}}$	$N_{\rm strc, 3}$	$M_{\rm strc, 3}$	$\frac{M_{\rm strc, 3}}{M_{\rm slump}}$
	["]	[°]	$[M_{\odot}]$		$[M_{\odot}]$	eiump		$[{\rm M}_{\odot}]$	eiump
G338.9380-0.4231:	$46\ge 20$	-12	221	4	3.15	0.014	4	12.33	0.056
G338.9362-0.4808:	$28 \ge 22$	52	197	2	1.36	0.007	1	4.69	0.024
G338.9371-0.4919:	$41 \ge 34$	134	1094	3	3.27	0.003	2	10.77	0.010
G338.9275-0.5018:	$39 \ge 26$	102	523	0	0.00	0.000	3	7.77	0.015
G338.8688-0.4796:	$32 \ge 23$	71	248	5	4.34	0.018	1	15.63	0.063
G338.7790-0.4591:	$39 \ge 23$	-24	176	4	3.97	0.022	4	15.95	0.090
G338.7314-0.4691:	$32 \ge 19$	90	116	3	4.76	0.041	3	13.14	0.114
G338.5519-0.4190:	$27\ge 24$	71	134	2	2.57	0.019	4	7.65	0.057
G338.4236-0.4101:	$28 \ge 26$	111	292	0	0.00	0.000	2	2.20	0.008
G338.3937-0.4053:	$42 \ge 31$	72	632	2	2.22	0.004	3	10.43	0.016
G338.3923-0.3972:	$34 \ge 19$	16	124	2	1.75	0.014	3	9.91	0.080
G338.3271-0.4096:	$36 \ge 27$	-20	534	4	3.72	0.007	3	11.02	0.021
G338.1991-0.4642:	$27\ge 25$	36	181	2	2.61	0.014	3	10.09	0.056
G338.1122-0.4632:	$41 \ge 25$	62	202	6	8.27	0.041	5	20.46	0.101
G338.0892-0.4474:	$30 \ge 25$	65	147	3	6.57	0.045	1	16.08	0.109
G338.3048-0.5223:	$47 \ge 22$	95	216	4	4.21	0.019	3	15.57	0.072
mean:			315	2.88	3.30	0.017	2.81	11.48	0.056
stddev:			261	1.63	2.16	0.014	1.17	4.66	0.037

5. Discussion

5.1. Scale-dependent fragmentation of Nessie

In the following, we discuss the scale-dependent fragmentation of Nessie (Fig. 10) in the context of the analytic gravitational fragmentation models. We showed that the upper limit of the average line-mass of Nessie is $(M/l) = 627 \text{ M}_{\odot} \text{ pc}^{-1}$. For a thermally supported filament at a temperature of T = 15 K the critical line-mass is $(M/l)_{\text{crit}} = 20 \text{ M}_{\odot} \text{ pc}^{-1}$. Thus, the filament is clearly thermally supercritical. There are no analytic theories that would self-consistently explore the evolution of such highly thermally super-critical filaments.

In the absence of directly applicable models, a common approach in the recent literature is to assume that the non-thermal motions provide a straightforward, idealized supporting force for the filament, increasing its critical line-mass (e.g., Jackson et al. 2010; Hernandez et al. 2012; Busquet et al. 2013; Beuther et al. 2015). This commonly leads to a conclusion that the line-masses of high line-mass filaments are close to their critical line-masses. This is true for Nessie, too. Jackson et al. (2010) showed that the non-thermal motions in Nessie increase the critical line-mass to $(M/l)_{\rm vir} = 525 \ {\rm M}_{\odot} \ {\rm pc}^{-1}$, which is similar to our observed value.

Building on the above agreement, observations are commonly compared to the predictions of gravitational fragmentation models developed for near-equilibrium cylinders. These models typically proceed from a static initial configuration with a linear perturbation analysis. In short, such models predict a periodic fragmentation pattern with a specific wavelength, i.e., the fragmentation pattern predicted by the models is not scaledependent. However, the fragmentation wavelength depends on the density of the filaments as described by Eqs. 13, 12, and 11; filaments with different densities have different fragmentation wavelengths. This should be kept in mind when interpreting the relationship between the data and models presented in Fig. 10.

In this context, the observed slope of the mean density – separation relationship in Nessie is in agreement with that of a non-thermal, self-gravitating cylinder that has a Larsonlike linewidth-size relation ($\sigma_v \propto \lambda^{0.5}$, Larson 1981; Solomon et al. 1987; Heyer & Brunt 2004; Shetty et al. 2012; Colombo et al. 2015). As the cloud shows non-thermal velocity dispersions (Jackson et al. 2010), this relation could be a result of turbulent motions within the cloud, but also systematic motions, such as collapse, could affect the linewidth. The observed median nearest-neighbor separations of the fragments are within a factor of two of the predictions of the isothermal and non-isothermal Jeans' fragmentation (Jeans 1902). However, the slope is significantly steeper than the observed one. Additionally, on the large scales the separations also are in agreement with the fragmentation model of a non-turbulent, self gravitating, infinite long



Fig. 11. Combined NIR and MIR extinction map $(l = 338.10^{\circ}, b = -0.45^{\circ})$ overlaid with the half power contour of two ATLASGAL GCSC sources (black ellipses) and their covered sources identified with clumpfind-2D from the scale 2 map. The white lines show the contours of the ATLASGAL emission.

cylinder Chandrasekhar & Fermi (1953); Inutsuka & Miyama (1992), but again the slope of the model is significantly steeper than observed. Note, that the cylindrical models predict central densities, which can only be seen as upper limits for the derived mean densities.

Previously, a change of fragmentation mode between large and small scales has been seen at the size-scale of $\sim 0.5 \text{pc}$, e.g., in the studies of the young high-mass cloud G11.11-0.12 (Kainulainen et al. 2013), the Taurus cloud (Hacar et al. 2013), and the integral-shaped filament in Orion (Teixeira et al. 2016; Kainulainen et al. 2017). While we do not detect such feature in Nessie, the data are in agreement with the presence of such a feature, i.e., cannot rule it out (c.f., Fig. 10). One possible explanation for the change of fragmentation modes can be changing influence of the environment (Pon et al. 2011). While on large scale fragmentation is driven by the characteristics of the cylindrical, filamentary structure, the smaller scales approach a more spherical shape, which is independent of larger scales. Also, recent numerical simulations have explored possibilities to explain scale-dependent fragmentation through dynamical processes (e.g., Clarke et al. 2017; Gritschneder et al. 2017).

5.2. Star formation potential

Ultimately, one would like to link the fragmentation in Nessie to star formation. To take the first step towards this, we estimated the young stellar object (YSO) content of Nessie using publicly available multi-band photometric catalogs. The detailed methods used to identify the YSOs and estimate the SFR are explained in Zhang et al. 2017, submitted. Here we give a short description of the method.

For the YSO selection we used near-infrared data (we did the PSF photometry on VVV images, VISTA Variables in the Via Lactea, Saito et al. 2012), Spitzer GLIMPSE (Galactic Legacy Mid-Plane Survey Extraordinaire, Benjamin et al. 2003;

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Churchwell et al. 2009) and MIPSGAL (Multiband Imaging Photometer Galactic Plane Survey, Carey et al. 2009; Gutermuth & Heyer 2015) archival catalogues, AllWISE catalogue (Wide-field Infrared Survey Explorer, Wright et al. 2010), Herschel Hi-GAL catalogue (Herschel infrared Galactic Plane Survey, Molinari et al. 2010, 2016), and Red MSX source catalogue (Midcourse Space Experiment, Lumsden et al. 2013, used to include massive protostars) and the methods from Gutermuth et al. (2009); Koenig & Leisawitz (2014); Saral et al. (2015); Robitaille et al. (2008); Veneziani et al. (2013). Our YSO selection scheme uses the SEDs of sources from 1 to 500 μ m and can efficiently mitigate the effects of contamination. In Nessie, we finally obtain 298 sources with the excessive infrared emission, of which 35 are classified as AGB candidates using the multi-color criteria.

Considering the distance of Nessie, it is necessary to correct the flux densities of the YSO candidates for extinction. We use the method suggested by Fang et al. (2013); Zhang et al. (2015) to estimate the foreground extinction towards each YSO candidate and de-redden their photometry. Here we also give a short description about this method.

- 1 For the sources with J, H, K_S detections, the extinction is obtained by employing the JHK_S color-color diagram. Figure 12 shows the J-H versus H-K_S color-color diagram of the YSO candidates in Nessie. Given the different origins of intrinsic colors of YSO candidates, the color-color diagram is divided into three subregions. In region 1, the intrinsic color of $[J-H]_0$ is simply assumed to be 0.6; in region 2, the intrinsic color of a YSO is obtained from the intersection between the reddening vector and the locus of main sequence stars (Bessell & Brett 1988); in region 3, the intrinsic color is derived from where the reddening vector and the classical T Tauri star (CTTS) locus (Meyer et al. 1997) intersects. Then the extinction values of YSO candidates are estimated from observed and intrinsic colors with the extinction law of Xue et al. (2016).
- 2 For other sources (outside these three regions or without detections in JHK_S bands), their extinction is estimated with the median extinction values of surrounding Class II sources that have extinction measurements in step 1.

Using the de-reddened SEDs, we re-classify the YSO candidates into Class I, Flat, and Class II sources based on their spectral indices and bolometric temperatures (Greene et al. 1994; Chen et al. 1995). Figure 13 shows the $K_S - [8.0]$ versus J-H color-color diagrams before and after de-reddening for Class I+Flat and Class II sources in Nessie.

Although we have removed some contamination during the YSO selection process, our YSO candidates in Nessie are still contaminated by the foreground and background sources.

The foreground contamination mainly includes the foreground AGBs and the foreground YSOs which are associated with the molecular clouds that are located between us and Nessie. We use the A_V values of YSOs obtained previously and the 3D extinction map (Marshall et al. 2006) to isolate the foreground contamination. Based on the distance of Nessie, we can estimate the foreground extinction in different lines of sight towards Nessie with the 3D extinction map. If the extinction value of a YSO is lower than the corresponding foreground extinction of Nessie, this YSO would have high probability to be a foreground contamination. We checked the YSOs in Nessie and marked the possible foreground contamination using this method. The fraction of foreground contamination in Nessie is 10% in Class I+Flat sources and 9% in Class II sources.

Our YSOs are also contaminated by background sources, including extragalactic objects, background AGBs, and background YSOs which are associated with the molecular clouds that are located behind Nessie. We think that the extragalactic contamination is not important in our YSOs because we are observing through the Galactic plane. Many background AGBs



Fig. 12. The H-K_S versus J-H color–color diagram for the YSO candidates in Nessie. The solid curves show the intrinsic colors for the main-sequence stars (black) and giants (red; Bessell & Brett 1988), and the dash–dotted line is the locus of T Tauri stars from Meyer et al. (1997). The dashed lines show the reddening direction, and the arrow shows the reddening vector. The extinction law we adopted is from Xue et al. (2016). Note that the dashed lines separate the diagram into three regions marked with numbers 1, 2, and 3 in the figure. We use different methods to estimate the extinction of YSO candidates in different regions (see the text for details).

have been removed using the multi-color criteria during the YSO identification process. The residual contamination of background AGBs is estimated with the control fields. We select five nearby fields with weak CO emission as the control fields and apply the YSO selection scheme to all the control fields to select YSOs. Assuming that there is no YSOs in each control field, all selected 'YSOs' in the control fields are actually contamination of AGBs (if neglecting the extragalactic contamination). With an assumption of a uniform distribution for AGB stars, we can estimate the number of residual background AGBs in the Nessie using the mean value of the surface density of background AGBs in five control fields. Combining the numbers of background AGBs identified by color criteria and estimated using control fields, we found that the fraction of background contamination is 22% in Class I+Flat sources and 11% in Class II sources. Note that we did not try to eliminate the contamination from background YSOs because they are difficult to remove without the information of radial velocities of YSOs.

After removing the contamination, we obtain 51 Class I and flat spectrum objects and 137 Class II sources in the Nessie. In order to calculate the star formation rate (SFR), we must estimate the total mass of YSOs in Nessie. In this work, we use different methods to estimate the total mass of Class I+Flat and Class II populations:

– We use the de-reddened photometry of Class II sources in Nessie to estimate the flux completeness. Figure 14 shows the K_S absolute magnitude histogram of Class II sources in Nessie. We simply adopt the peak position of histogram as the completeness of K_S band (~ 1 mag). Figure 15 shows the $M_{\rm K_S} - M_*$ relation for Class II sources constructed from YSO models presented by Robitaille et al. (2006). Using this relation, we transfer the K_S band completeness to the mass completeness of 1.48±0.65 M_{\odot} . Assuming a universal IMF (Kroupa 2001), we estimated the number of Class II sources to be 1282^{+1228}_{-614} and the total mass of Class II sources to be $698.4^{+711.8}_{-355.9}$ M $_{\odot}$.

For Class I+Flat sources, we used the observed luminosity functions constructed by Kryukova et al. (2012) as the template to estimate the total number of Class I+Flat sources. We calculate the bolometric luminosities of Class I+Flat sources using the trapezoid rule to integrate over the finitely sampled de-reddened SEDs (Dunham et al. 2008, 2015). Figure 16 shows the de-reddened luminosity function of Class I+Flat sources in Nessie and the corresponding luminosity completeness that is calculated with the method suggested by Kryukova et al. (2012) is also marked with the red line. As a comparison, we also plot the luminosity function of Class II sources in the Nessie. Assuming an universal luminosity function, we estimate the total number of Class I+Flat sources in Nessie to be 185^{+52}_{-51} . Assuming the average mass of 0.5 solar mass for each Class I/Flat source, we estimated the total mass of Class I+Flat sources to be $92.7^{+25.8}_{-25.7}\,\mathrm{M}_\odot.$

Adopting the lifetime of Class II sources, 2 Myr (Evans II et al. 2009), as the star formation time-scale, we obtain a star formation rate of SFR = $389^{+364}_{-182} \,\mathrm{M_{\odot} \, Myr^{-1}}$ for Nessie. The star formation efficiency within the star-forming time-scale is estimated by the total mass of YSOs, $M_{\rm YSOs}$, and the gas mass of Nessie, $M_{\rm Nessie}$, SFE = $M_{\rm YSOs}/(M_{\rm Nessie}+M_{\rm YSOs}) = 0.018^{+0.017}_{-0.008}$. The uncertainty is mainly from the uncertainty of transferring K_s magnitudes to stellar masses and the small number of observed Class I and Class II sources. To place these values in context, the SFR of Nessie is comparable to those of the most active nearby star forming regions like Perseus (150 M_{\odot} \,\mathrm{Myr^{-1}}), Orion A (715 M_{\odot} \,\mathrm{Myr^{-1}}) and Orion B (159 \,\mathrm{M_{\odot} \, Myr^{-1}}; all values from Lada et al. 2010).

