Sternentstehung - Star Formation

Winter term 2017/2018

Henrik Beuther & Thomas Henning

17.10	*Today: Introduction & Overview*	*(H.B.)*
24.10	*Physical processes I*	*(H.B.)*
31.10	no lecture – Reformationstag	
07.11	**Physcial processes II**	**(H.B.)**
14.11	Molecular clouds as birth places of stars	(H.L.)
21.11	Molecular clouds (cont.), Jeans Analysis	(H.B.)
28.11	Collapse models I	(H.B.)
05.12	Collapse models II	(T.H.)
12.12	Protostellar evolution	(T.H.)
19.12	Pre-main sequence evolution & outflows/jets	(T.H.)
09.01	Accretion disks I	(T.H.)
16.01	Accretion disks II	(T.H.)
23.01	High-mass star formation, clusters and the IMF	(H.B.)
30.01	Planet formation	(T.H.)
06.02	Examination week, no star formation lecture	

Book: Stahler & Palla: The Formation of Stars, Wileys

More Information and the current lecture files: http://www.mpia.de/homes/beuther/lecture_ws1718.html

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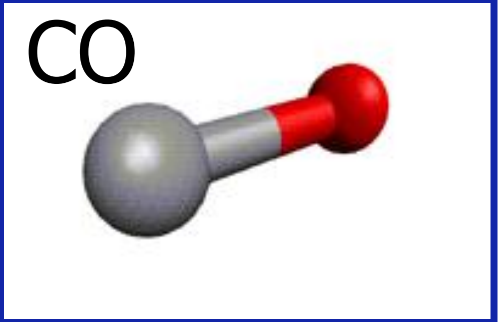
Last Week

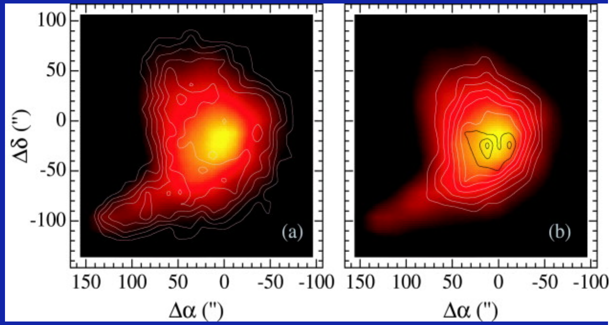
- Main tools: Spectral line emission, and thermal emission and extinction from dust.



- Molecules interesting for themselves and chemistry.

- However, also extremely useful to trace physical processes.

- Molecules deplete on grains at low temperatures.



- Discussed main cooling and heating processes.

- Discussed basic line radiation transfer and column density determination.

Molecular column densities II

Solving for upper level column density N_u

$$N_u \sim \tau / (1 - e^{-\tau}) \int T \, dv$$

N_u relates to the total column density N_{tot}

$$N_{tot} = N_u / g_u \, \exp(E_u / kT) \;\; Q$$

For linear molecule (CO), partition function Q can be approximated

$$Q = kT/hB.$$

For more complex molecules Q can become very complicated.

Conversion from CO to H_2 column densities

Classical way to derive conversion factors from CO to H_2 column densities:

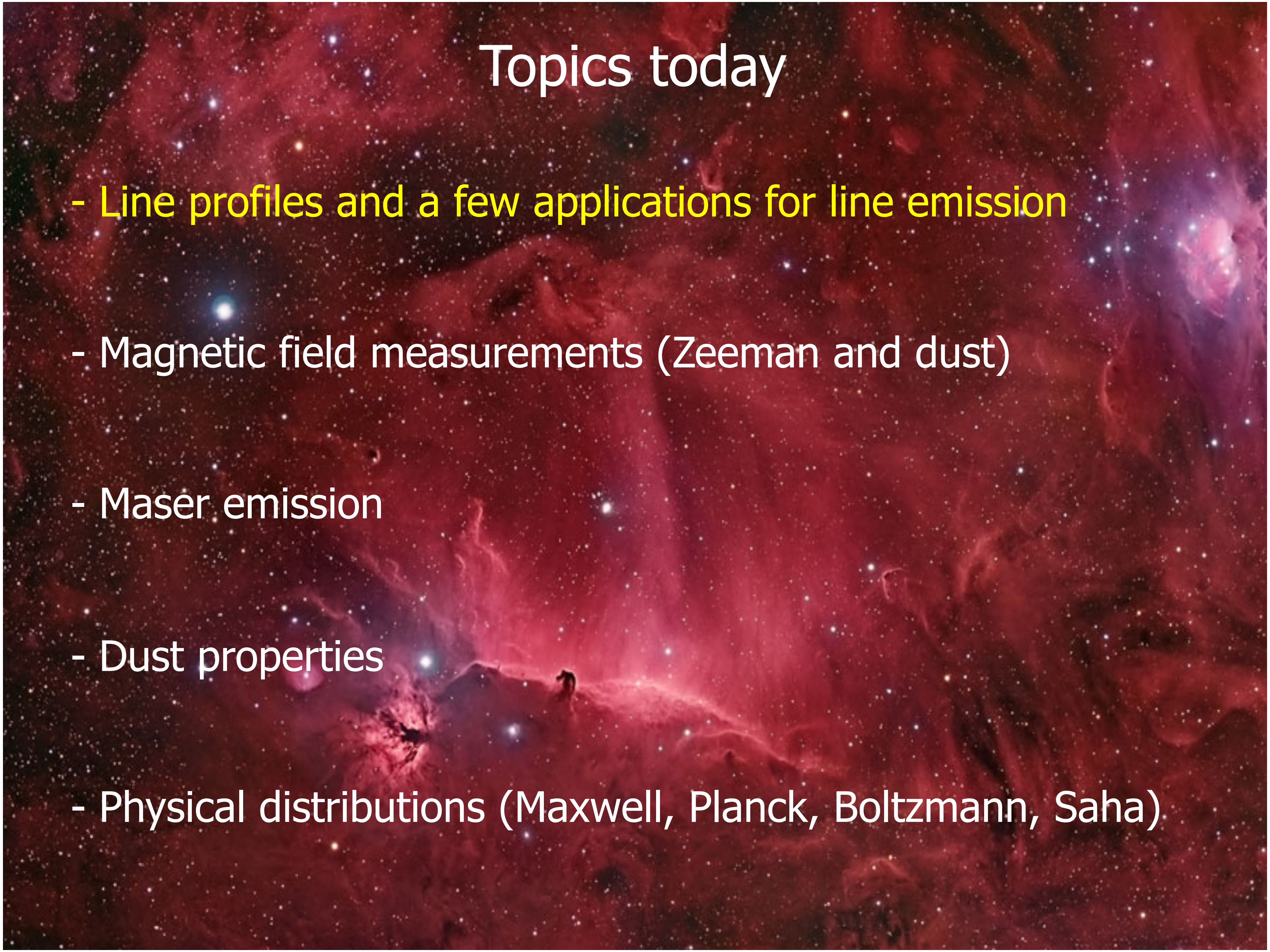
1) Derive ratio between colour excess E_{B-V} and optical extinction A_v

   $A_v = 3.1\ E_{B-V}$ (Savage and Mathis, 1979)

2) The ratio $N(H_2)/E_{B-V}$: One can measure the H_2 column density, e.g., directly from UV Absorption lines.

3) The ratio $N(CO)/A_v$: In regions of molecular gas emission, one can estimate A_v by star counts in the Infrared regime

⇒ Combining these three ratios: CO → H_2 column densities.



Topics today

- Line profiles and a few applications for line emission

- Magnetic field measurements (Zeeman and dust)

- Maser emission

- Dust properties

- Physical distributions (Maxwell, Planck, Boltzmann, Saha)

Line broadening

- Natural line broadening: Disturbance of molecule by zero-point vibrations of electromagnetic field

$$d\nu = 32\pi^3\nu^3 \mu^2/(3hc^3)$$

(μ: Dipole moment)

For CO(1-0) → $d\nu \sim 3.5\text{x}10^{-8}$ Hz or $dv \sim 9\text{x}10^{-14}$ km/s → Negligable!

- Pressure broadening: Arises from collisions between molecules.
  Quantum-mechanical problem of intermolecular forces.
  → At densities of star-forming regions negligable.

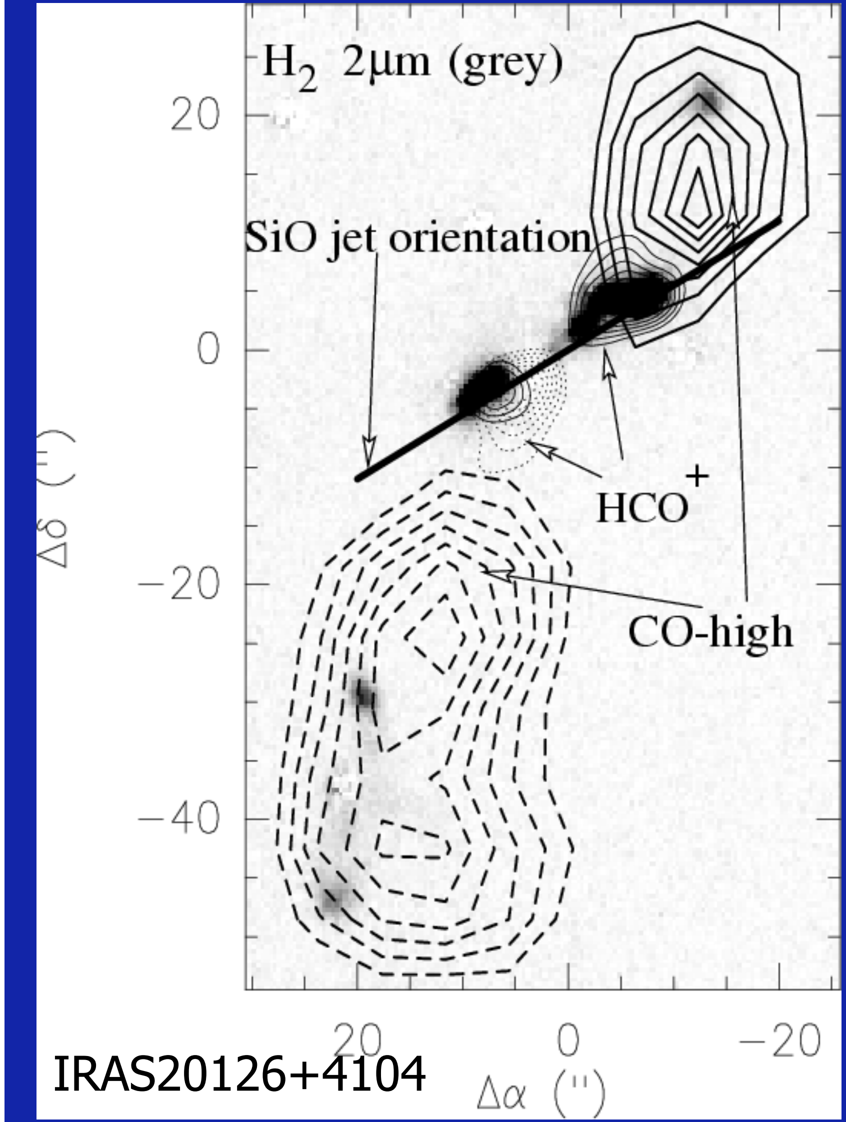
- Thermal line broadening: Thermal motions of gas cause doppler broadening:

$$dv = \sqrt{(8\ln 2\, kT/m_{mol})}$$

→ $dv(NH_3@30K) \sim 0.28$ km/s

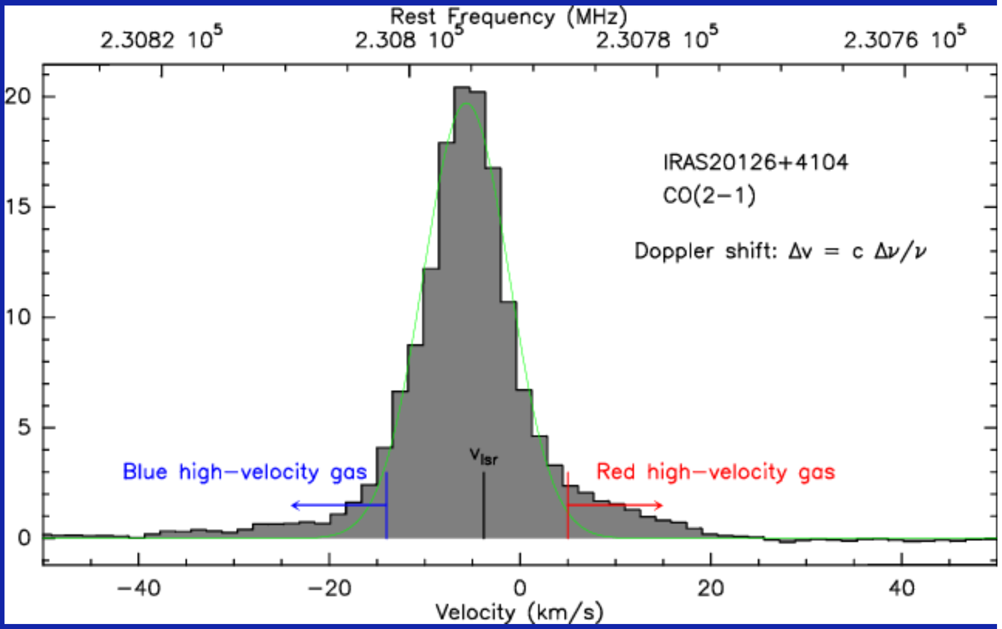
→ Other physical effects: Line broadening due to outflow motions, rotation …

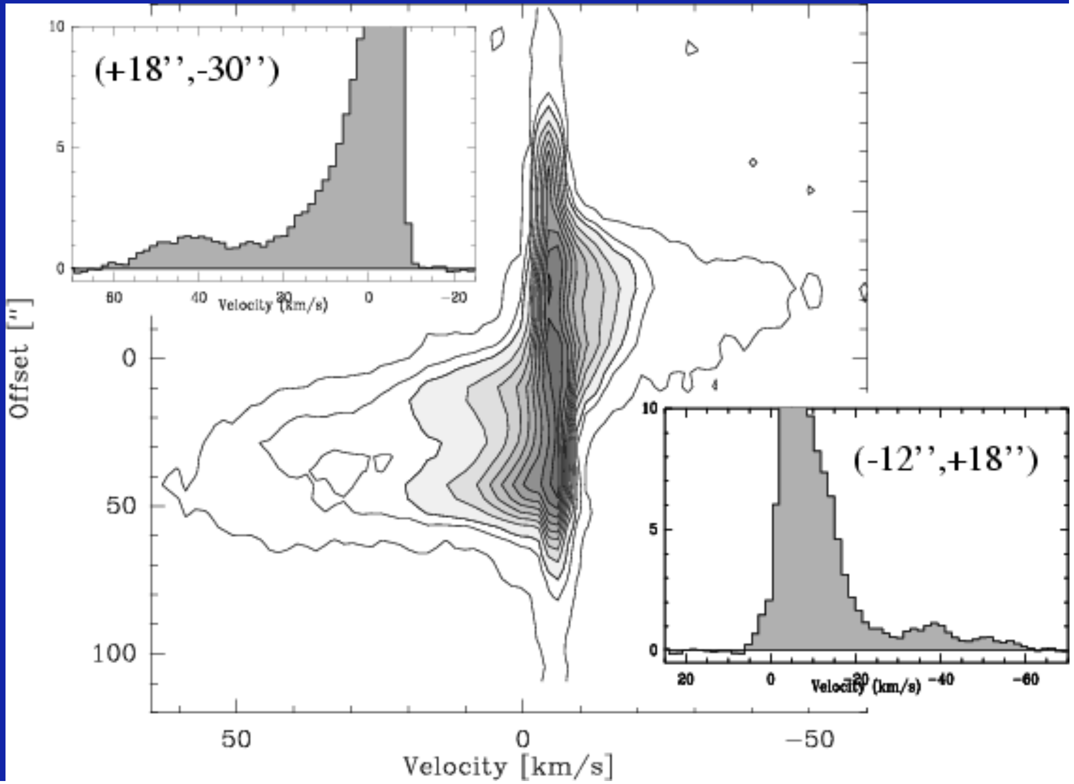
Molecular outflows



IRAS20126+4104

Lebron et al. 2006





Molecular gas structure of the Galaxy based on CO observations

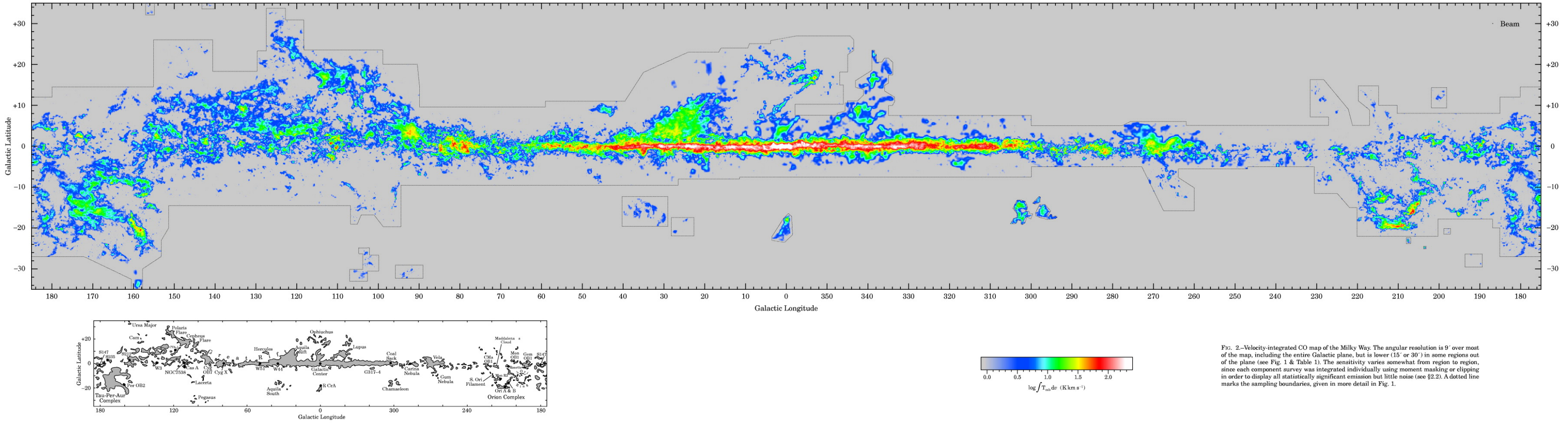


FIG. 2.–Velocity-integrated CO map of the Milky Way. The angular resolution is 9′ over most of the map, including the entire Galactic plane, but is lower (15′ or 30′) in some regions out of the plane (see Fig. 1 & Table 1). The sensitivity varies somewhat from region to region, since each component survey was integrated individually using moment masking or clipping in order to display all statistically significant emission but little noise (see §2.2). A dotted line marks the sampling boundaries, given in more detail in Fig. 1.