It is immediately interesting to compare this direct SFR estimate to other measures commonly linked with the star formation potential of molecular clouds. One such measure is the mass of dense gas in the cloud (e.g., Kainulainen et al. 2009; Lada et al. 2010). Specifically, Lada et al. (2010) found that in the Solar Neighborhood clouds (distance ≤ 500 pc) SFRs correlate best with the mass above a column density threshold of $A_V \approx 7.3$ mag. Adopting this threshold results in the dense gas mass of $M_{\rm dg} = 8.7 \cdot 10^3 \, {\rm M}_{\odot}$ in Nessie. Following the prescription of Lada et al. (2010) for the Solar Neighborhood clouds, the SFR of SFR = $4.6 \cdot 10^{-8} \, {\rm yr}^{-1} \cdot M_{\rm dg} = 400 \, {\rm M}_{\odot} \, {\rm Myr}^{-1}$ follows. This is in agreement with the SFR derived from the YSOs; in Nessie the mass of dense gas above $A_V \approx 7.3$ mag is a reasonable predictor of the SFR.

Yet another measure commonly connected with SFR is the dense core population of the molecular clouds (e.g., Motte et al. 1998; Alves et al. 2007; Marsh et al. 2016). To analyze this population in Nessie, we can take the advantage of the high spatial resolution of our column density map: we can directly count the cores that might form stars or multiple stellar systems and estimate their mass. The mass enclosed in the dense structures smaller than ~ 0.1 pc is likely to take part in star formation processes. Therefore, the number of structures at the smallest scale of the wavelet-filtered map $(i = 2, \sim 0.08 \text{pc})$ provides a first-order estimate for the number of stars forming in the cloud in the near future. To account for possible accretion processes during the collapse of a core, we assume the gas at the scales i = 2 and i = 3 (size < 0.16 pc) can participate in the collapse. This will then give an upper limit for the mass available for star formation. The mass of stars formed by these cores is then estimated by assuming a star formation efficiency of SFE = 30%(e.g., Alves et al. 2007; Rathborne et al. 2009; André et al. 2010). This results in the stellar mass of $M_{i=2,3} = 409 \text{ M}_{\odot}$. Adopting



Fig. 13. The observed (*left*) and de-reddened (*right*) $K_S - [8.0]$ versus J-H color-color diagrams for Class I+Flat (red) and Class II (green) sources in Nessie. The black arrows show the extinction vectors.



Fig. 14. K_S absolute magnitude (M_{K_S}) histogram of Class II sources in Nessie.

again the star formation time of $t_{\rm SF} \approx 2$ Myr leads to a star formation rate of SFR = $M_*/t_{\rm SF} = 205 \ {\rm M}_{\odot} \ {\rm Myr}^{-1}$ for the Nessie cloud. This estimate is within a factor of 2 of the values derived previously. We can also simply use the number of detected cores to gain a crude estimate of the star formation potential. If we assume that each structure at scale i = 2 will form at least one star, Nessie will form 523 stars. This is within a factor of two of the actual number of (completeness corrected) Class I and II sources. If we further divide the total mass in the cores in Nessie by 523, the predicted average mass of a star of 0.78 M $_{\odot}$ follows; this is relatively close to the mean stellar mass of 0.5 M $_{\odot}$ of the initial mass function (e.g., Kroupa 2002). Altogether, the above considerations suggest that the dense core population identified from Nessie using the approach of this paper is a reasonable proxy of Nessie's star formation potential.





Fig. 15. The relation between stellar mass and K_S absolute magnitude of Class II source. The black dots represent the Robitaille et al. (2006) Stage 2 models with $0.001 < M_{\rm disk}/M_* < 0.01, 0.08 < M_* < 7 M_{\odot}$, and 30° < inclination angle < 60° . The red curve shows the robust polynomial fitting while the gray region shows the 1 σ uncertainty of the fitting. The CTTS in L1641 from Fang et al. (2013) are marked with green filled circles. Most of CTTS are located in the gray region, which confirms that this $M_{\rm K_S} - M_*$ relation for Class II sources is consistent with the observational results.

6. Conclusions

We analyzed the column density structure of the (projected) 67 pc long filamentary Nessie cloud using a combined near- and mid-infrared extinction mapping method on data of the VVV survey and 8 μ m *Spitzer*/GLIMPSE images. Our results are as follows:

1. We derived a high-resolution (~ 0.03 pc), high dynamic range $(N({\rm H_2})$ = 3 – 100 $\cdot\,10^{21}~{\rm cm^{-2}})$ column density map



Fig. 16. De-reddened luminosity functions of Class I+Flat (top panel) and Class II (bottom panel) sources in Nessie. The red vertical line shows the de-reddened luminosity completeness.

for Nessie and estimated the distance towards it to be d = 3.5 kpc based on near-infrared source-counts. The mass of Nessie is $4.2 \cdot 10^4 \text{ M}_{\odot}$, considering regions above $N(\text{H}_2) \gtrsim 3 \cdot 10^{21} \text{ cm}^{-2}$. This leads to the mean line-mass of about 627 $M_{\odot} pc^{-1}$.

- We analyzed the fragmentation of the cloud across a wide 2.range of scales between 0.1 - 10 pc and detected fragmentation at all scales. We characterize the fragments and find that their masses decrease and densities increase as a function of size-scale. At the smallest scale, the typical masses of the fragments are 0.4 M_{\odot} and mean densities are $\sim 10^4$. The mean densities of the fragments decrease with their nearestneighbor separations, following approximately a power-law with an exponent of -0.96 ± 0.05 . The previous determination of the 4 pc fragmentation length by Jackson et al. (2010) is in agreement with this relationship, however, our data shows that determining the fragmentation length at any one particular scale does not capture the full, scale-dependent picture of fragmentation in Nessie.
- 3. In the context of analytic gravitational fragmentation models, the observed nearest-neighbor separations are within a factor of two of the Jeans' length at all size-scales. However, the slope of the observed mean density - separation relationship is significantly shallower than the scale-dependency of the Jeans' length. The observed relationship is in agreement with a gravitationally fragmenting near-equilibrium cylinder that is supported by non-thermal motions that exhibits a Larson-like velocity-size scaling, i.e., a power-law with an exponent of 0.5. This scaling could result, e.g., from turbulent motions in the cloud, because the cloud shows clearly non-thermal velocity dispersions (Jackson et al. 2010).
- 4. We estimated the SFR of Nessie to be $389 \ M_{\odot} \ Myr^{-1}$ based on the number of identified YSOs in the cloud. An estimate based on the number of $\sim~0.1~pc$ scale column density "cores" yields 205 $\rm M_\odot\,Myr^{-1}.$ We also estimate the SFR based on the total amount of dense gas $(A_V > 7.3 \text{ mag})$; Lada et al. 2012) in the cloud, resulting in 400 M_{\odot} Myr⁻ These results suggest that both the number of dense cores and the amount of dense gas above $A_{\rm V} > 7.3$ mag are relatively good proxies of the star-forming content of Nessie. We

further derive the SFE of 0.018 for Nessie. These numbers indicate that the star-forming content of Nessie is similar to the Solar neighborhood giant molecular clouds like Orion A.

5. The ATLASGAL clumps identified in Nessie typically harbor 2 - 3 small-scale structures (< 0.16 pc). These structures contain about 7% of the mass of the parental clump. However, this is a lower limit as the extinction mapping is susceptible for incompleteness arising from mid-infrared bright objects, like foreground stars, and warm/hot gas.

We showed that the filamentary Nessie cloud has scaledependent fragmentation characteristics. These characteristics are in agreement with some of the predictions of gravitational fragmentation models. However, self-consistent scale-dependent fragmentation models are needed to gain understanding of the structure and evolution of filamentary clouds.

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References

- Abreu-Vicente, J., Ragan, S., Kainulainen, J., et al. 2016, A&A, 590, 20 pp
- Alves, J., Lombardi, M., & Lada, C. J. 2007, A&A, 462, L17
- André, P., Di Francesco, J., Ward-Thompson, D., et al. 2014, Protostars and Planets VI: From Filamentary Networks to Dense Cores in Molecular Clouds : Toward a New Paradigm for Star Formation (University of Arizona Press), 24
- André, P., Men'shchikov, A., Bontemps, S., et al. 2010, A&A, 518, 1
- Arzoumanian, D., André, P., Didelon, P., et al. 2011, A&A, 529, 1
- Benjamin, R. A., Churchwell, E., Babler, B. L., et al. 2003, ASP, 115, 953
- Bessell, J. & Brett, M. 1988, ASP, vol. 100, 1134
- Beuther, H., Ragan, S. E., Johnston, K., et al. 2015, A&A, 584, A67 Bohlin, R. C., Savage, B. D., & Drake, J. F. 1978, ApJ, 224, 132
- Busquet, G., Zhang, Q., Palau, A., et al. 2013, ApJ, 764, L26
- Butler, M. J. & Tan, J. C. 2012, ApJ, 754, 5
- Cardelli, J. A., Clayton, G. C., & Mathis, J. S. 1989, ApJ, 345, 245 Carey, S. J., Noriega-Crespo, A., Mizuno, D. R., et al. 2009, PASP,
- 121, 76 Chandrasekhar, S. & Fermi, E. 1953, ApJ, 118, 116C
- Chen, H., Myers, P. C., Ladd, E. F., & Wood, D. O. S. 1995, ApJ, 445.377
- Churchwell, E., Babler, B. L., Meade, M. R., et al. 2009, ASP, 121
- Clarke, S. D., Whitworth, A. P., Duarte-Cabral, A., & Hubber, D. A. 2017, MNRAS, 468, 2489

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- Colombo, D., Rosolowsky, E., Ginsburg, A., Duarte-Cabral, A., & Hughes, A. 2015, MNRAS, 454, 2067
- Csengeri, T., Urquhart, J. S., Schuller, F., et al. 2014, A&A, 565, A75 Cutri, R. M., Skrutskie, M. F., van Dyk, S., et al. 2003, "The IRSA 2MASS All-Sky Point Source Cat.
- Dunham, M. M., Allen, L. E., Evans, II, N. J., et al. 2015, ApJS, 220, 11
- Dunham, M. M., Crapsi, A., Evans, II, N. J., et al. 2008, ApJS, 179, 249
- Evans II, N. J., Dunham, M. M., Jørgensen, J. K., et al. 2009, ApJ, 181, 321
- Fang, M., Kim, J. S., van Boekel, R., et al. 2013, ApJS, 207, 39 pp
- Froebrich, D., Murphy, G. C., Smith, M. D., Walsh, J., & Del Burgo,
- C. 2007, MNRAS, 378, 1447 Goldsmith, P. F., Heyer, M., Narayanan, G., et al. 2008, ApJ, 680, 428
- Goodman, A. A., Alves, J., Beaumont, C. N., et al. 2014, ApJ, 797, 53
- Goodman, A. A., Pineda, J. E., & Schnee, S. L. 2009, ApJ, 692, 91
- Greene, T. P., Wilking, B. A., Andre, P., Young, E. T., & Lada, C. J. 1994, ApJ, 434, 614
- Gritschneder, M., Heigl, S., & Burkert, A. 2017, ApJ, 834, 202
- Gutermuth, R. A. & Heyer, M. 2015, AJ, 149, 64
- Gutermuth, R. A., Megeath, S. T., Myers, P. C., et al. 2009, ApJS, 184, 18
- Hacar, A., Tafalla, M., Kauffmann, J., & Kovács, A. 2013, A&A, 554, A55
- Hennebelle, P. & Falgarone, E. 2012, A&AR, 20, 55
- Henshaw, J. D., Caselli, P., Fontani, F., et al. 2016, MNRAS, 463, 146 Hernandez, A. K., Tan, J. C., Kainulainen, J., et al. 2012, ApJ, 756,
- L13
- Heyer, M. & Brunt, C. 2004, ApJ, 615, L45
- Inutsuka, S. & Miyama, S. M. 1992, ApJ, 338, 392
- Ioannidis, G. & Froebrich, D. 2012, MNRAS, 425, 1380
- Jackson, J. M., Finn, S. C., Chambers, E. T., Rathborne, J. M., & Simon, R. 2010, ApJ, 719, L185 Jeans, J. H. 1902, Philos. Trans. R. Soc. A Math. Phys. Eng. Sci.,
- 199, 1
- Johnstone, D., Fiege, J., Redman, R. O., et al. 2003, ApJ, 588, L37 Juvela, M., Pelkonen, V.-M., Padoan, P., & Mattila, K. 2008, A&A, 480.445
- Kainulainen, J., Alves, J., Beuther, H., Henning, T., & Schuller, F. 2011, A&A, 536, A48
- Kainulainen, J., Beuther, H., Henning, T., & Plume, R. 2009, A&A, 508, L35
- Kainulainen, J., Federrath, C., & Henning, T. 2014, Science, 344, 183
- Kainulainen, J., Ragan, S. E., Henning, T., & Stutz, A. 2013, A&A, 557. A120, 1
- Kainulainen, J., Stutz, A. M., Stanke, T., et al. 2017, A&A, 600, A141 Kainulainen, J. & Tan, J. C. 2013, A&A, 549, A53, 1
- Kauffmann, J., Bertoldi, F., Bourke, T. L., Evans, N. J., & Lee, C. W. 2008, A&A, 487, 993
- Koenig, X. & Leisawitz, D. 2014, ApJ, 791, 27 pp
- Kroupa, P. 2001, MNRAS, 322, 231 Kroupa, P. 2002, Science (80-.)., 295, 82
- Kryukova, E., Megeath, S. T., Gutermuth, R. A., et al. 2012, AJ, 144, $28~{\rm pp}$
- Lada, C. J., Forbrich, J., Lombardi, M., & Alves, J. F. 2012, ApJ, 745, 190
- Lada, C. J., Lombardi, M., & Alves, J. F. 2010, ApJ, 724, 687
- Larson, R. B. 1981, MNRAS, 194, 809
- Li, G.-X., Urquhart, J. S., Leurini, S., et al. 2016, A&A, 591, 24 pp
- Lombardi, M. 2005, A&A, 438, 169
- Lombardi, M. & Alves, J. 2001, A&A, 377, 1023 Lombardi, M., Alves, J., & Lada, C. J. 2006, A&A, 454, 781
- Lumsden, S. L., Hoare, M. G., Urquhart, J. S., et al. 2013, ApJS, 208, 11
- Marsh, K. A., Kirk, J. M., André, P., et al. 2016, MNRAS, 459, 342
- Marshall, D. J., Robin, A. C., Reylé, C., et al. 2006, A&A, 453, 635 McKee, C. F. & Ostriker, E. C. 2007, ARA&A, 45, 565
- Menten, K. M., Reid, M. J., Forbrich, J., & Brunthaler, A. 2007, A&A, 474.515
- Meyer, M. R., Calvet, N., & Hillenbrand, L. A. 1997, AJ, 114, 288
- Molinari, S., Schisano, E., Elia, D., et al. 2016, A&A, 591, A149
- Molinari, S., Swinyard, B., Bally, J., et al. 2010, PASP, 122, 314
- Motte, F., Andre, P., & Neri, R. 1998, A&A, 336, 150
- Myers, P. C. 2009, ApJ, 700, 1609
- Ossenkopf, V. & Henning, T. 1994, A&A, 291, 943 Padoan, P., Haugbølle, T., & Nordlund, Å. 2014, ApJ, 797, 20 pp