Dame et al. 2001

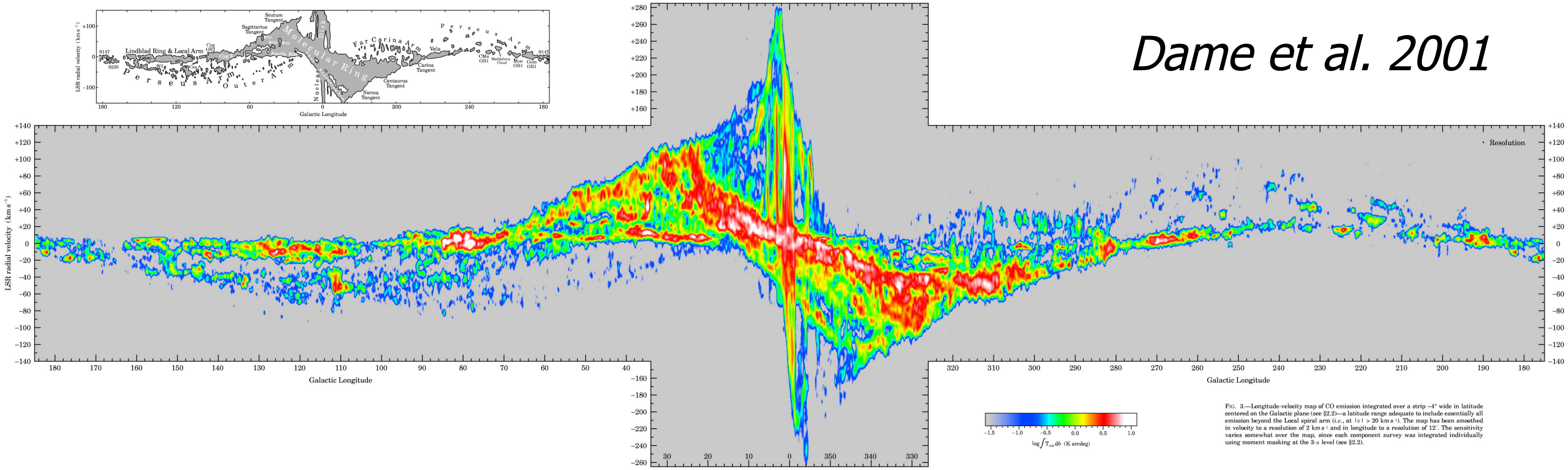
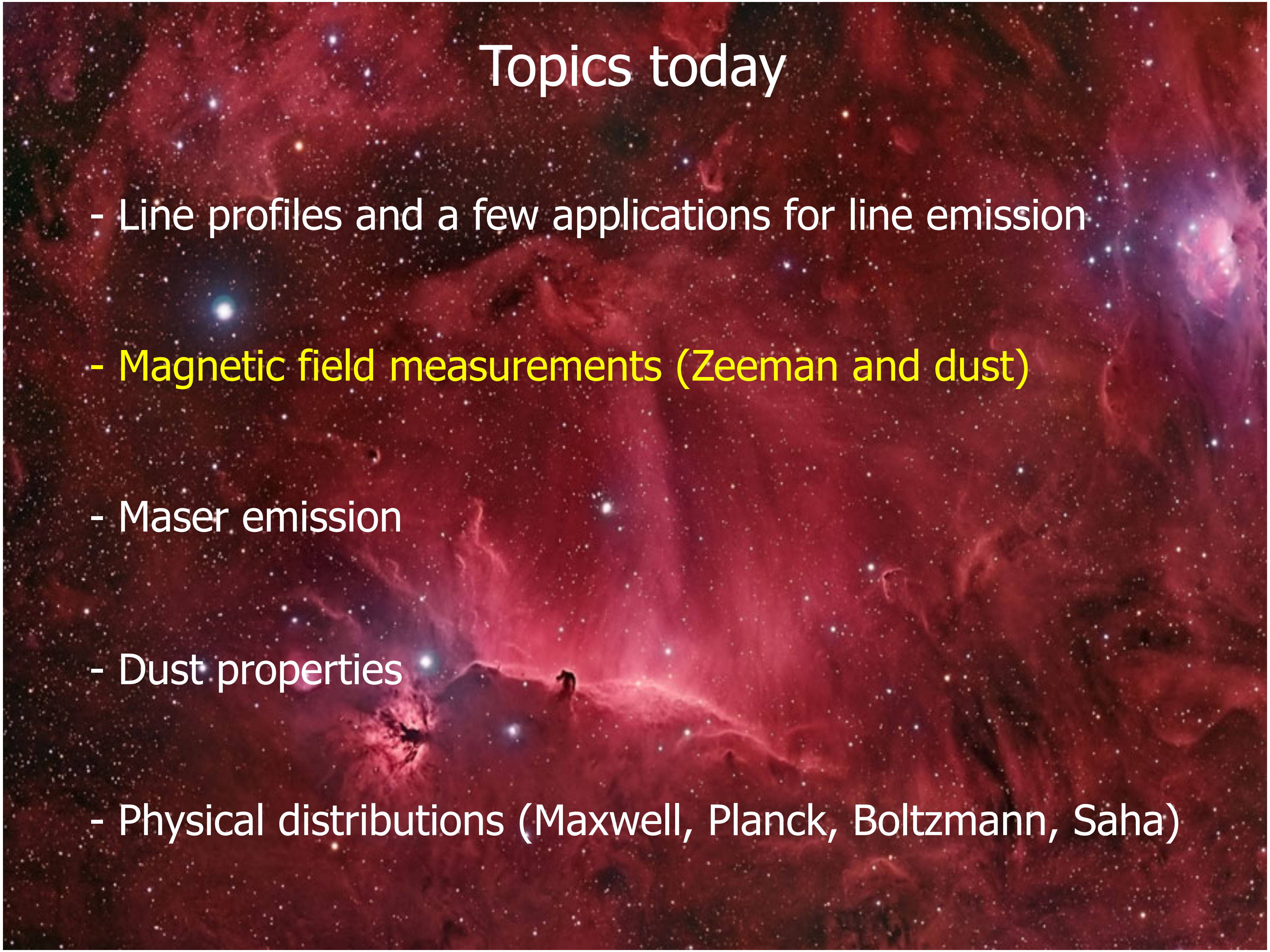


FIG. 3.—Longitude-velocity map of CO emission integrated over a strip ~4° wide in latitude centered on the Galactic plane (see §2.2)—a latitude range adequate to include essentially all emission beyond the Local spiral arm (*i.e.*, at $|v| > 20$ km s^{-1}). The map has been smoothed in velocity to a resolution of 2 km s^{-1} and in longitude to a resolution of 12′. The sensitivity varies somewhat over the map, since each component survey was integrated individually using moment masking at the 3-s level (see §2.2).



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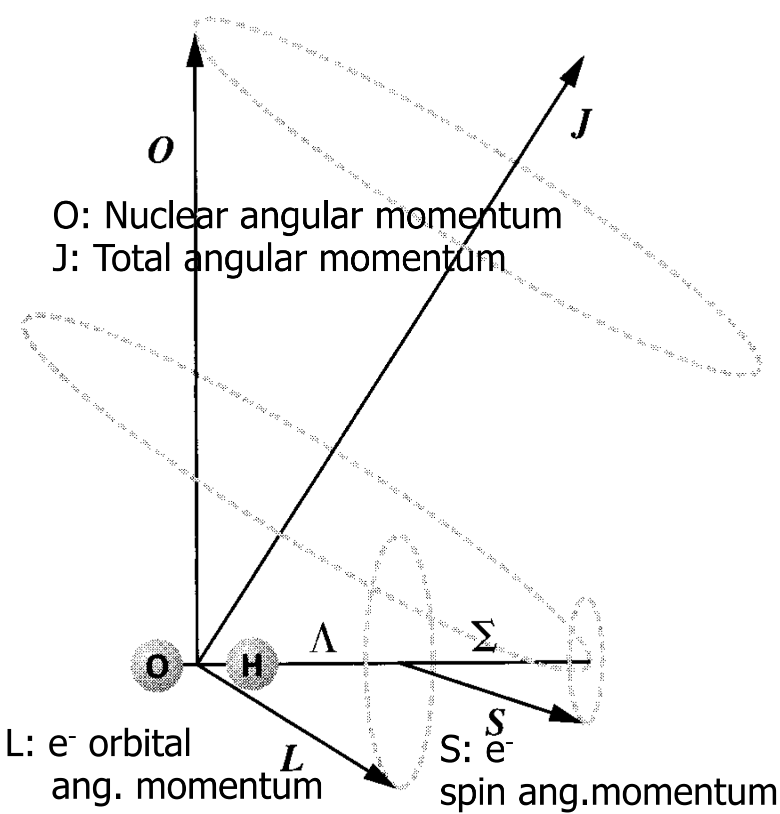
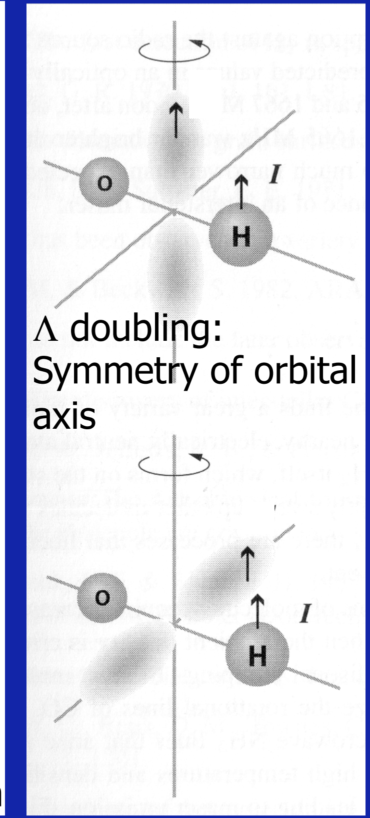
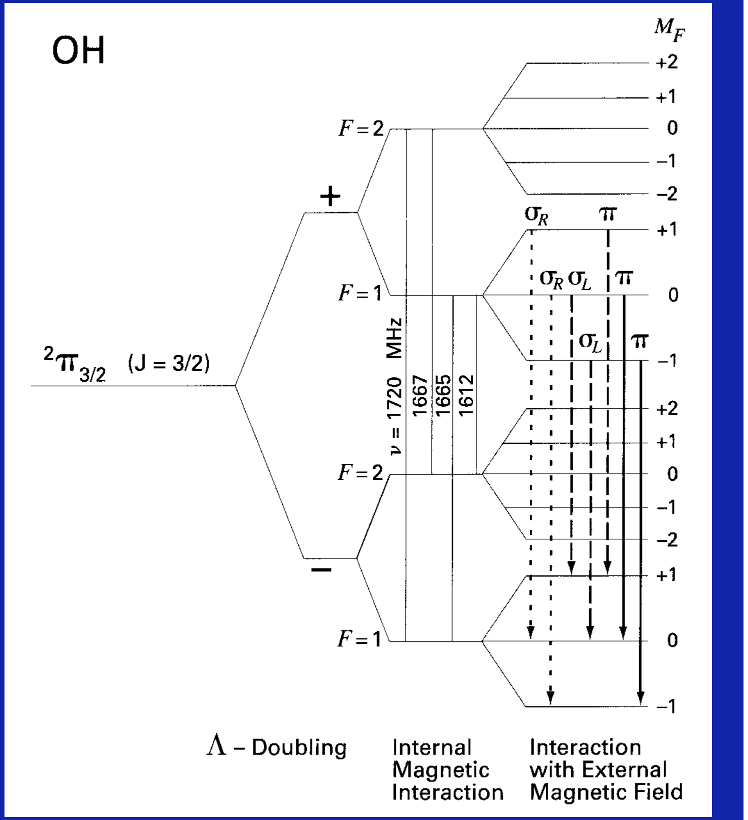
The Zeeman effect and magnetic fields I





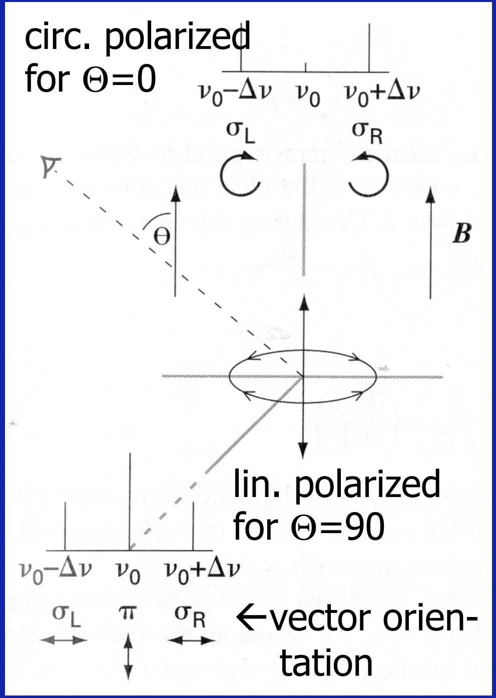
Λ doubling: Symmetry of orbital axis

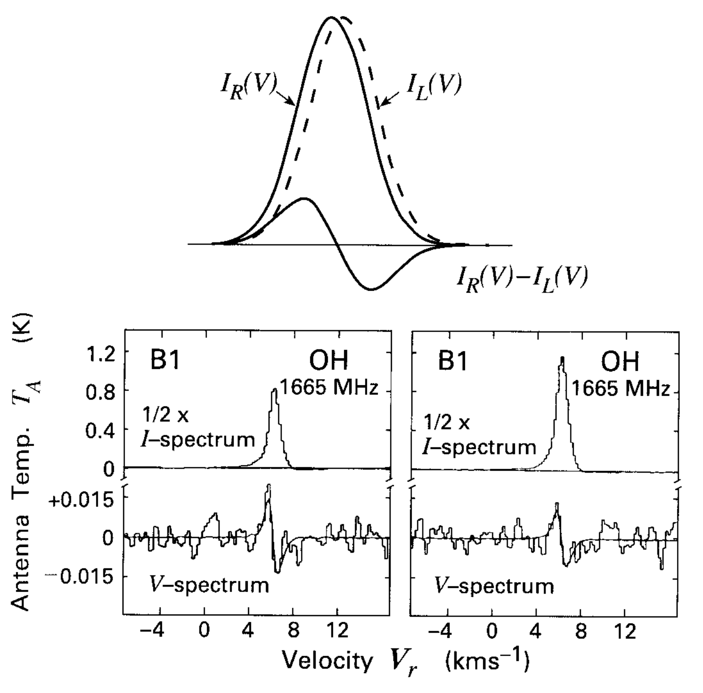


- OH free radical with unpaired free e$^-$ and hence non-zero electronic angular momentum. Highly reactive in lab, but can survive in space.
- Λ doubling → symmetry difference of the e$^-$ orbital to rotation axis
  → energy splitting.
- Interaction between spins of e$^-$ and H nucleus → magn. hyperfine splitting.

External magnetic field → produces Zeeman splitting.

The Zeeman effect and magnetic fields II





$\Delta\nu_{mag} \sim B$

usually of the order a few 10 μG

←Subtraction of left- from right-handed circ. pol. results in Stokes-V spec.

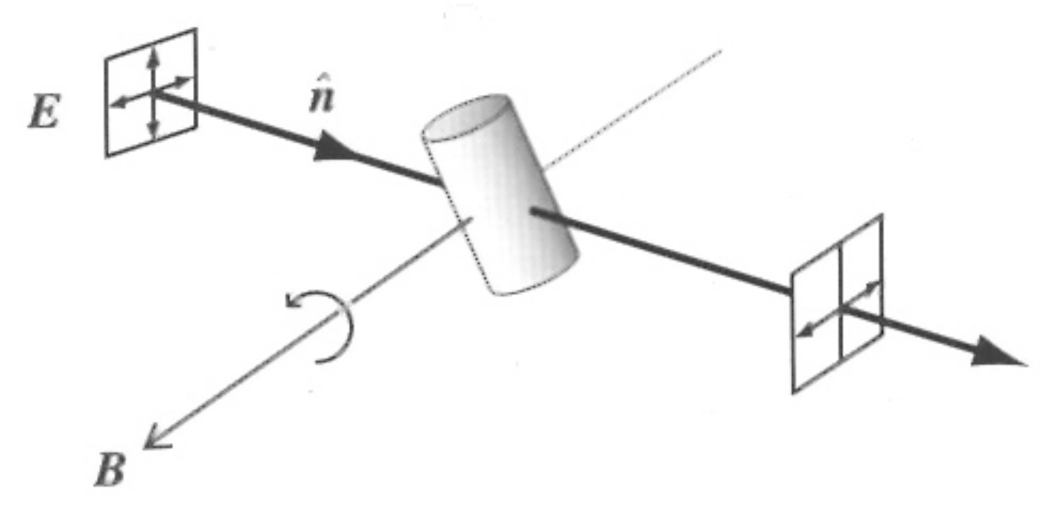
- Zeeman splitting of $\Delta\nu_{mag} = (b/2)\ B$ (b: constant, B: Magnetic field)

- Orientation between **B** and line of sight causes different polarization properties
  → Θ=0 (l.o.s) circular pol., Θ=90 lin. pol., in reality elliptical polarization.

- Thermal line broadening complicates matter: $\Delta\nu_{mag}/\Delta\nu_{therm} \sim 10^{-3} B(\mu G)$

  → One measures two polarizations differentially.
  → Only sensitive to the B component along the line of sight.

IR dust polarization & dichroic extinction



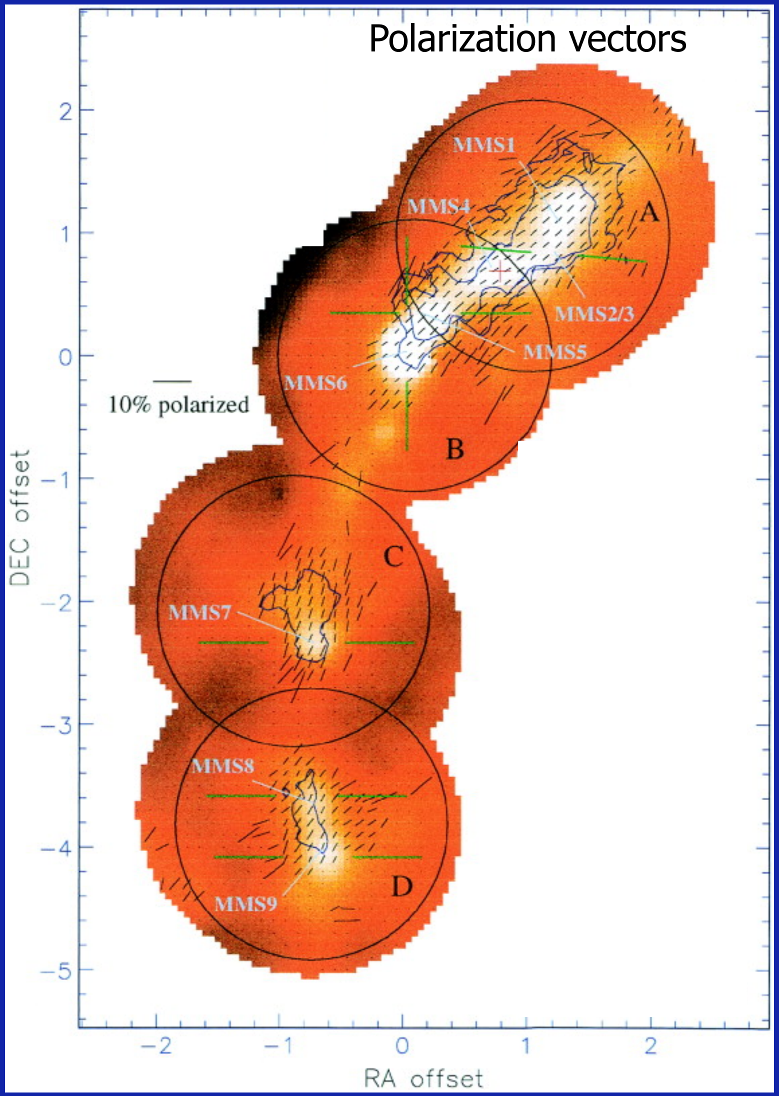
- Incident radiation unpolarized.

- Grains no spheres → small electric charge and paramagnetic
  → rotate about short axis, and magn. moment **M** points along rotation axis.

- External **B** field **M**x**B** forces grain's short axis to align with **B**.

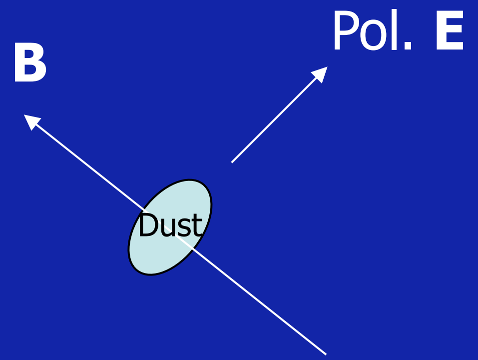
- Absorption best along major axis
  → Polarization afterwards largely along axis of magn. field.

Dust polarization and magnetic fields



Polarized submm continuum emission

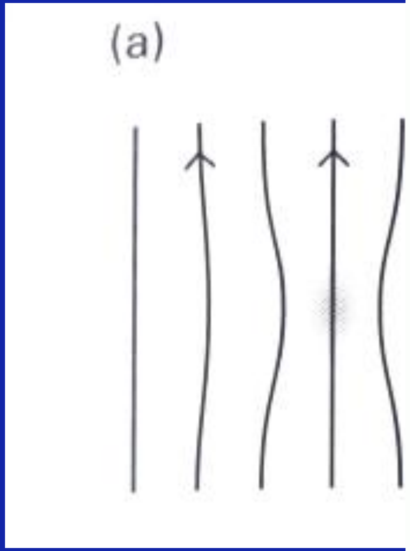
In contrast, thermal dust emission at (sub)mm wavelengths perpendicular to magnetic field!