Article number, page 18 of 30

- Peretto, N. & Fuller, G. A. 2009, A&A, 505, 405
- Pillai, T., Wyrowski, F., Carey, S. J., & Menten, K. M. 2006, A&A, 450, 569
- Pon, A., Johnstone, D., & Heitsch, F. 2011, ApJ, 740, 88
- Rachford, B. L., Snow, T. P., Tumlinson, J., et al. 2002, ApJ, 577, 221
- Ragan, S. E., Bergin, E. A., & Gutermuth, R. A. 2009, ApJ, 698, 324
- Ragan, S. E., Henning, T., Tackenberg, J., et al. 2014, A&A, 568, A73
- Rathborne, J. M., Lada, C. J., Muench, A. A., et al. 2009, ApJ, 699, 742
- Reach, W. T., Rho, J., Tappe, A., et al. 2005, AJ, 131, 1479 Robin, A. C., Reyl, C., & Derri, S. 2003, A&A, 409, 523
- Robitaille, T. P., Meade, M. R., Babler, B. L., et al. 2008, AJ, 136, 2413
- Robitaille, T. P., Whitney, B. A., Indebetouw, R., Wood, K., & Denzmore, P. 2006, ApJS, 167, 256
- Saito, R. K., Hempel, M., Minniti, D., et al. 2012, A&A, 537, A107, 1
- Saral, G., Hora, J. L., Willis, S. E., et al. 2015, ApJ, 813, 21 pp
- Savage, B. D., Bohlin, R. C., Drake, J. F., & Budich, W. 1977, ApJ, 216, 291
- Schisano, E., Rygl, K. L. J., Molinari, S., et al. 2014, ApJ, 791, 27
- Schneider, S. & Elmegreen, B. G. 1979, ApJ, 41, 87
- Schuller, F., Menten, K. M., Contreras, Y., et al. 2009, A&A, 504, 415 Schultz, G. V. & Wiemer, W. 1975, A&A, 43, 133
- Shetty, R., Beaumont, C. N., Burton, M. G., Kelly, B. C., & Klessen, R. S. 2012, MNRAS, 425, 720
- Skrutskie, M. F., Cutri, R. M., Stiening, R., et al. 2006, AJ, 131, 1163
- Solomon, P. M., Rivolo, A. R., Barrett, J., & Yahil, A. 1987, ApJ, 319, 730
- Starck, J.-L. & Murtagh, F. 2002, Astronomical image and data analysis (Springer- Verlag Berlin Heidelberg GmbH), 338 Stetson, P. B. 1987, PASP, 99, 191
- Stutz, A. M. & Gould, A. 2016, A&A, 590, A2
- Takahashi, S., Ho, P. T. P., Teixeira, P. S., Zapata, L. A., & Su, Y.-N. 2013, ApJ, 763, 57 Teixeira, P. S., Takahashi, S., Zapata, L. A., & Ho, P. T. P. 2016,
- A&A, 587, A47
- Veneziani, M., Elia, D., Noriega-Crespo, A., et al. 2013, A&A, 549, 8 pp
- Wang, K., Testi, L., Burkert, A., et al. 2016, ApJS, 226, 9
- Wang, K., Zhang, Q., Testi, L., et al. 2014, MNRAS, 439, 3275
- Wienen, M., Wyrowski, F., Menten, K. M., et al. 2015, eprint arXiv:1503.00007
- Williams, J. P., de Geus, E. J., & Blitz, L. 1994, ApJ, 428, 693
- Wright, E. L., Eisenhardt, P. R. M., Mainzer, A. K., et al. 2010, AJ, 140. 1868 Xue, M., Jiang, B. W., Gao, J., et al. 2016, ApJS, 224, 23
- Zhang, Q., Wang, K., Lu, X., & Jimenez-Serra, I. 2015, eprint arXiv:1503.03017
- Zucker, C., Battersby, C., & Goodman, A. 2015, eprint arXiv:1506.08807



Fig. B.1. Comparison of K_S magnitudes between the stacked and 16 s data. The red line indicates the one-to-one correlation. The shown stars have an photometric uncertainty lower than 0.05 mag.

Appendix A: Used observations

For the calculation of the near-infrared extinction map of the Nessie filament we use observations conducted by the VVV (VISTA Variables in the Via Lactea) survey (Saito et al. 2012) in the JHK_S photometric bands. This calibrated and reduced (science ready) data is publicly available in the ESO archive. The exact observations used in this study are listed in Table A.1.

Appendix B: Photometry of different observations

For the photometry of the near-infrared data we use a set of different observations (see Appendix A), which show different spatial resolutions due to different conditions. Therefore, the pointspread-function (PSF) for point sources will be different in the single observations, and also effect the stacked data. This might be especially relevant in the K_s filter where we use a larger set of observations. To test the significance of this quality difference we compare the results of photometry in the K_s filter performed on tile 068 of the stacked data, one 80 s exposure (from 2010-05-09), and the lowest resolution 16 s exposure (from 2011-08-31). For all data-sets we used the same parameters as described before and also calibrated the found magnitudes with the 2MASS data (Skrutskie et al. 2006; Cutri et al. 2003). We then identified stars seen in the stacked and 16 s data, and stacked and 80 s data, and plotted the derived magnitudes against each other.

We find a good correlation for the three data-sets. However, there is a larger scatter for the 16 s data because of the lower sensitivity of the data. In general, we find a significantly increased number of stars for the longer exposure or stacked data. Especially, more faint stars are detected because of the higher sensitivity of the data. As the number of sources is important for our applied method of near-infrared extinction measurement we except small uncertainties introduced by the PSF fitting on the stacked data as they are not significant, which we could show here.

Appendix C: Reference color correction

For the estimation of the dust extinction of a molecular cloud we need to calculate the average color of the stars behind the cloud.



Fig. B.2. Comparison of K_S magnitudes between the stacked and 80 s data. The red line indicates the one-to-one correlation. The shown stars have an photometric uncertainty lower than 0.05 mag.

For good measurements we need to address two problems. First the cloud itself causes a strong shift of the star colors. This is exactly the effect we want to measure, but a direct measure of the color on the farther side of the cloud is impossible. Therefore, we assume the colors of stars in a nearby cloud-free region are the same as behind the cloud. Second, diffuse dust in the Milky Way causes a steady dust reddening with distance from the observer. Therefore, stars located in between the cloud and the observer will confuse the measurement of the background color and need to be removed. We address this problem by statistical subtraction of foreground stars in the JHK_S color-color-space. We first bin the stars in the J-H and H-K_S colors and scale the numbers with the size of the reference field, which leads to a 2-dimensional histogram shown in Fig. C.1. Then, we do the same for stars located towards the highest extinction regions of the cloud. These stars are either in front of the cloud and show almost now color excess or they are behind the cloud, in which case the show a strong color excess and can be ignored. Again, we scale the number of stars per bin with the area in which they where observed. We subtract the number of stars per bin of foreground histogram from the number of stars in the corresponding bin of the reference field histogram. The resulting histogram is shown in Fig. C.2 and represents the distribution of star colors behind the cloud. Some bins show negative number of stars, but neighbouring bins show still 'unreddend' stars, so they cancel in deriving the average J-H, and $H-K_S$ colors.

Appendix D: mid-infrared near-infrared correlation

For the combination of the near- and mid-infrared extinction maps we convolve the mid-infrared data (FWHM = 2.4") to the significantly lower resolution of the near-infrared data (FWHM = 48"). Then, we perform a pixel-to-pixel comparison between the two maps to investigate their correlation. Fig. D.1 shows only a poor correlation of the data and a large scatter. For $A_V^{NIR} \lesssim 10$ mag the mid-infrared extinction at most positions is underestimated by a factor of \sim 5, but at some positions the data is correlated. This can be explained by the spatial filtering of the mid-infrared mapping, which is not able to trace the diffuse cloud component. Therefore, the correlation arises only from the very inner parts of the filament. Additionally, at extinctions higher than $A_V^{NIR} \lesssim 5-10$ mag the near-infrared data

Table A.1. List of observations

Filter	RA	DEC	Exposure time	beamsize	date
	hh:mm:ss.ss	dd:mm:ss.ss	S	//	
near-in	frared VVV til	le d068			
J	16:40:50.52	-47:19:13.08	80	0.82	2010-03-27
J	16:40:50.52	-47:19:13.08	80	0.93	2010-05-09
Η	16:40:50.52	-47:19:13.08	80	0.84	2010-03-27
Η	16:40:50.52	-47:19:13.08	80	0.94	2010-05-09
K_S	16:40:50.52	-47:19:13.08	80	0.82	2010-03-27
K_S	16:40:50.52	-47:19:13.08	80	0.96	2010-05-09
K_S	16:40:50.52	-47:19:13.08	16	0.93	2010-03-06
K_S	16:40:50.52	-47:19:13.08	16	0.93	2010-06-26
\mathbf{K}_{s}	16:40:50.52	-47:19:13.08	16	0.74	2011-05-14
K_S	16:40:50.52	-47:19:13.08	16	0.82	2011-05-15
K_S	16:40:50.52	-47:19:13.08	16	0.74	2011-05-16
\mathbf{K}_{S}	16:40:50.52	-47:19:13.08	16	0.71	2011-05-16
\mathbf{K}_{S}	16:40:50.52	-47:19:13.08	16	0.90	2011-05-18
\mathbf{K}_{S}	16:40:50.52	-47:19:13.08	16	1.09	2011-08-31
K_S	16:40:50.52	-47:19:13.08	16	0.93	2011-09-01
K_S	16:40:50.52	-47:19:13.08	16	0.88	2011-09-05
K_S	16:40:50.52	-47:19:13.08	16	0.93	2011-09-17
\mathbf{K}_{S}	16:40:50.52	-47:19:13.08	16	0.81	2011-09-21
noonin	franced VVVV +:1	la 4060			
near-m	16.46.25 56	16 0009 46.12.07 22	80	0.70	2010 02 27
J	10.40.20.00 16.46.25.56	-40:13:07.32	80	0.79	2010-03-27
J TT	10:40:20.00	-40:13:07.32	80 80	0.90	2010-00-09
П	10:40:20.00	-40:13:07.32	80 80	0.81	2010-05-27
II V	10:40:20:00	-40:13:07.32	80	0.89	2010-03-09
	10.40.20.00 16.46.25.56	-40:13:07.32	80	0.03	2010-03-27
	10.40.20.00 16.46.25.56	-40:13:07.32	00 16	0.07	2010-03-09
	10.40.25.50 16.46.25.56	-40.13.07.32	10 16	1.02	2010-03-00
	10.40.20.00 16.46.25.56	-40:13:07.32	10 16	0.79	2010-06-16
	10.40.20.00 16.46.25.56	-40:13:07.32	10 16	0.97	2011-00-15 2011 05 14
	10.40.20.00 16.46.25.56	-40:13:07.32	10 16	0.71	2011-05-14 2011 05 15
r_S	10:40:20.00	-40:10:07.02	10	0.03	2011-00-10
	10:40:20.00 16:46:25 FC	-40:10:07.32	10 16	0.72	2011-00-10
K_S	10:40:20.00	-40:13:07.32	10 16	1.04	2011-08-09
\mathbf{n}_S	10:40:25.50	-40:13:07.32	10	1.01	2011-09-06

begins to underestimate the extinction, because of a lower number of background stars in the line-of-sight. A similar behavior of the correlation can be seen in the study of Kainulainen & Tan (2013).

Appendix E: ATLASGAL clumps

Here we show cut-outs from the combined near- and midinfrared extinction map of the 16 ATLASGAL GCSC sources contained in Nessie. In section 4.4 we describe how these parsecscale structures identified from ATLASGAL (white contours) break down into possibly star-forming substructures. Therefore, we show the positions of identified scale i = 2 structures with black crosses.

Appendix F: Properties of the identified structures

Here we show the properties of the identified small-scale structures, which are likely to become star formation sites. The shown properties are the results of the clumpfind-2D algorithm applied to the column density map of scale (i = 2).



Fig. C.1. JHK_s color-color histogram of the reference field before correction.



Fig. C.2. JHK $_s$ color-color histogram of the reference field after correction.



Fig. D.1. Pixel-to-pixel comparison of the mid- and nearinfrared extinction values restricted to the filament area (polygon in Fig. 5). The black line indicates the one-to-one correlation.



Fig. E.1. Half power ellipses of 16 GCSC ATLASGAL sources (black) overlaid on combined near- and mid-infrared extinction maps. The crosses mark the position of substructures detected on a scale-map (s = 2) by the clumpfind-2D algorithm within the ATLASGAL sources. The white lines indicate the contours of the ATLASGAL emission.