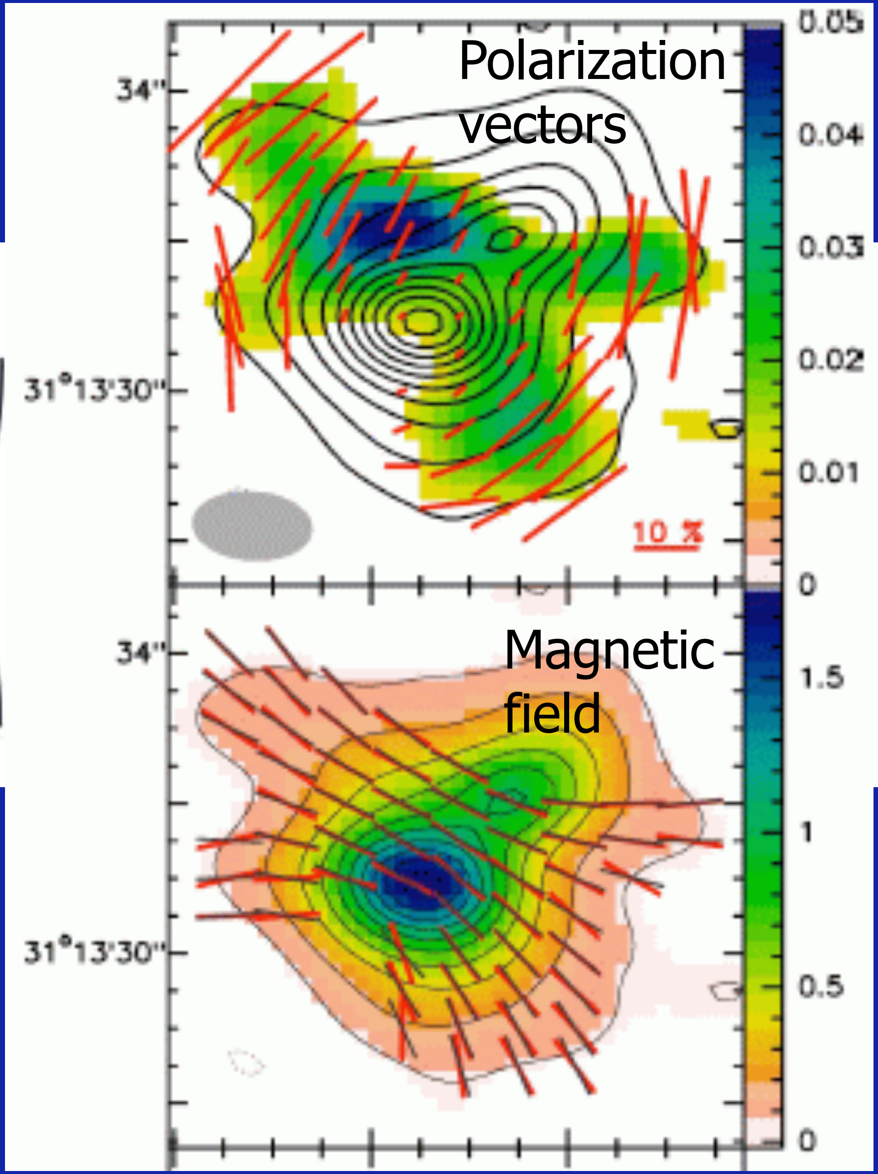


Molecular filaments can collapse along their magnetic field lines.

Matthews & Wilson 2000

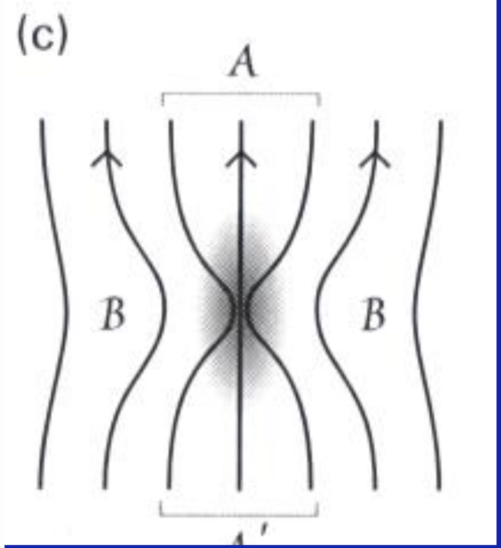
Ambipolar diffusion





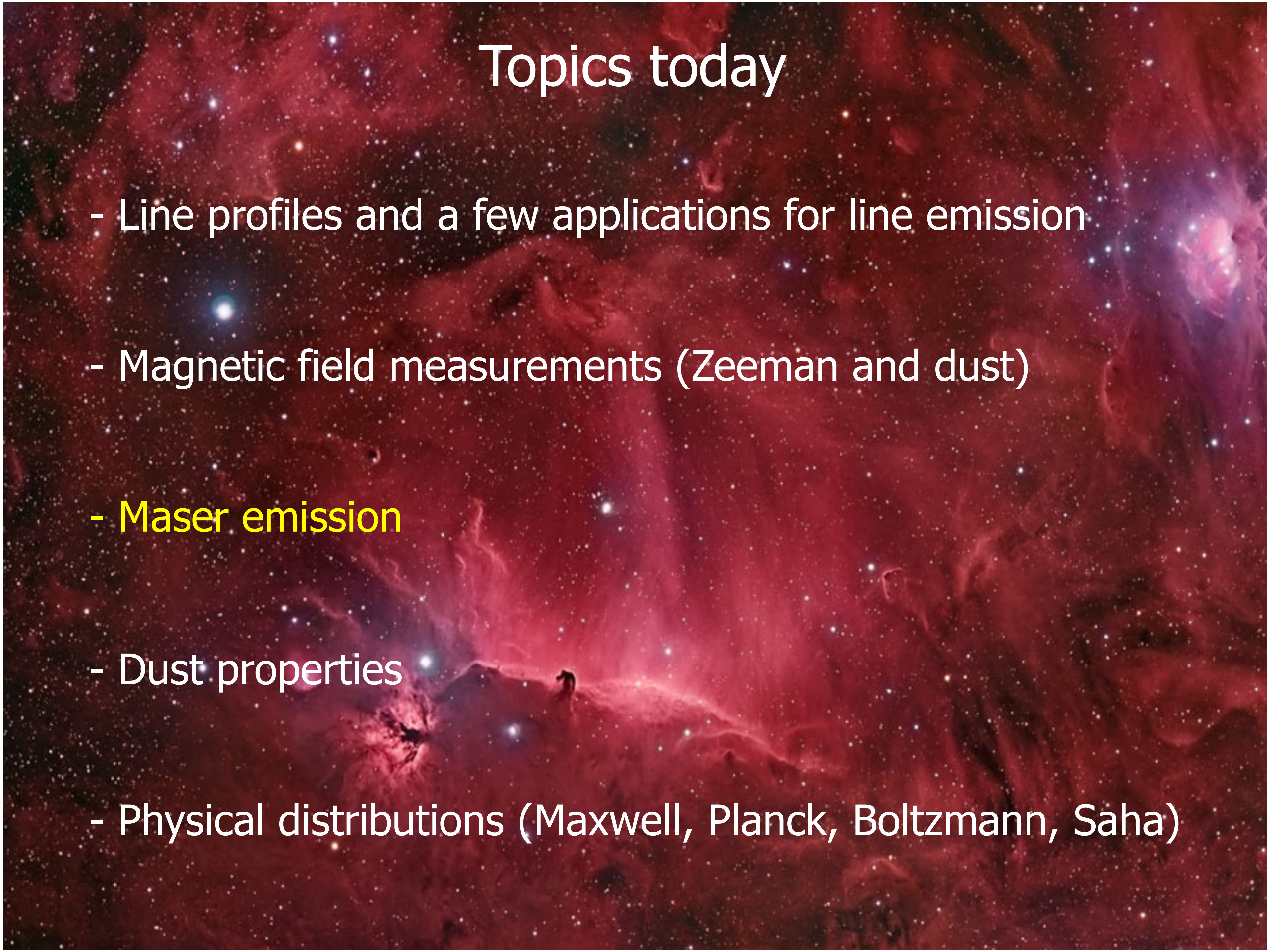
Polarization vectors

Magnetic field



B_{pos} ~ 5mG

Girart et al. 2006



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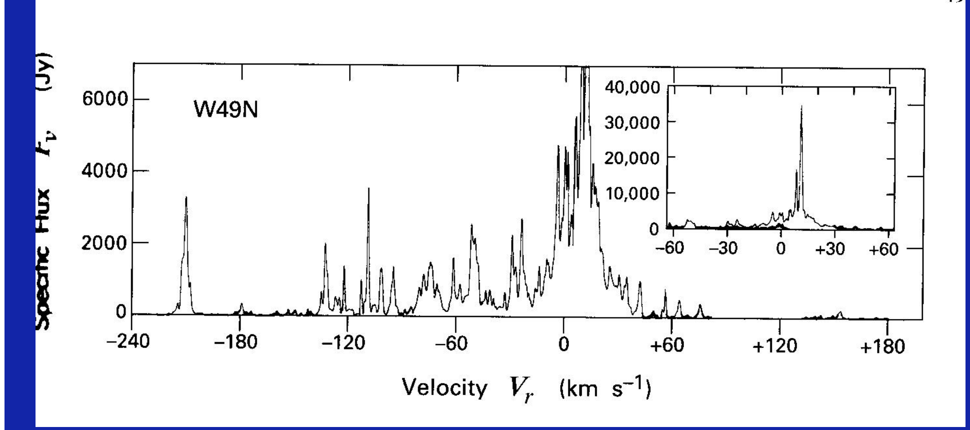
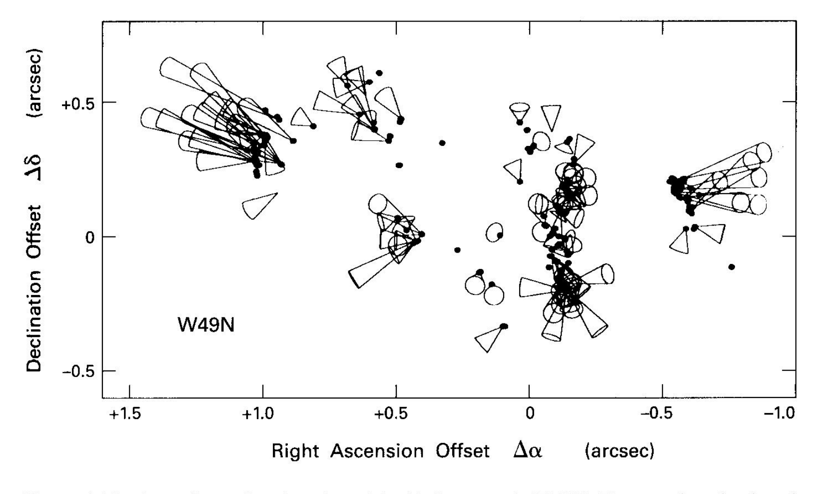
Molecular Masers I





- Rayleigh-Jeans limit: $T = c^2/2k\nu^2 I$ with $I=F/\Omega$ (Ω: solid angle).

- Small spot diameters (~ some AU) → T as high as 10^{15}K
  → no thermal equilibrium and no Boltzmann distribution.

- Narrow line-width with potential broad velocity distribution (many components).

- They allow to study proper motions.