Table F.1. structures identified on the i = 2 scale-map

$\begin{array}{ c c c c c c c c c c c c c c c c c c c$	ID	l	b	$N(\mathrm{H}_2)_{\mathrm{peak}}$	$\mathrm{FWHM}_{\mathbf{x}}$	$\mathrm{FWHM}_{\mathrm{y}}$	R	$N({\rm H}_2)_{\rm tot}$	$N_{\rm pix}$
9 338.11 0.47 10.38 4.38 2.93 3.39 194.42 36 10 338.00 -0.45 9.66 2.87 2.67 2.65 127.39 22 13 38.70 -0.46 8.63 7.02 2.49 3.87 173.85 47 20 338.73 -0.47 8.38 3.43 3.19 3.14 106.113 44 24 338.09 -0.45 8.01 3.21 2.67 2.82 113.50 25 29 338.64 -0.46 7.55 4.21 2.61 3.09 108.18 306 338.71 -0.46 7.05 4.01 4.30 3.52 134.24 39 4338.64 -0.46 6.99 5.20 3.13 3.61 155.90 411 53 338.10 -0.46 6.89 5.20 3.13 3.61 152.90		[°]	[°]	$\left[\frac{1}{\mathrm{cm}^2}\right]$	pix	pix	pix	$\left[\frac{1}{\mathrm{cm}^2}\right]$	
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	9	338.11	-0.47	10.38	4.38	2.93	3.39	194.42	36
$ \begin{array}{ccccccccccccccccccccccccccccccccccc$	10	338.08	-0.45	9.78	4.96	6.44	3.95	224.46	49
	11	338.09	-0.45	9.66	2.87	2.67	2.65	127.39	22
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	10	338.80	-0.47	8.71	10.33	3.52	4.62	232.36	67
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	17	338.70 338 73	-0.40	8.03	1.02	2.49 3.10	3.87 3.14	173.80	47 21
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	20 24	338.00	-0.47	0.00 8.13	3.43 4 51	5.19 6.27	3.14	120.11 160.13	31 44
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	24 26	338.09	-0.45	8.01	3.21	2.67	2.82	11350	25
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	29	338.65	-0.46	7.77	5.00	3.77	3.19	118.37	32
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	33	338.71	-0.46	7.55	4.21	2.61	3.09	108.18	30
$ \begin{array}{ccccccccccccccccccccccccccccccccccc$	36	338.11	-0.46	7.38	2.27	4.12	2.82	85.24	25
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	43	338.64	-0.46	7.14	2.93	2.09	2.33	82.64	17
46 338.11 -0.47 6.95 4.01 4.30 3.52 134.24 39 48 338.08 -0.46 6.89 5.20 3.13 3.61 125.90 41 50 338.69 -0.46 6.87 2.74 5.54 3.61 125.90 41 51 338.34 -0.50 6.84 3.37 2.89 2.88 89.20 26 53 338.09 -0.43 6.79 1.80 4.40 2.65 83.35 22 54 338.10 -0.45 6.73 3.88 1.90 2.52 80.81 200 56 338.08 -0.44 6.66 3.01 2.54 2.71 76.39 239 63.31 18 60 338.55 -0.42 6.45 6.23 4.03 3.61 133.05 41 67 338.87 -0.48 6.15 2.36 3.23 5.62 6.61 18.33 12.71 195 338.12 <td>44</td> <td>338.73</td> <td>-0.46</td> <td>7.09</td> <td>3.01</td> <td>3.62</td> <td>2.93</td> <td>102.15</td> <td>27</td>	44	338.73	-0.46	7.09	3.01	3.62	2.93	102.15	27
48 338.08 -0.45 6.92 3.21 2.83 2.52 72.30 20 49 338.69 -0.46 6.89 5.20 3.13 3.61 125.90 41 50 338.87 -0.48 6.87 2.74 5.54 3.61 153.59 41 51 338.09 -0.43 6.73 3.88 1.90 2.52 80.81 20 56 338.08 -0.44 6.69 5.41 2.96 3.24 111.74 33 58 338.31 -0.51 6.66 3.01 2.54 2.71 76.39 23 59 338.08 -0.44 6.27 1.92 2.17 1.95 43.95 12 71 338.87 -0.48 6.15 2.36 3.13 2.33 50.64 33.85 7.04 6.01 5.25 2.75 3.19 95.52 32 32 71 338.87 -0.48 5.84 2.25 3.41 2.46 72.41 19 37 338.39 -0.40 5.79 <t< td=""><td>46</td><td>338.11</td><td>-0.47</td><td>6.95</td><td>4.01</td><td>4.30</td><td>3.52</td><td>134.24</td><td>39</td></t<>	46	338.11	-0.47	6.95	4.01	4.30	3.52	134.24	39
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	48	338.08	-0.45	6.92	3.21	2.83	2.52	72.30	20
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	49	338.69	-0.46	6.89	5.20	3.13	3.61	125.90	41
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	50 51	338.87 220 24	-0.48	0.87 6.84	2.74	5.54	3.61	153.59	41 26
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	53	338.04	-0.30	6 79	1.80	2.89	2.00 2.65	83.05	$\frac{20}{22}$
56 338.08 -0.44 6.69 5.41 2.90 3.24 0111.74 33 58 338.31 -0.51 6.66 3.01 2.54 2.71 76.39 23 59 338.11 -0.47 6.61 1.91 3.39 2.39 63.31 18 60 338.55 -0.42 6.45 6.23 4.03 3.61 133.05 41 67 338.08 -0.43 6.04 2.23 2.42 2.26 52.61 16 78 338.87 -0.48 6.15 2.36 3.23 2.46 67.43 19 77 338.08 -0.43 6.00 2.70 1.72 2.11 45.62 14 83 38.87 -0.46 5.92 2.34 3.13 2.33 50.45 17 85 338.87 -0.48 5.74 2.83 5.21 3.09 78.61 30 80 338.02 -0.44 5.74 2.60 3.44 3.44 3.64 3.25 67.07 22 3	54	338.09	-0.45	6 73	3.88	1.40	$\frac{2.00}{2.52}$	80.81	$\frac{22}{20}$
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	56	338.08	-0.44	6.69	5.41	2.96	3.24	111.74	33
59338.11 -0.47 6.61 1.91 3.39 2.39 63.31 18 60338.55 -0.42 6.45 6.23 4.03 3.61 133.05 41 67338.09 -0.44 6.27 1.92 2.17 1.95 43.95 12 71338.09 -0.44 6.15 2.36 3.23 2.46 67.43 19 77338.08 -0.43 6.00 2.23 2.42 2.26 52.61 16 78338.78 -0.47 6.01 5.25 2.75 3.19 95.52 32 79338.08 -0.43 6.00 2.70 1.72 2.11 45.62 14 83338.78 -0.48 5.84 2.25 3.44 2.46 72.41 19 87 338.39 -0.40 5.79 2.66 3.43 2.65 65.60 22 88 338.69 -0.44 5.74 2.60 2.47 2.33 53.72 17 90 338.32 -0.41 5.73 4.68 3.34 3.19 105.23 32 95 338.32 -0.51 5.67 8.75 2.10 3.57 11.7 40 96 338.19 -0.48 5.64 2.84 3.06 2.65 67.07 22 97 338.24 -0.44 5.63 2.63 4.44 3.04 88.73 29 100 338.10 -0.45 5.59	$\overline{58}$	338.31	-0.51	6.66	3.01	2.54	2.71	76.39	23
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	59	338.11	-0.47	6.61	1.91	3.39	2.39	63.31	18
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	60	338.55	-0.42	6.45	6.23	4.03	3.61	133.05	41
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	67	338.09	-0.44	6.27	1.92	2.17	1.95	43.95	12
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	71	338.87	-0.48	6.15	2.36	3.23	2.46	67.43	19
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	77	338.08	-0.43	6.04	2.23	2.42	2.26	52.61	16
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	78	338.87	-0.47	6.01	5.25	2.75	3.19	95.52	32
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	79	338.08	-0.43	6.00 5.00	2.70	1.72	2.11	45.62	14
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	00 85	338.87	-0.40	0.92 5.84	2.34	0.10 2.44	2.33	50.45 72.41	10
338.612 0.44 5.74 2.63 5.15 2.03 5.16 3.09 78.61 30 89 338.09 -0.45 5.74 2.60 2.47 2.33 53.72 17 90 338.32 -0.41 5.73 4.68 3.34 3.19 105.23 32 95 338.32 -0.51 5.67 8.75 2.10 3.57 115.17 40 96 338.19 -0.48 5.64 2.84 3.06 2.65 67.07 22 97 338.24 -0.44 5.63 2.63 4.44 3.04 88.73 29 100 338.11 -0.49 5.62 3.73 3.68 3.14 85.00 31 102 338.11 -0.45 5.59 6.57 3.81 3.39 107.10 36 106 338.46 -0.43 5.50 5.86 2.02 3.04 85.06 29 108 338.10 -0.45 5.49 3.55 2.47 2.39 58.37 18 110 338.85 -0.47 5.47 2.49 2.68 2.246 59.34 19 111 338.09 -0.44 5.45 3.26 2.60 2.65 70.43 22 116 338.60 -0.44 5.44 2.69 3.57 2.71 75.55 23 117 338.60 -0.44 5.44 2.69 3.57 2.71 75.55 23 </td <td>87</td> <td>338.39</td> <td>-0.48</td> <td>5.84 5.79</td> <td>2.25 2.66</td> <td>3.44 3.43</td> <td>2.40 2.65</td> <td>65.60</td> <td>$\frac{19}{22}$</td>	87	338.39	-0.48	5.84 5.79	2.25 2.66	3.44 3.43	2.40 2.65	65.60	$\frac{19}{22}$
89338.09 -0.45 5.74 2.60 2.47 2.33 53.72 17 90338.32 -0.41 5.73 4.68 3.34 3.19 105.23 32 95338.32 -0.51 5.67 8.75 2.10 3.57 115.17 40 96338.19 -0.48 5.64 2.84 3.06 2.65 67.07 22 97 338.24 -0.44 5.63 2.63 4.44 3.04 88.73 29 100 338.13 -0.49 5.62 3.73 3.68 3.14 85.00 31 102 338.10 -0.45 5.59 6.57 3.81 3.39 107.10 36 106 338.46 -0.43 5.50 5.66 2.02 3.04 85.06 29 108 338.10 -0.45 5.49 3.55 2.47 2.39 58.37 18 110 338.85 -0.47 5.47 2.49 2.68 2.46 59.34 19 111 338.30 -0.44 5.44 2.69 3.57 2.71 75.55 23 117 338.60 -0.44 5.44 2.69 3.57 2.71 75.55 23 117 338.64 -0.46 5.43 3.59 1.88 2.39 58.47 18 120 338.67 -0.45 5.40 4.36 2.46 2.71 63.15 23 118 338.64 -0.46 <	88	338.62	-0.44	5.74	3.83	5.21	$\frac{2.00}{3.09}$	78.61	30
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	89	338.09	-0.45	5.74	2.60	2.47	2.33	53.72	17^{-10}
95 338.32 -0.51 5.67 8.75 2.10 3.57 115.17 40 96 338.19 -0.48 5.64 2.84 3.06 2.65 67.07 22 97 338.24 -0.44 5.63 2.63 4.44 3.04 88.73 29 100 338.13 -0.49 5.62 3.73 3.68 3.14 85.00 31 102 338.11 -0.45 5.60 2.71 2.89 2.65 66.99 22 103 338.10 -0.45 5.59 6.57 3.81 3.39 107.10 36 106 338.46 -0.43 5.50 5.86 2.02 3.04 85.06 29 108 338.10 -0.45 5.49 3.55 2.47 2.39 58.37 18 110 338.50 -0.47 5.47 2.49 2.68 2.46 59.34 19 111 338.33 -0.51 5.46 2.03 2.91 2.33 50.69 17 113 338.09 -0.44 5.45 3.26 2.60 2.65 70.43 22 116 338.40 -0.46 5.44 2.69 3.57 2.71 75.55 23 117 338.60 -0.44 5.44 2.69 3.57 2.71 75.55 23 118 338.64 -0.46 5.33 2.96 2.73 2.52 64.19 20 122 338.84 <td< td=""><td>90</td><td>338.32</td><td>-0.41</td><td>5.73</td><td>4.68</td><td>3.34</td><td>3.19</td><td>105.23</td><td>32</td></td<>	90	338.32	-0.41	5.73	4.68	3.34	3.19	105.23	32
96 338.19 -0.48 5.64 2.84 3.06 2.65 67.07 22 97 338.24 -0.44 5.63 2.63 4.44 3.04 88.73 29 100 338.13 -0.49 5.62 3.73 3.68 3.14 85.00 31 102 338.11 -0.45 5.60 2.71 2.89 2.65 66.99 22 103 338.10 -0.45 5.59 6.57 3.81 3.39 107.10 36 106 338.46 -0.43 5.50 5.86 2.02 3.04 85.06 29 108 338.10 -0.45 5.49 3.55 2.47 2.39 58.37 18 110 338.85 -0.47 5.47 2.49 2.68 2.46 59.34 19 111 338.33 -0.51 5.46 2.03 2.91 2.33 50.69 17 113 38.09 -0.44 5.45 3.26 2.60 2.65 70.43 22 116 338.20 -0.46 5.43 3.59 1.88 2.39 58.47 18 120 338.64 -0.46 5.43 3.59 1.88 2.39 58.47 18 120 338.67 -0.45 5.32 2.92 3.32 2.52 64.19 20 122 338.84 -0.46 5.33 6.25 3.01 3.24 89.47 33	95	338.32	-0.51	5.67	8.75	2.10	3.57	115.17	40
97 338.24 -0.44 5.63 2.63 4.44 3.04 88.73 29 100 338.13 -0.49 5.62 3.73 3.68 3.14 85.00 31 102 338.11 -0.45 5.60 2.71 2.89 2.65 66.99 22 103 338.10 -0.45 5.59 6.57 3.81 3.39 107.10 36 106 338.46 -0.43 5.50 5.86 2.02 3.04 85.06 29 108 338.10 -0.45 5.49 3.55 2.47 2.39 58.37 18 110 338.85 -0.47 5.47 2.49 2.68 2.46 59.34 19 111 338.33 -0.51 5.46 2.03 2.91 2.33 50.69 17 113 338.09 -0.44 5.45 3.26 2.60 2.65 70.43 22 116 338.64 -0.46 5.44 2.69 3.57 2.71 75.55 23 117 338.06 -0.44 5.45 3.29 5.47 18 2.39 58.47 18 120 338.67 -0.45 5.40 4.36 2.46 2.71 63.15 23 121 338.08 -0.44 5.35 2.96 2.73 2.52 64.19 20 122 338.34 -0.40 5.33 6.25 3.01 3.24 89.47 33 125 3	96	338.19	-0.48	5.64	2.84	3.06	2.65	67.07	22
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	97	338.24	-0.44	5.63	2.63	4.44	3.04	88.73	29
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	100	338.13	-0.49	5.62	3.73	3.68	3.14	85.00	31
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	102	338.11	-0.45	5.60	2.71	2.89	2.65	66.99	22
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	103	338.10	-0.45	5.59 5.50	0.57 5.86	3.81	3.39	107.10	30 20
100 338.10 -0.45 5.45 3.53 2.41 2.35 53.51 110 110 338.85 -0.47 5.47 2.49 2.68 2.46 59.34 19 111 338.33 -0.51 5.46 2.03 2.91 2.33 50.69 17 113 338.09 -0.44 5.45 3.26 2.60 2.65 70.43 22 116 338.20 -0.46 5.44 2.69 3.57 2.71 75.55 23 117 338.60 -0.44 5.44 1.83 5.45 2.82 79.80 25 118 338.64 -0.46 5.43 3.59 1.88 2.39 58.47 18 120 338.67 -0.45 5.40 4.36 2.46 2.71 63.15 23 121 338.08 -0.44 5.35 2.96 2.73 2.52 64.19 20 122 338.4 -0.40 5.33 6.25 3.01 3.24 89.47 33 125 338.82 -0.45 5.32 2.92 3.32 2.52 59.44 20 126 338.81 -0.48 5.31 3.19 2.31 2.59 60.68 21 129 338.60 -0.44 5.27 3.64 2.43 2.71 75.71 23 134 338.60 -0.44 5.22 3.72 2.35 2.82 70.46 25 <td< td=""><td>100</td><td>338.40 338.10</td><td>-0.45</td><td>5.30 5.40</td><td>0.00 3.55</td><td>2.02 2.47</td><td>5.04 2.30</td><td>60.00 58.37</td><td>29 18</td></td<>	100	338.40 338.10	-0.45	5.30 5.40	0.00 3.55	2.02 2.47	5.04 2.30	60.00 58.37	29 18
110338.33 -0.41 5.46 2.16 2.16 2.16 5.061 17 113338.33 -0.44 5.45 3.26 2.60 2.65 70.43 22 116338.20 -0.46 5.44 2.69 3.57 2.71 75.55 23 117338.60 -0.44 5.44 1.83 5.45 2.82 79.80 25 118338.64 -0.46 5.43 3.59 1.88 2.39 58.47 18 120338.67 -0.45 5.40 4.36 2.46 2.71 63.15 23 121338.08 -0.44 5.35 2.96 2.73 2.52 64.19 20 122338.44 -0.40 5.33 6.25 3.01 3.24 89.47 33 125338.82 -0.45 5.32 2.92 3.32 2.52 59.44 20 126338.81 -0.48 5.31 3.19 2.31 2.59 60.68 21 129338.69 -0.46 5.28 3.21 2.11 2.39 52.69 18 130338.60 -0.44 5.27 3.64 2.43 2.71 75.71 23 134338.82 -0.47 5.23 3.11 3.57 2.65 61.68 22 135338.94 -0.42 5.22 2.34 4.97 2.99 83.16 28 136 338.50 -0.42 5.21 <td>110</td> <td>338.85</td> <td>-0.43</td> <td>5.49 5.47</td> <td>2.49</td> <td>2.47</td> <td>$\frac{2.39}{2.46}$</td> <td>59.34</td> <td>10</td>	110	338.85	-0.43	5.49 5.47	2.49	2.47	$\frac{2.39}{2.46}$	59.34	10
113338.09 -0.44 5.453.262.602.65 70.43 22116338.20 -0.46 5.44 2.69 3.57 2.71 75.55 23 117338.60 -0.44 5.44 1.83 5.45 2.82 79.80 25 118338.64 -0.46 5.43 3.59 1.88 2.39 58.47 18 120338.67 -0.45 5.40 4.36 2.46 2.71 63.15 23 121338.08 -0.44 5.35 2.96 2.73 2.52 64.19 20 122338.34 -0.40 5.33 6.25 3.01 3.24 89.47 33 125338.82 -0.45 5.32 2.92 3.32 2.52 59.44 20 126338.81 -0.48 5.31 3.19 2.31 2.59 60.68 21 129338.69 -0.46 5.28 3.21 2.11 2.39 52.69 18 130338.60 -0.44 5.27 3.64 2.43 2.71 75.71 23 134 338.82 -0.47 5.23 3.11 3.57 2.65 61.68 22 135 338.94 -0.42 5.22 2.34 4.97 2.99 83.16 28 136 338.27 -0.43 5.21 3.78 3.16 2.71 69.47 23 139 338.33 -0.44 5.19	111	338.33	-0.51	5.46	2.03	2.91	2.33	50.69	17
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	113	338.09	-0.44	5.45	3.26	2.60	2.65	70.43	$\frac{1}{22}$
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	116	338.20	-0.46	5.44	2.69	3.57	2.71	75.55	23
118 338.64 -0.46 5.43 3.59 1.88 2.39 58.47 18 120 338.67 -0.45 5.40 4.36 2.46 2.71 63.15 23 121 338.08 -0.44 5.35 2.96 2.73 2.52 64.19 20 122 338.34 -0.40 5.33 6.25 3.01 3.24 89.47 33 125 338.82 -0.45 5.32 2.92 3.32 2.52 59.44 20 126 338.81 -0.48 5.31 3.19 2.31 2.59 60.68 21 129 338.69 -0.46 5.28 3.21 2.11 2.39 52.69 18 130 338.60 -0.44 5.27 3.64 2.43 2.71 75.71 23 134 338.82 -0.47 5.23 3.11 3.57 2.65 61.68 22 135 338.94 -0.42 5.22 2.34 4.97 2.99 83.16 28 136 338.27 -0.43 5.22 3.72 2.35 2.82 70.46 25 138 338.50 -0.42 5.21 3.78 3.16 2.71 69.47 23 139 338.33 -0.44 5.16 3.26 4.76 3.09 95.21 30 147 338.60 -0.44 5.16 3.26 4.76 3.09 95.21 30 <td< td=""><td>117</td><td>338.60</td><td>-0.44</td><td>5.44</td><td>1.83</td><td>5.45</td><td>2.82</td><td>79.80</td><td>25</td></td<>	117	338.60	-0.44	5.44	1.83	5.45	2.82	79.80	25
120 338.67 -0.45 5.40 4.36 2.46 2.71 63.15 23 121 338.08 -0.44 5.35 2.96 2.73 2.52 64.19 20 122 338.34 -0.40 5.33 6.25 3.01 3.24 89.47 33 125 338.82 -0.45 5.32 2.92 3.32 2.52 59.44 20 126 338.81 -0.48 5.31 3.19 2.31 2.59 60.68 21 129 338.69 -0.46 5.28 3.21 2.11 2.39 52.69 18 130 338.60 -0.44 5.27 3.64 2.43 2.71 75.71 23 134 338.82 -0.47 5.23 3.11 3.57 2.65 61.68 22 135 338.94 -0.42 5.22 2.34 4.97 2.99 83.16 28 136 338.27 -0.43 5.22 3.72 2.35 2.82 70.46 25 138 338.50 -0.42 5.21 3.78 3.16 2.71 69.47 23 139 338.33 -0.46 5.19 3.75 6.70 2.99 73.09 28 145 338.61 -0.44 5.16 3.26 4.76 3.09 95.21 30 147 338.60 -0.44 5.12 3.34 2.18 2.46 65.92 19 <td< td=""><td>118</td><td>338.64</td><td>-0.46</td><td>5.43</td><td>3.59</td><td>1.88</td><td>2.39</td><td>58.47</td><td>18</td></td<>	118	338.64	-0.46	5.43	3.59	1.88	2.39	58.47	18
121 338.08 -0.44 5.35 2.96 2.73 2.52 64.19 20 122 338.34 -0.40 5.33 6.25 3.01 3.24 89.47 33 125 338.82 -0.45 5.32 2.92 3.32 2.52 59.44 20 126 338.81 -0.48 5.31 3.19 2.31 2.59 60.68 21 129 338.69 -0.46 5.28 3.21 2.11 2.39 52.69 18 130 338.60 -0.44 5.27 3.64 2.43 2.71 75.71 23 134 338.82 -0.47 5.23 3.11 3.57 2.65 61.68 22 135 338.94 -0.42 5.22 2.34 4.97 2.99 83.16 28 136 338.27 -0.43 5.22 3.72 2.35 2.82 70.46 25 138 338.50 -0.42 5.21 3.78 3.16 2.71 69.47 23 139 338.33 -0.46 5.19 3.75 6.70 2.99 73.09 28 141 338.09 -0.46 5.19 3.75 6.70 2.99 73.09 28 145 338.61 -0.44 5.16 3.26 4.76 3.09 95.21 30 147 338.60 -0.44 5.12 3.34 2.18 2.46 65.92 19 <td< td=""><td>120</td><td>338.67</td><td>-0.45</td><td>5.40</td><td>4.36</td><td>2.46</td><td>2.71</td><td>63.15</td><td>23</td></td<>	120	338.67	-0.45	5.40	4.36	2.46	2.71	63.15	23
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	121	338.08	-0.44	5.35	2.96	2.73	2.52	64.19	20
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	122	338.34	-0.40	5.33	6.25	3.01	3.24	89.47	33
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	120	338.82 220 01	-0.45	0.32 5.21	2.92	3.32 9.21	2.52	59.44 60.68	20
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	120	338.60	-0.46	5.01	3.19	2.31 2.11	2.09 2.30	52.60	21 18
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	$120 \\ 130$	338.60	-0.44	$5.20 \\ 5.27$	3.64	2.43	$\frac{2.55}{2.71}$	75.71	$\frac{10}{23}$
$ \begin{array}{c ccccccccccccccccccccccccccccccccccc$	134	338.82	-0.47	5.23	3.11	3.57	2.65	61.68	$\frac{-3}{22}$
$ \begin{array}{c ccccccccccccccccccccccccccccccccccc$	135	338.94	-0.42	5.22	2.34	4.97	2.99	83.16	28^{-}
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	136	338.27	-0.43	5.22	3.72	2.35	2.82	70.46	25
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	138	338.50	-0.42	5.21	3.78	3.16	2.71	69.47	23
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	139	338.33	-0.41	5.21	3.61	1.76	2.33	55.73	17
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	141	338.09	-0.46	5.19	3.75	6.70	2.99	73.09	28
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	145	338.61	-0.44	5.16	3.26	4.76	3.09	95.21 65.00	30
	147 148	338 78	-0.44 -0.46	$\frac{0.12}{5.19}$	3.34 4.06	$\frac{2.18}{4.37}$	$2.40 \\ 3.43$	00.92 94.78	$\frac{19}{37}$