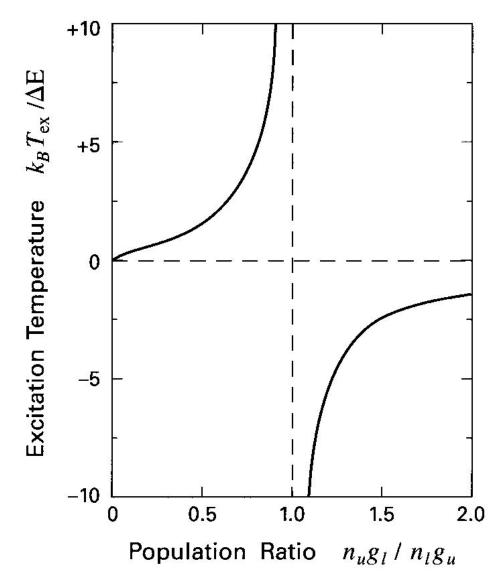
Molecular Masers II

- Excitation temperature for Boltzmann: $n_u/n_l = g_u/g_l \exp(-h\nu/kT_{ex})$.
- Maser activity requires population inversion: $n_u/g_u > n_l/g_l$.

  → Negative excitation temperatures
- In thermal conditions at a few 100K, for typical microwave lines

  $E_{line} = h\nu/k < T_{kin} \sim T_{ex} \rightarrow n_u/g_u \sim n_l/g_l$

  → Only a relatively small shift is required in get population inversion

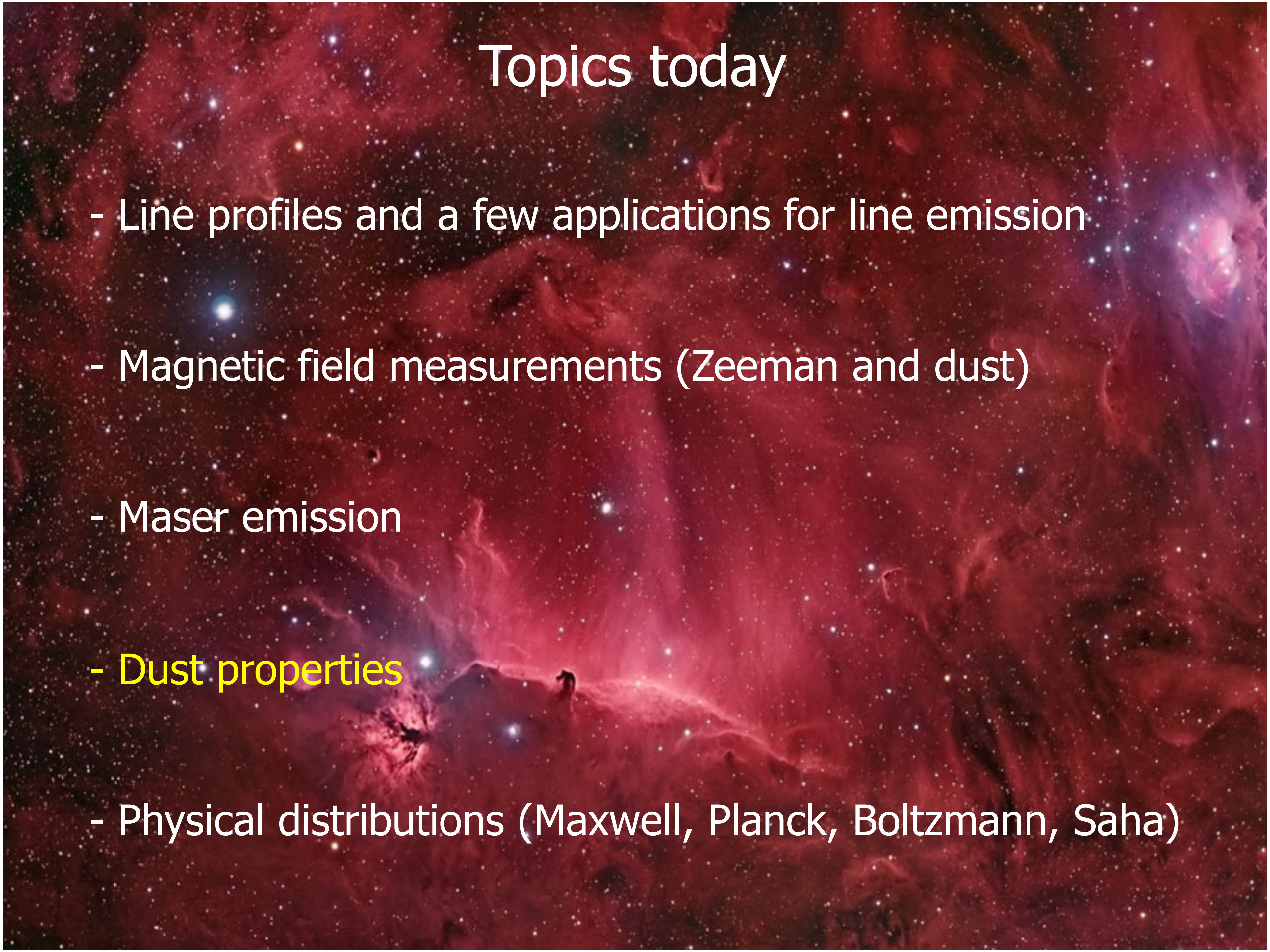


$T_{ex}/E_{line} = -1/\ln(n_u g_l/n_l g_u)$

Rising T_{ex}
→ Level populations approach each other
→ Only "overcome the border".

Pumping mechanisms, e.g.:
- Collisional pumping in shocks of protostellar jets for H_2O masers.
- Radiative pumping in shocks or from protostars (e.g., CH_3OH masers)

  → In both cases, very high densities and temperatures required.



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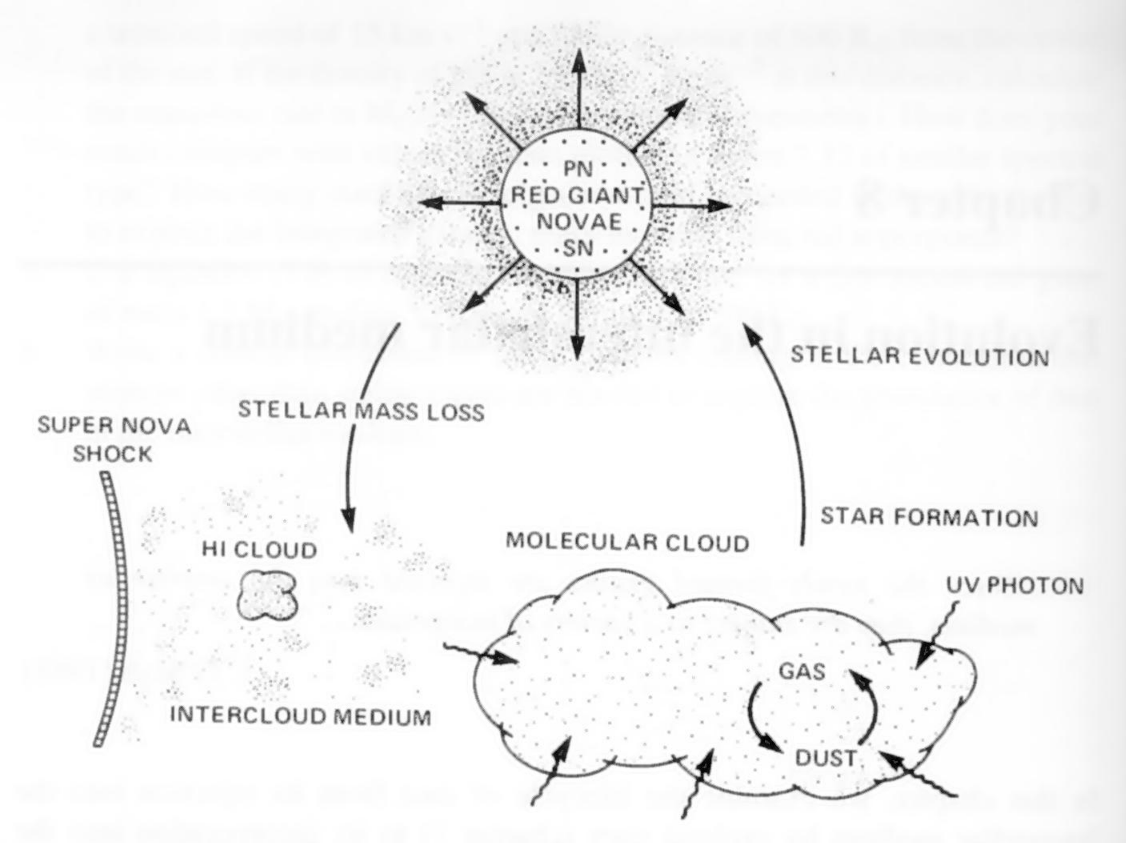
Dust within the big circle of matter



Dust composition:

Graphite	C
Silicon carbide	SiC
Enstatite	$(Fe,Mg)SiO_3$
Olivine	$(Fe,Mg)_2SiO_4$
Iron	Fe
Magnetite	Fe_3O_4

Size distribution:

Between 0.005 and 1μm

$n(a) \sim a^{-3.5}$ (a: size)

(Mathis, Rumpl, Nordsieck 1977)

Gas to dust mass ratio:

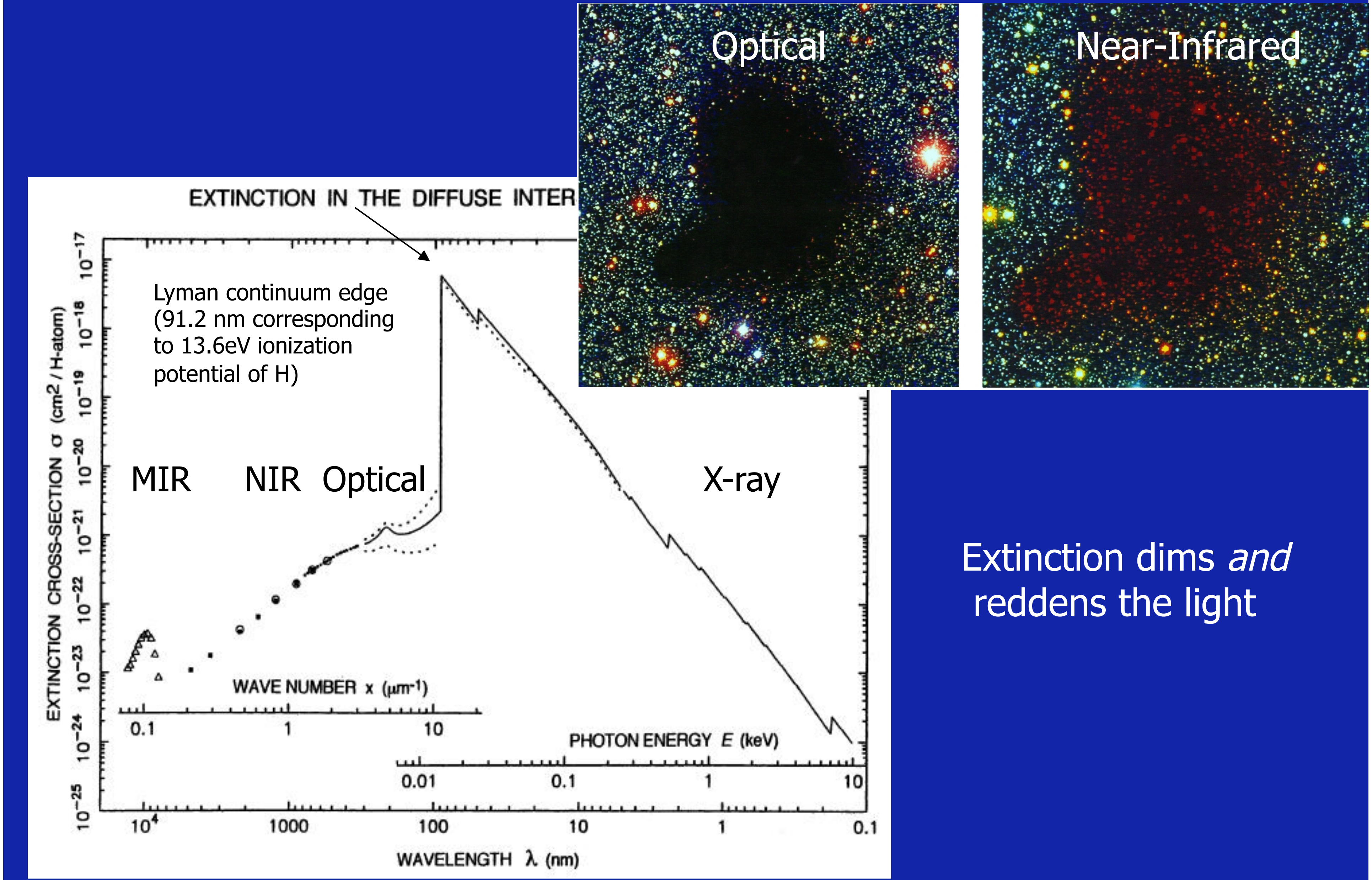
Canonical 1:100

Recent work suggests 1:150 (Draine et al. 2011)

Producers:

- Outer atmosph. of Red Giants, Planetary Nebulae (PN)
- Supernovae (SN)
- More recent: Dust can also form in the general interstellar medium (ISM).

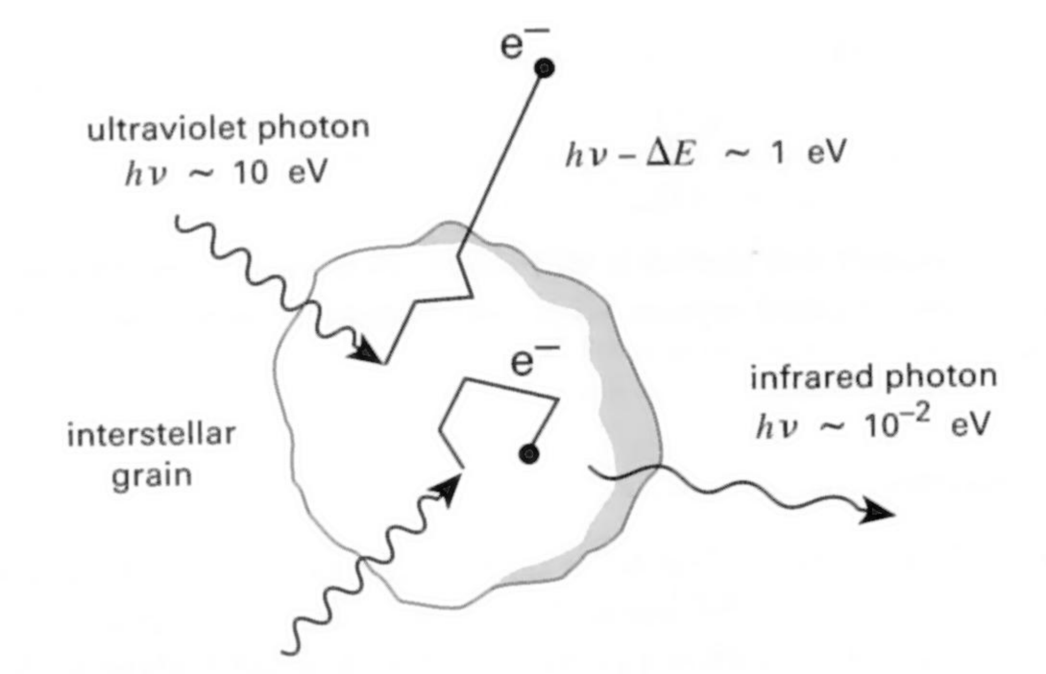
Interstellar dust: Extinction at shorter wavelengths



Lyman continuum edge (91.2 nm corresponding to 13.6eV ionization potential of H)

Extinction dims *and* reddens the light

Dust action at longer wavelengths: Re-emission

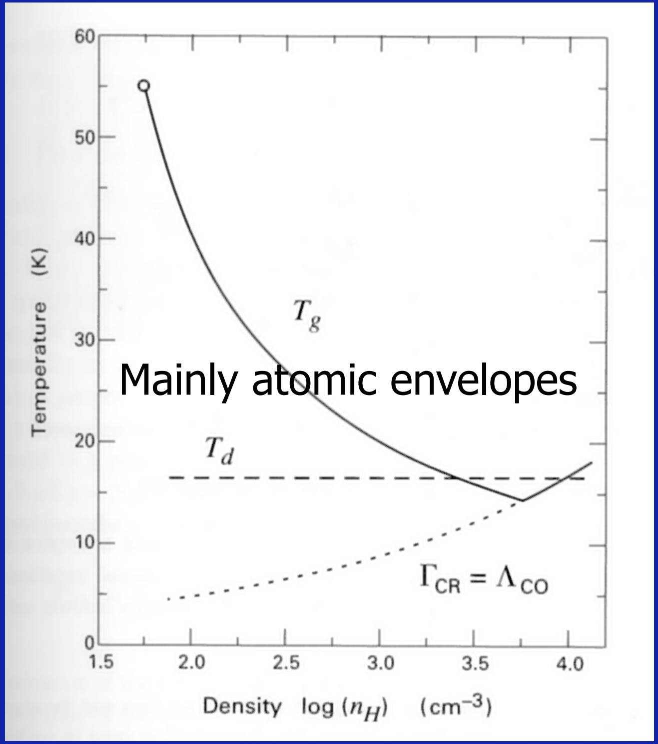


Dust grain hitted by UV photon:

1) Photoelectrical effect → give energy to e^- → leaves grain and heats gas.

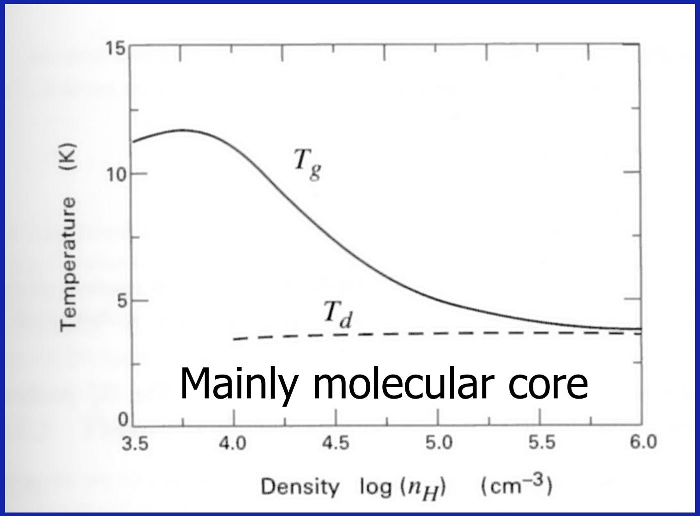
2) Excites lattice vibrations → transformed to (far)-IR photons and re-emitted.

Dust and gas coupling



Mainly atomic envelopes

Cosmic ray heating more sensitive to density than CO cooling. Hence T_g rises again.



Mainly molecular core

Γ – heating rate Λ – Cooling rate

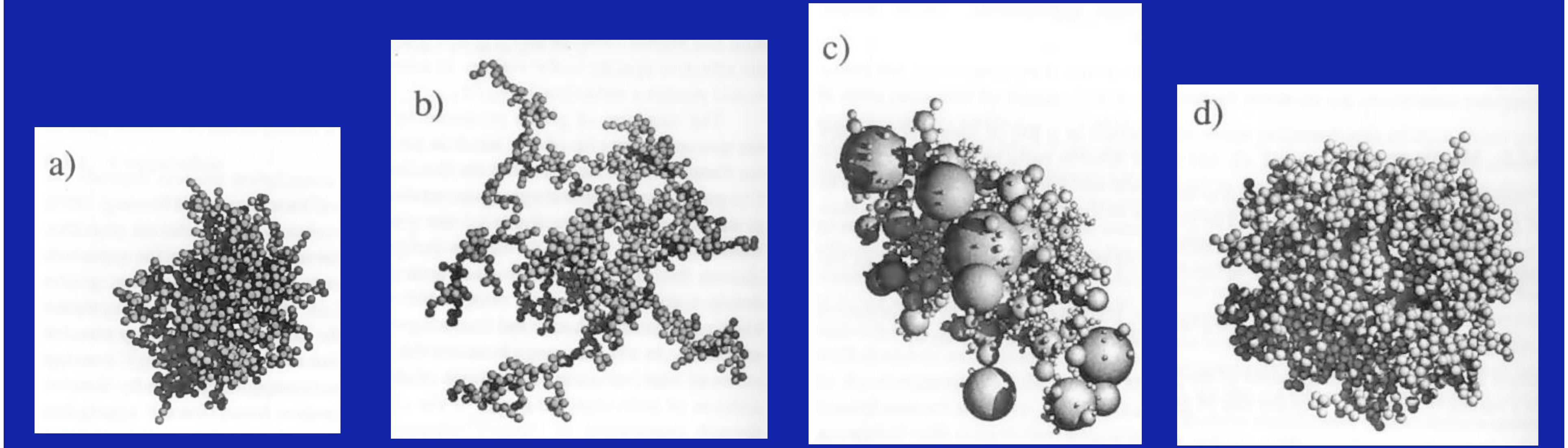
- Low densities: gas and dust de-coupled; at high densities coupled.

- Low densities gas cooling mainly CO; high densities via CO & dust.