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		Tabl				s page		
ID	l	b	$N({\rm H}_2)_{\rm peak}$	$\mathrm{FWHM}_{\mathbf{x}}$	$FWHM_y$	R	$N({\rm H_2})_{\rm tot}$	$N_{\rm pix}$
	[°]	[°]	$\left[\frac{1}{\mathrm{cm}^2}\right]$	$_{\rm pix}$	pix	pix	$\left[\frac{1}{cm^2}\right]$	
149	338.87	-0.48	5.12	3.81	1.96	2.39	48.98	18
154	338.85	-0.47	5.03	2.84	2.09	2.33	47.11	17
157	338.48	-0.43	4.95	4.17	2.94	3.14	87.01	31
159	338.29	-0.43	4.93	3.24	3.61	3.09	81.85	30
160	338.30	-0.52	4 92	3.00	4 62	3 29	98.33	34
161	338.62	-0.44	4 91	3.04	5.80	3.09	79.39	30
164	338.52	-0.43	4 89	2.45	2.04	2.11	38 29	14
168	338.62	-0.43	4.85	2.40 2.20	2.04 2.72	2.11 2.33	17 18	17
170	338.60	0.44	4.83	2.20	2.12	2.00	56.86	10
170	338 30	0.52	4.00	J.11 4 16	2.49	2.40	72.20	25
170	220.30	-0.52	4.02	4.10	2.31	2.02	12.29	20
170	000.09 000 ED	-0.44	4.70	4.77	2.31	2.00	78 70	20
101	000.0U	-0.42	4.11	5.09 7.50	4.07	2.00	10.19	20
101	000.04 000.75	-0.45	4.70	1.52	5.05	0.07	100.69	40
102	000.70 000.70	-0.40	4.70	0.73	0.10	0.74 0.22	51.00	44
183	338.34	-0.51	4.75	2.98	2.28	2.33	51.20	17
184	338.76	-0.48	4.74	2.20	4.51	2.71	57.15	23
185	338.61	-0.44	4.73	6.08	3.39	3.57	104.06	40
189	338.70	-0.48	4.71	4.08	3.40	3.19	81.20	32
190	338.11	-0.46	4.71	3.54	1.95	2.33	42.64	17
191	338.61	-0.44	4.70	3.15	2.90	2.71	62.21	23
195	338.30	-0.52	4.66	2.44	6.96	3.14	80.88	31
196	338.09	-0.45	4.66	6.14	2.50	3.14	80.21	31
197	338.10	-0.46	4.66	2.83	2.30	2.26	45.35	16
199	338.73	-0.47	4.62	3.08	6.62	2.76	57.63	24
201	338.57	-0.44	4.61	3.60	2.26	2.46	50.63	19
202	338.65	-0.45	4.61	2.44	1.79	1.95	33.23	12
204	338.17	-0.47	4.59	2.19	2.51	2.19	39.57	15
208	338.18	-0.46	4.56	2.20	1.96	1.95	35.85	12
211	338.31	-0.52	4.56	5.37	2.54	2.71	57.95	23
213	338.10	-0.46	4.55	2.73	3.91	2.76	66.28	24
217	339.04	-0.39	4.54	4.51	2.87	3.14	85.78	31
218	338.33	-0.41	4.52	3.38	1.75	2.11	41.53	14
219	338.93	-0.49	4 52	2.79	4 27	3 14	84 45	31
210	338.31	-0.51	4 50	1.74	2.43	1.95	32.52	12
$220 \\ 227$	338.86	-0.47	1.00	4 71	3 95	3 /8	86.50	38
221	338.08	0.44	4.46	3.36	1.70	0.40	43.44	17
220	228 22	-0.44	4.40	2.30	1.79	$2.00 \\ 9.11$	28.08	14
229 021	000.00 000 01	-0.41	4.40	2.90	1.04	2.11	30.90 20 E7	14
201	220.01	-0.40	4.40	2.20	2.34	2.05	30.37	10
230	000.0U	-0.52	4.45	3.70	0.27 0.01	2.02	02.30	20
239	338.20	-0.40	4.42	4.80	3.31	3.19	76.03	32
240	338.93	-0.43	4.42	2.64	2.19	2.26	42.37	10
242	338.77	-0.46	4.40	3.02	2.80	2.46	51.70	19
246	338.19	-0.46	4.37	3.46	2.83	2.71	64.63	23
248	338.24	-0.44	4.36	1.55	3.30	2.03	37.13	13
249	338.29	-0.43	4.36	2.86	2.19	2.26	41.10	16
252	338.31	-0.43	4.35	1.68	3.33	2.19	38.06	15
253	338.45	-0.42	4.35	7.54	2.76	3.29	78.07	34
256	338.20	-0.48	4.32	2.47	4.63	2.88	67.15	26
257	338.62	-0.44	4.31	2.13	2.44	1.95	34.91	12
259	338.09	-0.42	4.31	3.15	2.10	2.39	46.75	18
267	338.77	-0.46	4.27	3.16	2.20	2.26	45.37	16
269	338.21	-0.48	4.27	2.08	1.87	1.87	30.97	11
271	338.18	-0.48	4.26	4.38	3.80	2.93	69.68	27
272	338.33	-0.41	4.26	5.35	4.27	3.39	99.61	36
273	338.18	-0.46	4.25	3.00	2.82	2.52	56.71	20
276	338.87	-0.49	4.23	3.27	4.13	2.88	62.19	26
277	338.87	-0.48	4.22	4.13	2.40	2.46	48.62	19
279	338.63	-0.45	4.21	2.34	3.10	2.33	42.57	17^{-10}
282	338 12	-0.47	4 20	$\frac{2.01}{3.42}$	3 57	3.09	72.84	30
285	338.00	-0.43	4.20	4 70	3 30	2.03	66 23	27
200	228 16	0.40	4.20	1.00	0.00 0.47	2.90 1 05	21 50	41 19
201	000.40 220 47	-0.40	4.10	1.09	2.41 0 50	1.90 9.92	01.00 46.07	14 16
200	000.47 220 00	-0.43	4.10	ა.ა∠ ე დე	2.08	2.20 0.10	40.97	10
289	338.90	-0.43	4.18	2.82	2.09	2.19	39.80	10
290	338.51 220.05	-0.42	4.17	5.28	2.82	3.09 0.00	09.12	30 95
291	338.85	-0.47	4.17	2.92	3.41	2.82	00.00	25
293	338.87	-0.49	4.15	2.24	2.26	2.03	32.41	13
						Contir	nued on nex	t page

Table F.1 – Continued from previous page

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		Tabl	e r.r = Con	tinued iro	m previou	s page		
ID	l	b	$N({\rm H}_2)_{\rm peak}$	$FWHM_{x}$	$FWHM_{v}$	R	$N(\mathrm{H}_2)_{\mathrm{tot}}$	$N_{\rm pix}$
	[0]	[0]	$\begin{bmatrix} 1 \\ \end{bmatrix}$	niv	niv	niv	$\begin{bmatrix} 1 \\ \end{bmatrix}$	pix
	[]		$\left[\frac{1}{cm^2}\right]$	pix	pix		$\left[\frac{1}{\mathrm{cm}^2}\right]$	
296	338.33	-0.51	4.13	2.07	2.52	2.11	38.05	14
298	338.26	-0.38	4.12	2.75	2.28	2.26	39.48	16
300	338.08	-0.45	4.12	2.29	5.21	2.19	34.42	15
303	338 10	0.47	4 11	7 73	3 5 8	2.03	60.83	27
303	000.19	-0.47	4.11	1.13	0.00	2.90	00.85	21 14
306	338.71	-0.47	4.09	2.51	2.26	2.11	38.51	14
307	338.09	-0.45	4.09	3.07	2.07	2.33	41.94	17
308	338.75	-0.47	4.09	2.99	3.47	2.88	67.70	26
310	338.81	_0.49	4.09	2.98	3.04	2.76	60.01	24
210	220 50	-0.43	4.00	2.50	0.04	2.10	74.00	24
312	338.30	-0.44	4.08	0.34	2.29	3.04	74.99	29
315	338.34	-0.41	4.08	3.59	1.91	2.11	34.87	14
316	338.16	-0.48	4.08	1.99	3.17	2.26	41.14	16
317	338 93	-0.49	4.08	3 41	2.55	2.46	$45 \ 47$	19
310	338 78	0.46	4.07	2 01	4.07	2.76	58 32	24
019	000.10	-0.40	4.07	2.91	4.07	2.10	45.95	24
326	338.32	-0.42	4.05	3.33	3.87	2.59	45.35	21
329	338.46	-0.43	4.05	3.33	1.96	2.39	43.65	18
330	338.93	-0.49	4.04	6.58	4.80	3.61	111.99	41
333	338 65	-0.45	4.03	4.66	2.04	2.65	54.46	22
227	220 06	0.10	4.00	1.00	2.01	2.00	55 51	21
557	338.80	-0.48	4.02	4.20	2.04	2.09	00.01	21
338	338.65	-0.45	4.02	3.83	4.01	2.88	60.45	26
339	338.58	-0.44	4.02	1.73	2.48	1.95	29.73	12
340	338.58	-0.43	4.02	1.65	5.11	2.59	52.32	21
342	338 24	-0.45	4.01	1 79	2 55	1 95	30.15	12
242	220.24	-0.40	4.01	1.15	2.00	0.00	21.01	12
343	338.87	-0.49	4.01	1.77	2.00	2.03	31.81	13
344	338.54	-0.43	4.00	2.98	2.09	2.33	38.68	17
345	339.09	-0.41	3.99	3.68	2.49	2.59	53.84	21
346	338.71	-0.48	3.99	2.20	2.50	2.03	33.55	13
350	228 21	0.42	3.07	2.20	2.00	$\frac{2.00}{2.10}$	30.75	15
350	000.01	-0.42	0.91	0.00	2.09	2.19	19.10	10
353	338.72	-0.47	3.96	3.01	2.18	2.26	42.21	16
358	338.32	-0.41	3.94	3.42	3.18	2.46	47.99	19
360	338.13	-0.49	3.93	5.36	2.08	2.65	51.31	22
361	338.00	-0.40	3 03	5.93	3 47	3 30	90.02	36
262	220.22	-0.40	2.00	1.00	2.10	0.00	25.02	15
303	330.01	-0.49	5.92	1.99	3.20	2.19	55.27	15
367	338.92	-0.50	3.91	1.74	3.51	2.33	43.16	17
368	338.62	-0.45	3.91	3.24	2.77	2.39	43.53	18
373	338.88	-0.49	3.88	3.06	2.38	2.33	39.56	17
374	338.05	0.42	3 88	3.84	4 10	2.00	67.02	26
974	220.50	-0.42	0.00	0.04	9.10	2.00	44.64	10
370	338.57	-0.44	3.87	2.80	2.99	2.40	44.04	19
379	338.13	-0.47	3.86	2.24	2.57	2.19	35.89	15
380	338.85	-0.47	3.85	3.28	2.52	2.39	42.99	18
386	338 86	-0.48	3.82	3 18	2.00	2.19	$36\ 43$	15
387	338 39	0.41	3.82	2 75	2.00	2 20	46 55	18
307	000.02	-0.41	0.02	2.10	2.12	2.39	40.00	10
389	338.11	-0.46	3.81	3.14	1.97	3.48	90.03	38
391	339.02	-0.40	3.81	5.29	2.79	3.14	79.56	31
393	338.68	-0.46	3.81	5.74	5.04	3.09	61.83	30
394	338 10	-0.46	3.81	2.87	3.85	2.93	70.69	27
307	338 85	0.45	3 80	2.01	3 10	2.00	46 74	18
397	336.65	-0.45	3.80	0.04	3.10	2.09	40.74	10
400	338.87	-0.47	3.79	2.38	2.87	2.33	40.43	17
402	338.78	-0.46	3.79	1.79	2.52	1.87	27.91	11
407	338.55	-0.42	3.77	3.22	2.70	2.39	44.22	18
408	338 78	-0.46	3 77	2.81	2.28	2.26	30 11	16
400	228 40	-0.40	2.76	2.01	2.20	2.20	24.17	10
409	338.40	-0.40	3.70	2.31	2.37	2.03	34.17	13
410	338.50	-0.42	3.76	3.23	2.94	2.39	44.86	18
416	338.25	-0.44	3.74	4.34	3.55	2.65	51.41	22
418	338.94	-0.49	3.74	2.97	2.28	2.33	41.33	17
410	228 22	0.40	3 74	2.50	1.03	2.03	21.01	12
419	000.00	-0.40	0.74	2.02	1.30	2.00	JI.JI 47 40	10
424	338.11	-0.46	3.13	2.23	3.18	2.52	47.40	20
427	338.81	-0.46	3.73	2.21	2.09	1.95	29.11	12
430	338.41	-0.41	3.70	4.55	2.54	2.59	50.33	21
432	338 55	-0 /3	3 70	1.62	3 16	2 11	30.81	14
495	990 A4	0.40	9.40	2.02	0.10	2.11 0 E0	16 00	
435	ə ə ð.94	-0.42	3.09	3.12	2.50	2.52	40.00	20
436	338.30	-0.48	3.69	8.93	3.74	3.24	67.69	33
437	338.42	-0.42	3.69	3.96	2.96	2.33	37.20	17
441	$338\ 47$	-0.43	3.68	4.43	6.44	3.19	70.14	32
119	220 04	0.40	967	9.90	0.11	9 10	16 10	10
443	JJ0.94	-0.49	0.07	ა.ა∠ ი იი	2.28	∠.40 0.11	40.19	19
445	338.34	-0.51	3.66	3.00	1.76	2.11	33.98	14
_447	338.92	-0.49	3.66	2.54	2.26	2.19	35.85	15

Table F.1 – Continued from previous page

Continued on next page

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		Tabl	e F.I – Con	tinued from	in previou	s page		
ID	l	b	$N(\mathrm{H}_2)_{\mathrm{peak}}$	$FWHM_x$	$FWHM_v$	R	$N(\mathrm{H}_2)_{\mathrm{tot}}$	$N_{\rm pix}$
	[°]	[°]	$\left[\frac{1}{2}\right]$	pix	pix	pix	$\left[\frac{1}{2}\right]$	1
119	220.10	0.40	2 66	5 59	2.02	2.76	$\frac{1 \text{ cm}^2 \text{ J}}{51.95}$	- 94
440	339.10	-0.40	3.00	0.08	2.95	2.70	51.55	24
455	338.87	-0.47	3.65	3.50	2.63	2.59	46.82	21
456	338.17	-0.47	3.64	2.70	2.84	2.33	42.97	17
458	338.11	-0.46	3.64	4.67	3.00	3.09	65.79	30
450	220 02	0.49	2 62	4.40	2.10	2.65	51 91	20
409	330.93	-0.42	0.00	4.40	2.19	2.05	01.21	10
464	338.41	-0.40	3.62	3.43	2.33	2.39	43.00	18
466	338.86	-0.50	3.62	4.56	2.46	2.82	59.96	25
467	338.64	-0.45	3.62	2.01	3.10	2.03	29.45	13
460	338.62	0.44	3 62	2.80	2 32	2 33	40.72	17
403	220.02	-0.44	0.02	2.00	2.02	2.00	40.12	10
472	338.38	-0.41	3.61	3.10	2.22	2.39	40.07	18
474	338.38	-0.40	3.60	8.46	1.77	3.52	85.07	39
476	338.43	-0.41	3.59	2.82	2.19	2.33	38.70	17
481	338 25	-0.42	3 59	6.31	4.20	2.88	51.53	26
400	220 72	0.12	2 50	2.65	9.49	2.00	20.00	12
462	338.73	-0.47	5.59	5.05	2.45	2.05	30.09	15
483	338.18	-0.46	3.58	3.29	2.87	2.26	39.31	16
484	339.09	-0.41	3.58	6.07	2.06	2.99	66.74	28
485	338.85	-0.47	3.58	1.73	2.72	1.87	27.13	11
187	228.28	0.43	3 58	4.00	1.68	2 30	41 53	18
401	000.40	-0.40	0.00	4.09	1.00	⊿.09 0.10	41.00	10
488	338.25	-0.44	3.58	7.14	5.04	3.19	78.22	32
494	338.63	-0.46	3.56	2.64	3.38	2.52	43.19	20
497	338.88	-0.54	3.56	2.31	1.74	1.87	25.55	11
400	338 32	-0 41	3 55	3 74	2.00	2 30	45.00	18
= <i>JJ</i>	220.04	0.41	0.00	9.1±	1.00	0.00 0.40	11.00	10
503	339.09	-0.41	3.53	3.84	1.94	2.40	44.80	19
504	338.16	-0.48	3.53	6.17	3.02	3.19	69.67	32
505	338.62	-0.45	3.53	4.27	6.16	3.19	64.71	32
508	338.26	-0.44	3 52	2.63	5.68	2 00	60.92	28
500	228 02	-0.49	2 5 2	2.05	0.00	2.55	41 61	20
509	338.93	-0.42	3.52	3.87	2.80	2.32	41.01	20
510	338.39	-0.40	3.52	4.35	2.13	2.39	38.16	18
511	338.32	-0.42	3.52	3.86	2.65	2.19	33.45	15
512	338.15	-0.48	3.51	2.03	2.86	2.03	29.11	13
E1E	220 70	0.10	9 51	2.00	2.00	2.00	47.61	20
515	338.12	-0.40	5.51	2.39	5.91	2.32	47.01	20
516	338.84	-0.45	3.50	7.58	3.22	3.57	92.03	40
519	339.02	-0.40	3.50	2.39	1.71	1.87	26.38	11
525	338.31	-0.50	3.48	3.09	3.07	2.26	33.52	16
526	338.00	-0.46	3 / 8	3 1 3	1 74	2 11	32 11	14
520	000.09	-0.40	9.40	0.10	1.14	2.11	02.44	14
527	338.21	-0.47	3.48	3.10	4.38	3.04	63.24	29
530	338.93	-0.49	3.47	3.38	2.63	2.46	43.78	19
534	338.15	-0.49	3.46	2.34	1.94	1.87	27.83	11
541	338 23	-0.49	3.45	2 19	2 79	2.26	35.14	16
E 40	220 20	0.49	9.44	6.02	2.10	0.00	60 50	26
548	338.32	-0.42	3.44	0.95	2.10	2.00	02.58	20
549	338.93	-0.49	3.44	2.35	3.64	2.19	34.08	15
551	339.08	-0.41	3.43	2.53	2.37	2.19	33.81	15
552	338 11	-0.48	3 43	2.75	2.52	2.39	41 90	18
555	228 19	0.49	2 /2	2.10	2.02	$\frac{2.55}{2.10}$	34.05	15
000	000.42	-0.42	0.40	0.01	2.00	2.19	04.00	10
556	338.55	-0.42	3.43	5.50	6.00	2.93	49.09	27
558	338.93	-0.50	3.41	3.18	2.20	2.19	34.37	15
559	338.87	-0.48	3.41	1.92	3.23	2.19	33.82	15
563	338 41	-0 41	3.41	1 97	2.45	2.03	29.67	13
EGF	990 40	0.11	9.10	1.01	6 19	2.00 9.94	20.01	10 9E
202	338.42	-0.41	3.40	4.70	0.13	3.34	(1.32	30
569	338.15	-0.48	3.39	5.04	2.66	2.71	50.21	23
573	338.10	-0.46	3.38	2.86	1.95	2.19	35.27	15
575	338.96	-0.43	3.37	4 13	2.74	2.65	52.43	22
576	338 75	0.46	2 27	9.91	2.12	2.00	38 80	16
570	000.70	-0.40	0.07	2.31	5.10	2.20	30.00	10
581	338.28	-0.43	3.36	3.05	2.02	2.26	37.93	16
583	338.85	-0.45	3.36	3.12	1.62	1.87	25.73	11
584	338.93	-0.42	3.35	2.38	2.02	1.87	25.95	11
EOO	220 04	0.10	9.00 9.95	2.00	9 71	2.01	55 90	25
500	000.94	-0.40	0.00	0.00	0.71	2.02	00.20 00.70	<u>∠</u> 0 1.4
593	339.03	-0.40	3.34	2.37	2.50	2.11	32.78	14
594	338.26	-0.43	3.34	3.57	2.11	2.39	41.79	18
595	338.24	-0.37	3.34	3.32	3.32	2.52	45.93	20
500	338 70	0.45	2 2/	0.10	2 80	2.02	20.00	-0 19
090	000.10	-0.40	0.04	2.10	2.00	4.00	29.00	10
605	338.25	-0.42	3.31	3.13	2.21	1.95	25.39	12
606	338.73	-0.46	3.30	2.62	2.71	2.26	35.77	16
608	338.65	-0.45	3.30	3.57	2.30	2.26	38.13	16
613	338 87	_0.48	3 20	2 60	1.60	1.87	24 71	11
010	000.01	-0.40	0.00	2.00	1.09	1.01	47.11	11
010	338.65	-0.45	3.29	3.40	3.37	2.40	41.44	19