- At very high densities gas and dust temperatures approach each other

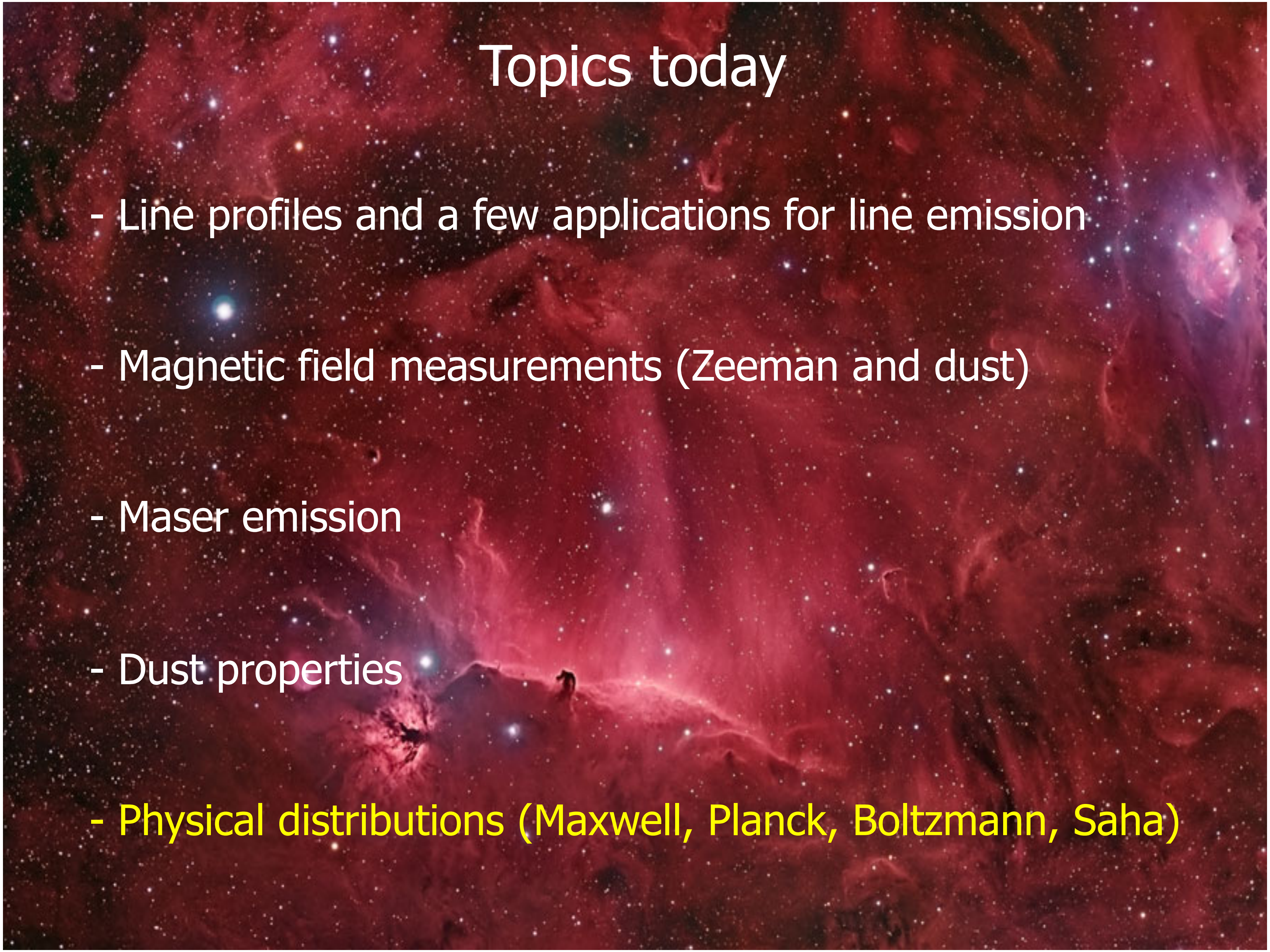
→ CO cooling becomes insignificant then!

Dust incarnations



Dust can grow and coagulate in very dense environments, e.g., disks.

Figures: Simulations of dust grain cluster growth for different initial parameters (gas and dust density, temperature, stickyness, grain charge, coagulation time …).
(From Dorschner & Henning 1995)



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Physical conditions : Micro-Level

A medium in thermodynamic eq. can be described by 4 distribution laws:

1.) MAXWELL distribution of the particle velocity contributions (kinetic energy):

$$N(v;T) = 4\pi\left(\frac{m}{2kT}\right)^{3/2} v^2 \exp\left(-\frac{mv^2}{2kT}\right)$$

v : particle velocities

2.) BOLTZMANN distribution of the population numbers of the particle energy levels:

$$\frac{N_o}{N_u} = \frac{g_o}{g_u}\exp\left(-\frac{E_o - E_u}{kT}\right)$$

$E_{o/u}$ → Energies of the upper (o) and lower (u) levels

$g_{o/u}$ → Corresponding statistical weights

3.) PLANCK radiation law (distribution of the photon energies):

$$B_\nu = \frac{2h\nu^3 / c^2}{\exp(h\nu / kT) - 1}$$

ν : photon frequencies

4.) SAHA equation (distribution of the ionisation levels in plasma):

$$\frac{N_{j+1} N_e}{N_j} = \frac{2\,U_{j+1}(T)}{U_j(T)}\left(\frac{2\pi m_e kT}{h^2}\right)^{3/2} \exp(-\chi_{j,j+1} / kT)$$

N_{j+1}, N_j - Number densities of (j+1)-fold and j-fold ionised particles

N_e - electron density

$\chi_{j,j+1}$ - ionisation energy needed to get from ionisation level j to $j+1$

U_{j+1}, U_j - partition function for both states

Physical conditions : Micro-Level

Are these distribution functions valid in the ISM?

General rule: time scale for processes leading to equilibrium short compared to time scales of disturbing processes

1. Example : Collisions between H-atoms:

   Consider:
   $T = 100$ K → mean v ~ 1 km/s; cross section $\sigma = \pi R_H^2 \sim \pi (0.1 \text{ nm})^2$

   → average time between two collision $\tau_s = (v\, \sigma\, n_H)^{-1}$
   → with HI density of 1 cm^{-3} → $\tau_s \sim 1000$ yrs

   → short compared to most interstellar processes (except shock fronts)

→ Maxwell distribution valid, introduction of kinetic temp. T_{kin} reasonable!

Physical conditions : Micro-Level

2. Example: Balance for energy level population numbers for ISM:

Correction factor to Boltzmann: $$\frac{1}{1 + (A_{21} / (n\,Q_{21}))}$$

(Pure Boltzmann only if $(n\,Q_{21}) >> A_{21}$)

A_{21} [s^{-1}] Einstein coefficient for spontaneous radiative decay
Q_{21} [$m^3\,s^{-1}$] collision rate
n [m^{-3}] number density

- In thin ISM collision rate small (Example 1).

- For dense cores: E.g. CO(1-0) at density $10^5 cm^{-3}$: $A_{21}=7.2x10^{-8}s^{-1}$, $Q_{21}=3.3x10^{-11}\ cm^3s^{-1}$

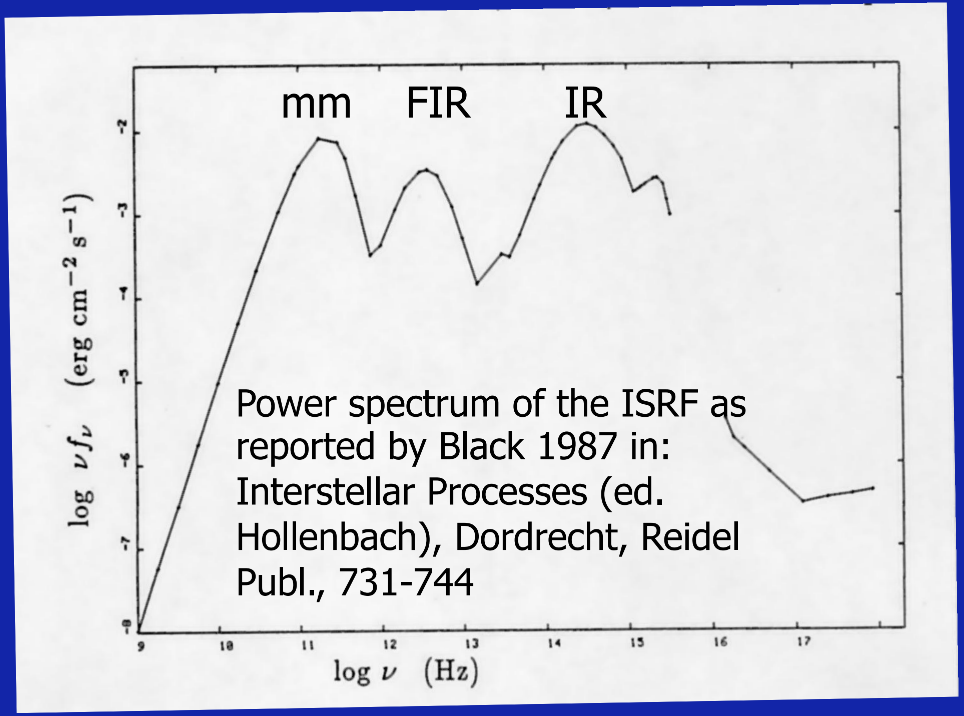
→ $A_{21} / (n\,Q_{21}) \sim 0.02$
→ Boltzmann distribution valid in dense cores!

Physical conditions : Micro-Level

3. Example : Interstellar radiation field (ISRF) :

Sum of emission contributions from all emitting objects (stars, dust, gas) in the nearer and further vicinity of the gas cloud

ISRF cannot be approximated by a black body (i.e., Planck function not applicable)
ISRF hence far from thermodynamic equilibrium …



Power spectrum of the ISRF as reported by Black 1987 in: Interstellar Processes (ed. Hollenbach), Dordrecht, Reidel Publ., 731-744

However: Dense cores and stars can be fitted relatively well with single or multiple black body functions.

Summary

- Line profiles (thermal and kinematic broadening) and some applications

- Magnetic fields are very important but difficult to measure:
  - Zeeman effect traces **B** component along line of sight.
  - Dust polarication traces **B** in plane of the sky.
    (Other magnetic field measurements possible.)

- Masers are non-thermal processes. Good for high spatial accuracy and proper motion studies.

- Dust important from many points of view:
  - Traces warm and cold components of ISM.
  - Important coolant at high densities.
  - Traces magnetic field.
  - Chemical catalyst.

- Physical distributions and their applicability to the ISM.

Sternentstehung - Star Formation

Winter term 2017/2018
Henrik Beuther & Thomas Henning

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More Information and the current lecture files: http://www.mpia.de/homes/beuther/lecture_ws1718.html
beuther@mpia.de, henning@mpia.de