Table F.1 – Continued from previous page

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		Tabl	e F.I – Con	tinued iro	in previou	s page		
ID	l	b	$N(\mathrm{H}_2)_{\mathrm{peak}}$	$FWHM_{x}$	$FWHM_v$	R	$N(\mathrm{H}_2)_{\mathrm{tot}}$	$N_{\rm pix}$
	[°]	[°]	$\left[\frac{1}{2}\right]$	pix	pix	pix	$\left[\frac{1}{2}\right]$	-
619	220.00	$\frac{11}{0.47}$	2 90	2.52	2.00	2.46	$\frac{1_{\rm cm^2}}{20.80}$	10
010	336.20	-0.47	3.29	2.00	3.00	2.40	39.60	19
620	338.09	-0.46	3.29	1.82	2.79	1.95	26.62	12
622	338.64	-0.45	3.28	2.54	3.15	2.52	46.68	20
623	338.30	-0.48	3.28	2.60	2.90	2.33	36.33	17
626	220.20	0.42	2.07	1.61	1 59	2.00 0.00	42.02	17
020	330.32	-0.42	0.27	1.01	4.00	2.00	42.92	10
627	338.09	-0.43	3.27	2.32	2.30	1.95	28.44	12
628	338.31	-0.42	3.27	2.62	2.46	2.11	32.06	14
632	338.67	-0.46	3.26	3.29	2.19	2.26	32.28	16
634	338 77	-0.46	3.26	3.01	1 17	2 03	61.89	27
696	220.11	-0.40	2.20	1.05	2.20	2.00	26 50	10
030	338.13	-0.45	3.20	1.85	3.80	2.20	30.59	10
639	338.74	-0.47	3.25	3.64	4.05	2.76	55.48	24
641	338.82	-0.47	3.25	1.83	4.12	1.95	24.01	12
650	338 33	-0.41	3 23	3 26	2.43	2.11	31.07	14
654	220 76	0.11	2.20	2.67	5 G A	2.11	49 56	
034	338.70	-0.40	3.23	5.07	5.04	2.05	42.00	22
657	338.82	-0.46	3.22	2.36	3.43	2.19	33.18	15
662	338.89	-0.44	3.21	2.99	2.63	2.26	35.37	16
664	339.04	-0.43	3 21	3 53	1 70	2 19	29.72	15
670	220 10	0.10	2 20	4.08	2.10	2.10	50.01	25
070	336.19	-0.47	3.20	4.90	2.61	2.62	50.91	20
671	338.77	-0.46	3.20	4.11	3.46	2.93	53.53	27
672	338.92	-0.50	3.20	2.43	1.75	1.87	24.94	11
676	338.31	-0.42	3.19	2.05	2.11	1.87	24.23	11
677	338 30	0.52	3 10	3 17	1.82	2 10	32.27	15
011	000.00	-0.52	0.19	0.17	1.02	2.13	02.21	10
681	338.93	-0.42	3.18	2.06	2.65	1.95	26.58	12
689	338.94	-0.42	3.17	2.46	3.88	1.95	24.56	12
695	338.15	-0.49	3.17	1.87	2.40	1.87	23.86	11
700	338.81	-0.46	3 15	5 73	4 56	2.03	50.88	97
705	220.01	0.10	9.14	0.10	1.00	2.00 9.11	00.00 07.95	14
705	338.23	-0.44	3.14	2.55	3.33	2.11	21.80	14
706	338.63	-0.46	3.14	4.05	2.33	2.71	55.49	23
709	338.15	-0.48	3.14	2.35	2.85	1.87	26.52	11
714	338.11	-0.45	3.13	3.05	2.41	2.33	35.36	17
724	338 38	0.40	3 10	3.01	2.46	2.00	34.84	16
724	000.00	-0.40	9.10	3.01	2.40	2.20	10 11	10
730	338.94	-0.42	3.10	3.94	2.95	2.76	49.44	24
736	338.31	-0.42	3.09	4.75	1.54	2.19	36.73	15
738	338.09	-0.44	3.08	2.19	2.36	2.03	27.83	13
743	338.09	-0.43	3.04	1.58	4 01	2.19	31 44	15
747	220.00	0.10	2.04	2.01	2.01	2.10	45 57	10
141	338.32	-0.50	3.04	3.21	2.91	2.39	45.57	21
756	338.87	-0.49	3.03	1.79	3.62	2.26	33.49	16
757	339.10	-0.41	3.02	3.32	2.94	2.19	30.79	15
761	339.00	-0.41	3.01	2.80	2.12	2.11	28.16	14
762	220 20	0.41	2 01	2.00	4.05	2.50	44.02	91
702	000.09	-0.41	3.01	0.01	4.05	2.09	44.03	21
766	338.92	-0.51	3.00	2.51	4.78	2.93	58.06	27
767	338.32	-0.44	3.00	2.55	3.28	2.19	31.35	15
771	338.45	-0.42	2.99	4.15	4.00	2.71	43.32	23
780	338.03	0.49	2.08	118	8.68	3 14	67 10	21
700	000.00	-0.49	2.30	4.10	0.00	0.14	40.00	01
786	338.20	-0.43	2.97	4.00	4.10	2.71	48.20	23
787	338.91	-0.53	2.97	3.53	2.00	2.19	30.76	15
789	338.12	-0.50	2.96	2.54	2.13	2.11	29.34	14
790	338.37	-0.40	2.96	3.62	3.54	2.71	47.90	23
702	330.01	_0.40	2.06	2.62	7.07	2 80	18 16	
190	009.01	-0.40	2.30	2.02	1.01	2.02	40.40	20 00
794	338.14	-0.48	2.96	5.21	3.54	2.71	41.30	23
795	338.18	-0.46	2.96	3.52	1.74	2.11	27.88	14
796	338.64	-0.45	2.96	2.71	2.31	1.95	26.54	12
800	338 71	-0.47	2.95	2 10	2 36	1.05	26.85	12
000	220.11	-0.41	2.50	2.10	2.00	1.50	20.00	12
801	338.95	-0.42	2.95	4.03	3.81	3.29	13.55	34
803	338.28	-0.43	2.95	4.34	2.31	2.52	38.05	20
805	338.52	-0.43	2.94	2.01	3.06	2.19	31.25	15
806	338 42	-0 41	2.94	2.43	2.17	2.03	25.85	13
000	990 CO	0.45	2.04	2.40	2.11	2.00 9.65	49.49	10
009	336.03	-0.40	2.94	5.87	3.43	2.00	42.43	22
812	338.94	-0.49	2.94	2.39	2.75	2.26	33.87	16
819	338.11	-0.45	2.92	3.38	2.40	2.33	32.89	17
821	338.32	-0.41	2.92	2.78	4.96	2.52	37.38	20
\$ <u>5</u> 1	338 55	0.49	2.02	2.10	2.00	<u>.</u> 9 29	35.00	17
044	000.00	-0.42	2.92	2.01	0.40	⊿. J J	33.02	11
823	339.00	-0.40	2.92	2.54	3.53	2.52	39.96	20
824	338.09	-0.46	2.92	2.71	2.71	2.39	38.23	18
825	339.00	-0.41	2.92	2.37	2.88	1.95	25.02	12
826	338 86	-0.50	2 02	2 80	2 03	2 10	30 56	15
040	000.00	-0.00	4.34	2.00	2.00	4.13	00.00	10

Table F.1 – Continued from previous page

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		Table	e F.I - Com	inued no	ii previou	s page		
ID	l	b	$N({ m H}_2)_{ m peak}$	$FWHM_x$	$FWHM_y$	R	$N({ m H}_2)_{ m tot}$	$N_{\rm pix}$
	[°]	[°]	$\begin{bmatrix} 1 \\ -2 \end{bmatrix}$	pix	pix	pix	$\begin{bmatrix} 1 \\ -2 \end{bmatrix}$	
833	338.00	0.44	$\frac{1 \text{ cm}^2 \text{ J}}{2 01}$	2.07	3 /1	2.03	$\frac{1 \text{ cm}^2 \text{ J}}{25.15}$	13
000	338.09	-0.44	2.91	2.07	0.41	2.05	20.10	10
835	338.85	-0.47	2.90	3.57	3.08	2.39	36.98	18
841	338.39	-0.40	2.90	3.31	3.20	2.52	41.88	20
843	338.41	-0.42	2.89	2.30	1.81	1.87	23.09	11
818	338 34	0.41	2.80	2.08	1.57	2.03	26.28	12
040	000.04	-0.41	2.03	1.50	2.10	1.05	20.20	10
851	338.09	-0.43	2.88	1.50	3.19	1.87	22.86	11
855	338.47	-0.42	2.88	3.92	2.29	2.39	39.05	18
860	338.91	-0.44	2.87	2.76	1.97	2.03	25.09	13
862	338 48	-0.43	2.87	1.82	3.44	2.03	24.75	13
002	220.00	-0.40	2.01	1.02	0.47	2.00	24.10	15
808	339.00	-0.41	2.80	2.94	2.47	2.19	30.87	10
873	338.77	-0.46	2.86	4.29	2.22	2.59	44.53	21
875	338.78	-0.45	2.86	2.38	3.50	2.11	27.94	14
881	338.39	-0.41	2.85	4.44	3.62	2.99	56.27	28
882	338 41	0.40	2.85	0.00	3.04	2.22	35.67	17
002	220.41	-0.40	2.00	2.22	0.40	2.00	00.10	10
886	338.96	-0.41	2.84	2.74	2.40	2.03	26.13	13
889	339.07	-0.41	2.84	3.20	2.86	2.19	30.98	15
894	338.86	-0.50	2.83	2.99	1.41	1.87	23.21	11
896	338 13	-0.46	2.83	3 60	2.03	2.03	25.17	13
807	220.72	0.10	2.00	0.00	2.00	2.00	20.11	10
897	338.73	-0.47	2.82	2.40	3.07	2.20	30.70	10
898	338.38	-0.39	2.82	6.66	3.16	3.34	65.60	35
900	338.93	-0.48	2.82	2.71	2.23	2.11	29.19	14
909	338.35	-0.41	2.80	5.46	1.97	2.76	47.88	24
010	228.00	0.11	2.00	1.64	2 50	2.10	25.99	19
912	336.09	-0.44	2.80	1.04	3.50	2.05	20.00	10
921	338.33	-0.51	2.78	5.40	2.97	2.59	37.55	21
922	338.14	-0.48	2.78	2.09	3.20	2.19	29.89	15
923	339.04	-0.42	2.78	4.01	1.68	2.11	28.54	14
024	338.82	0.45	2.78	2 70	3 66	2 50	30.57	21
924	220.02	-0.49	2.10	2.10	0.00	2.03	09.01	10
927	338.31	-0.43	2.78	2.18	3.32	2.26	33.61	16
930	338.08	-0.42	2.77	2.29	2.81	2.03	26.13	13
931	338.93	-0.50	2.77	4.96	1.52	2.26	30.70	16
033	338.83	-0.46	2 77	2 47	3 37	2 59	41.85	21
004	220.00	-0.40	2.11	2.47	4 10	2.09	26.00	17
934	338.32	-0.42	2.11	1.97	4.12	2.33	30.09	17
936	338.08	-0.45	2.76	1.66	2.60	1.87	21.71	11
939	338.19	-0.48	2.76	1.74	2.69	1.87	23.72	11
940	338 85	-0.47	2.76	2.98	3.87	2.39	37.16	18
041	220 10	0.17	2.76	2.00	2.66	2.00	20.54	16
941	330.10	-0.47	2.70	2.40	2.00	2.20	30.34	10
944	338.83	-0.45	2.75	6.34	7.60	3.74	85.64	44
949	338.82	-0.46	2.74	3.20	2.57	2.39	37.41	18
950	338.10	-0.46	2.74	3.28	2.50	2.33	36.22	17
053	338 78	0.46	2.74	4.17	2.04	2.46	38.05	10
955	330.10	-0.40	2.14	4.17	2.04	2.40	38.05	19
954	338.40	-0.40	2.74	6.99	3.35	2.99	47.57	28
956	338.81	-0.48	2.73	5.81	5.17	2.93	52.59	27
963	338.41	-0.41	2.73	2.17	2.39	1.87	21.75	11
970	338.60	-0.49	2.72	1 / 9	4 69	2 11	26 56	14
079	220.05	-0.40	2.12	1.45	1.05	0.11	20.00	14
973	338.41	-0.40	2.72	1.57	3.93	2.11	20.85	14
980	338.37	-0.40	2.70	2.67	3.47	2.03	25.73	13
981	338.20	-0.47	2.70	3.26	5.64	2.65	44.85	22
987	338.31	-0.50	2.69	2.55	2.70	2.26	30.31	16
004	338 37	0.40	2.60	2 56	4.04	2 10	30.63	15
1000	000.01	-0.40	2.03	2.00	4.04	2.13	30.03	10
1002	338.09	-0.46	2.68	4.93	2.33	2.39	35.03	18
1005	338.93	-0.48	2.68	2.58	2.44	2.19	29.21	15
1009	338.86	-0.47	2.68	2.71	3.94	2.26	29.48	16
1011	339.03	-0.40	2.67	2 44	3.07	2 26	31.96	16
1010	990 04	0.40	2.01	2.11	2.07	2.20	40.61	10
1010	338.94	-0.42	2.00	3.03	3.90	2.82	49.01	20
1022	338.11	-0.47	2.65	1.75	3.64	2.26	32.96	16
1025	338.32	-0.40	2.65	5.67	3.46	2.59	40.37	21
1027	$338 \ 76$	-0.46	2.64	2.44	2.77	2.19	28.37	15
1021	330.10	0.40	2.01	<u> </u>	2.11	2.10	20.01	19
1001	330.00	-0.42	2.00	4.11	0.00	2.00	24.00	10
1032	338.31	-0.52	2.63	2.77	1.89	2.03	26.27	13
1033	338.93	-0.47	2.63	3.04	3.09	2.26	28.61	16
1037	338.92	-0.47	2.63	3.53	2.26	2.11	25.46	14
10/0	338 65	0.46	2.00	1 72	2.20	11 9 11	25.00	1/
1040	000.00	-0.40	2.02	1.10	0.97	4.11 0.05	20.92	14
1046	338.31	-0.40	2.62	4.60	2.31	2.65	39.79	22
1049	338.23	-0.46	2.61	2.09	2.92	2.19	27.37	15
1051	338.39	-0.40	2.61	2.43	2.66	1.95	21.65	12
1052	338 19	_0 /1	2 60	2 24	2.28	1 05	22 77	19
1000	000.44	-0.41	2.00	2.04	2.20	1.30	22.11	14

Table F.1 – Continued from previous page

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		Tabl	e F.I - Con	tinued iro	m previou	s page		
ID	l	b	$N({\rm H}_2)_{\rm peak}$	$FWHM_{x}$	$FWHM_{v}$	R	$N({\rm H}_2)_{\rm tot}$	$N_{\rm pix}$
	[0]	[0]	$\left[\frac{1}{2}\right]$	nix	nix	nix	$\begin{bmatrix} 1 \\ -1 \end{bmatrix}$	1
1056	11	050	$\frac{1 \text{ cm}^2 \text{ J}}{2.60}$	4.00	2.24	2.46	$\frac{1_{\rm cm}^2}{26.40}$	10
1050	338.78	-0.50	2.60	4.00	2.24	2.40	30.40	19
1058	338.73	-0.48	2.60	2.31	2.08	1.87	20.14	11
1059	338.86	-0.49	2.60	2.60	3.83	2.33	32.30	17
1061	339.03	-0.42	2.59	2.68	3.08	2.26	30.22	16
1069	220 00	0.12	2.00	2.00	0.00	2.20	26.06	10
1002	330.02	-0.45	2.59	3.33	2.01	2.39	30.00	18
1072	338.32	-0.51	2.57	6.73	3.80	2.65	39.47	22
1074	338.82	-0.48	2.57	2.42	3.00	2.33	33.41	17
1075	338.18	-0.46	2.57	2.87	2.45	2.03	23.19	13
1077	338 18	0.47	2.57	1.88	3 66	2.03	22.67	13
1077	330.10	-0.47	2.57	1.00	5.00	2.05	22.07	10
1079	338.92	-0.52	2.56	2.09	6.18	2.59	38.36	21
1089	338.93	-0.50	2.54	2.79	1.70	1.95	22.25	12
1098	338.08	-0.42	2.53	2.58	4.52	2.52	37.33	20
1099	$338\ 32$	-0.51	2.53	6.36	2.29	2.39	31.33	18
1100	228.02	0.01	2.00	4.19	2.20	2.00	40.69	10
1100	338.92	-0.49	2.05	4.15	2.40	2.00	40.02	22
1103	338.78	-0.47	2.53	1.82	2.66	1.95	23.11	12
1107	338.19	-0.47	2.52	4.23	4.07	2.46	35.50	19
1109	338.68	-0.46	2.52	3.00	3.41	2.59	37.29	21
1110	338.66	0.46	2 52	2.36	3.03	2 33	30.75	17
1110	220.00	-0.40	2.52	2.50	0.00	2.00	01.00	17
1111	338.23	-0.54	2.51	2.72	2.88	2.33	31.96	17
1112	338.92	-0.49	2.51	3.51	1.76	2.19	29.06	15
1115	338.75	-0.47	2.51	2.49	3.75	2.19	27.81	15
1118	338.96	-0.44	2.50	2.76	2.17	1.87	20.08	11
1100	220.20	0.11	2.00	2.70	2.11	0.00	20.00	17
1122	338.35	-0.41	2.50	3.72	5.95	2.33	31.33	17
1132	338.11	-0.47	2.48	3.15	1.93	1.95	22.37	12
1140	339.03	-0.40	2.47	2.94	4.73	2.33	28.47	17
1144	338.68	-0.48	2.46	2.10	2.74	2.03	23.67	13
1154	338 76	0.48	2.10	2.10	2.11 2.07	2.00	23.01	13
1104	330.70	-0.40	2.40	0.10	2.07	2.05	20.11	10
1155	339.01	-0.41	2.45	3.40	2.87	2.03	23.04	13
1157	338.42	-0.41	2.45	3.02	2.06	2.11	24.91	14
1163	338.86	-0.48	2.44	3.54	3.71	2.52	34.61	20
1164	338 33	0.51	2 44	2 70	1 77	1.05	21.37	19
1104	220.00	-0.51	2.44	2.13	2.11	1.50	21.57	12
1100	338.79	-0.45	2.44	3.51	2.90	2.20	29.53	10
1167	338.65	-0.45	2.44	2.83	4.72	2.52	37.71	20
1170	338.32	-0.40	2.44	3.62	2.71	2.46	34.45	19
1174	338 31	-0.41	2.43	1.76	3.56	2.11	$25\ 21$	14
1196	220 02	0.11	2.10	2.16	2.00	1.97	20.21	11
1100	330.93	-0.44	2.41	2.10	2.09	1.01	20.10	11
1193	338.41	-0.43	2.40	2.91	4.72	2.76	43.26	24
1199	338.43	-0.41	2.38	3.00	2.42	2.26	28.97	16
1201	338.13	-0.47	2.38	2.13	3.51	2.26	29.05	16
1202	228 20	0.46	0.38	9.15	2.04	1.87	10.05	11
1203	000.29	-0.40	2.30	2.10	2.94	1.07	19.90	11
1204	338.75	-0.50	2.38	2.94	1.00	1.95	22.68	12
1205	338.81	-0.47	2.38	4.46	2.38	2.33	28.46	17
1206	338.80	-0.46	2.38	4.55	1.80	2.11	23.51	14
1210	338 30	-0.41	2 37	2.82	1.63	1.87	19.66	11
1011	000.00	-0.44	2.01	2.02	2.00	0.10	10.00	15
1211	338.31	-0.44	2.37	2.29	3.39	2.19	28.20	15
1216	338.34	-0.50	2.36	1.40	3.81	1.95	23.55	12
1226	338.64	-0.44	2.34	2.71	1.73	1.95	21.87	12
1242	338.07	-0.43	2.33	4.49	4.30	2.59	37.37	21
1246	338.84	0.48	2.00	1.62	2.74	1.87	10.24	11
1240	000.04	-0.40	2.32	1.02	2.74	1.07	19.24	11
1248	339.04	-0.41	2.32	1.74	2.79	1.95	21.91	12
1249	338.25	-0.44	2.32	2.62	3.13	2.03	22.78	13
1251	339.02	-0.40	2.32	3.15	4.01	2.39	31.33	18
1252	330.00	-0.41	2 32	3.96	1 03	2 11	25 44	14
1050	000.00	-0.41	2.02	0.00	1.55	2.11	20.44	15
1258	338.12	-0.47	2.31	2.82	2.50	2.19	25.60	15
1259	338.69	-0.48	2.31	1.80	2.68	1.87	18.46	11
1261	338.20	-0.48	2.31	2.76	7.84	2.39	30.11	18
1262	338 70	-0.46	2.31	4 21	3.62	1.95	20.75	12
1000	990 00	0.40	2.01	1.41 9.10	1 71	1.05	20.10	10
1200	330.09	-0.40	2.30	3.10	1./1	1.90	22.14	12
1271	338.11	-0.46	2.30	3.51	1.85	2.03	23.63	13
1275	338.16	-0.47	2.29	1.88	2.94	2.03	23.12	13
1278	338.75	-0.50	2.29	2.67	9.65	3.04	48.86	29
1991	228 75	0.46	2.20	9.75	0.00	2.01	20.00	16
1201	000.70 000.70	-0.40	2.29	2.10	2.70	2.20	∠9.0ð	10
1288	338.68	-0.46	2.28	3.35	1.92	2.19	26.11	15
1289	338.94	-0.48	2.28	5.99	1.85	2.19	25.31	15
1290	339.03	-0.40	2.27	3.28	2.19	2.11	23.38	14
1202	338.02	-0.43	2.27	6 72	1 35	2 76	12 14	24
1490	000.92	-0.40	4.41	0.14	4.00	2.10	44.14	24

Table F.1 – Continued from previous page

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Table F 1 –	Continued	from	previous pag	ρ
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		Tabi		tillueu 110	in previou	s page		
ID	l	b	$N(\mathrm{H}_2)_{\mathrm{peak}}$	$FWHM_x$	$FWHM_v$	R	$N(\mathrm{H}_2)_{\mathrm{tot}}$	$N_{\rm pix}$
	[0]	[0]	$\left[\frac{1}{2}\right]$	nix	nix	nix	$\begin{bmatrix} 1 \\ -1 \end{bmatrix}$	1
1907	11	0.44	$\frac{1 \text{ cm}^2}{2.04}$	PIX	1.44	<u> </u>	$\frac{1 \text{ cm}^2 \text{ J}}{22.45}$	1.9
1307	338.08	-0.44	2.24	5.37	1.44	2.03	23.45	13
1309	338.28	-0.43	2.24	3.77	1.77	2.03	22.34	13
1312	338.21	-0.49	2.24	2.75	2.46	1.87	20.47	11
1314	338 50	0.45	0.03	3.06	4 20	2.26	27 20	16
1014	000.09	-0.45	2.20	0.00	4.20	2.20	21.29	10
1315	338.39	-0.40	2.23	3.21	2.08	2.03	22.14	13
1318	338.77	-0.47	2.23	4.02	2.83	2.03	20.97	13
1322	338.80	-0.48	2.22	4.08	1.91	2.19	26.50	15
1224	220.06	0.10	2.22	1.00	2.05	2.10	20.00	19
1024	339.00	-0.42	2.22	1.77	2.95	2.03	22.80	10
1325	338.82	-0.47	2.22	4.72	1.58	2.03	23.03	13
1336	338.24	-0.46	2.21	3.37	2.03	2.11	23.68	14
1341	338 24	-0.45	2.20	2.75	6.28	2.59	35.06	21
1949	220.21	0.10	2.20	2.10	0.20 9 EE	1.05	01.05	10
1545	330.33	-0.54	2.19	2.10	5.55	1.95	21.50	12
1346	338.61	-0.44	2.19	2.50	2.28	1.87	18.92	11
1348	338.31	-0.49	2.19	2.14	5.71	2.39	29.30	18
1351	338.00	-0.43	2 18	2.04	2.78	1.87	18 20	11
1057	220.00	-0.40	2.10	2.04	2.10	1.07	10.25	10
1357	339.06	-0.41	2.18	2.29	2.16	1.95	21.11	12
1359	338.58	-0.45	2.17	1.87	4.60	2.46	31.17	19
1361	339.05	-0.41	2.17	2.60	2.69	2.03	22.80	13
1363	338 55	0.49	2.17	1 69	2 70	1.97	10.60	11
1000	000.00	-0.42	2.11	1.00	2.19	1.07	10.50	11
1364	338.97	-0.41	2.17	2.00	3.08	1.87	19.79	11
1365	338.93	-0.49	2.17	5.81	11.05	2.88	44.12	26
1366	338.52	-0.43	2.17	1.89	2.52	1.95	21.52	12
1375	228 11	0.46	2.16	3.46	4.06	2.10	26.28	15
1010	000.11	-0.40	2.10	5.40	4.00	2.13	20.20	10
1380	338.51	-0.41	2.15	5.50	2.41	2.52	32.87	20
1382	338.21	-0.48	2.15	2.33	3.05	1.95	20.87	12
1384	338.31	-0.46	2.14	2.61	2.04	1.87	18.32	11
1386	338.61	0.45	2.14	1 82	2.01	1.87	10.15	11
1000	336.01	-0.45	2.14	1.02	2.42	1.07	19.10	11
1392	338.21	-0.46	2.13	7.33	1.18	2.52	33.94	20
1396	338.18	-0.49	2.12	2.00	2.43	1.87	18.83	11
1399	338 39	-0.40	2.12	2.20	2.23	1.87	18 85	11
1400	220.02	0.19	9.19	2.15	2.20	1.07	18.06	11
1400	336.93	-0.48	2.12	3.13	2.10	1.07	10.90	11
1405	338.12	-0.46	2.11	2.77	2.00	1.87	19.07	11
1408	338.13	-0.49	2.10	7.10	5.57	2.46	31.00	19
1411	339.02	-0.40	2.10	8 15	3.09	2.93	4473	27
1410	220.72	0.10	2.10	0.10	2.05	0.11	04.19	14
1419	338.12	-0.40	2.08	2.14	5.05	2.11	24.18	14
1421	339.09	-0.41	2.08	6.14	1.73	2.03	22.01	13
1426	338.66	-0.45	2.07	3.70	1.57	1.95	18.95	12
1/128	339.05	-0.41	2.07	4 41	1.05	2 33	28.24	17
1420	220.70	-0.41	2.01	1.11	1.50	1.07	10.49	11
1458	338.70	-0.40	2.05	2.20	2.01	1.07	19.42	11
1442	338.12	-0.50	2.05	2.24	2.15	1.87	18.54	11
1444	338.80	-0.48	2.04	3.64	2.13	1.87	18.07	11
1454	338 03	-0.50	2.03	4 36	5 94	2.76	39 36	24
1400	220.00	0.00	2.00	9.70	1 75	1.07	10 05	⊥⊥ I 11
1400	338.93	-0.48	2.02	2.18	1.75	1.87	18.35	11
1464	338.95	-0.42	2.01	5.24	3.37	2.59	34.32	21
1466	339.10	-0.41	2.01	2.80	3.20	1.87	17.33	11
1473	$338\ 24$	-0.45	2.00	2.75	2.79	1.95	20.27	12
1/7/	220 00	0 50	1 00	2.10	1 79	1.00	10.00	11
14(4	338.92	-0.50	1.99	3.29	1.13	1.81	18.02	11
1475	338.63	-0.46	1.99	1.68	3.27	1.87	19.41	11
1478	338.75	-0.48	1.98	3.83	4.53	2.19	25.00	15
1485	338 02	-0 47	1 97	2 21	2 50	1 95	19.61	12
1400	990 AF	0.41	1.05	2.21 1 55	4.60	1.05	10.01	10
1492	338.25	-0.41	1.95	1.55	4.60	1.95	18.79	12
1495	338.91	-0.45	1.94	5.13	1.67	2.11	22.28	14
1513	338.79	-0.50	1.90	3.47	3.03	1.95	18.84	12
1515	338 68	-0.46	1 80	3 22	2.30	1 95	18 85	12
1515	000.00	0.40	1.00	0.44	2.03	1.05	10.00	11
1517	338.16	-0.48	1.89	3.66	2.43	1.87	18.60	11
1522	338.97	-0.41	1.86	3.08	2.27	1.87	17.50	11
1527	338.83	-0.46	1.85	3.54	4.50	1.95	19.37	12
1525	222 61	0.45	1 70	0.01	2 99	1.00	17 19	11
1000	000.01	-0.40	1.19	2.29	0.44	1.01	17.14	11
1538	338.80	-0.48	1.77	3.13	2.48	1.87	17.05	11
1539	338.79	-0.47	1.76	4.52	1.12	1.87	17.07	11
1540	338.09	-0.43	1.76	4.58	3.86	1.95	17.82	12
1546	330 00	0.40	1 79	5 75	1 51	2.00	20.96	14
1040	JJ0.0U	-0.49	1.70	0.70	1.01	⊿.11	20.80	14